Survey of Present and Future Ground-Based Imaging Systems

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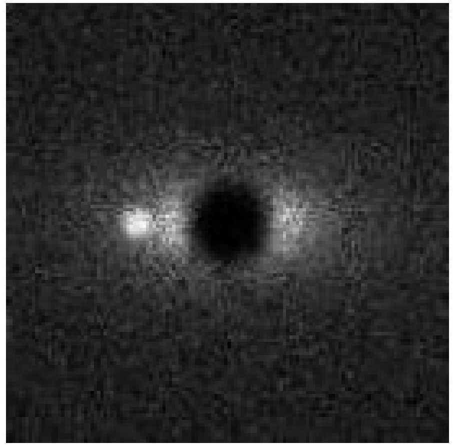
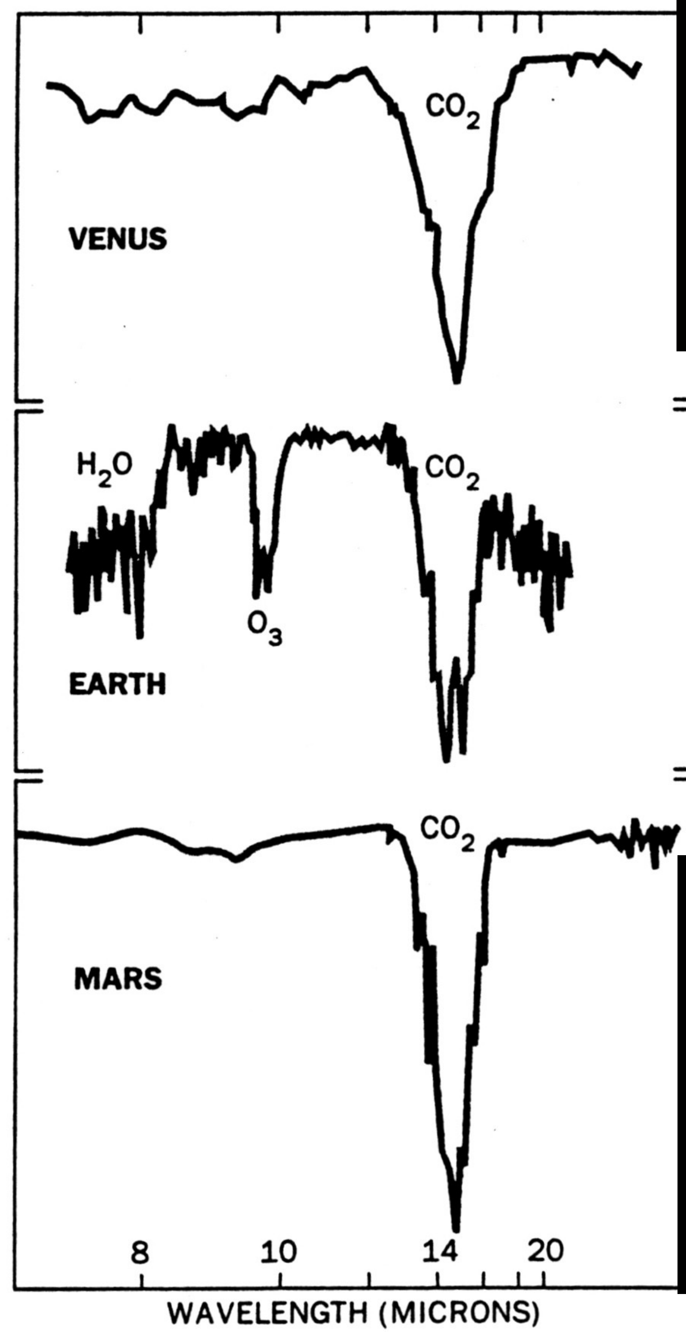
Subaru Telescope

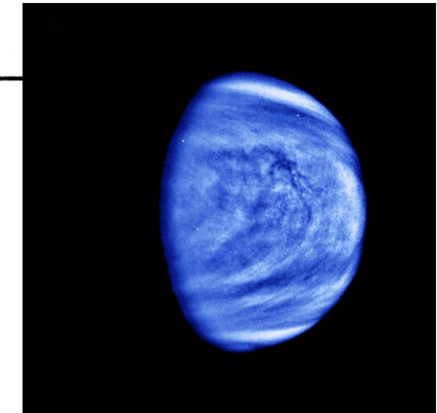
With material from :

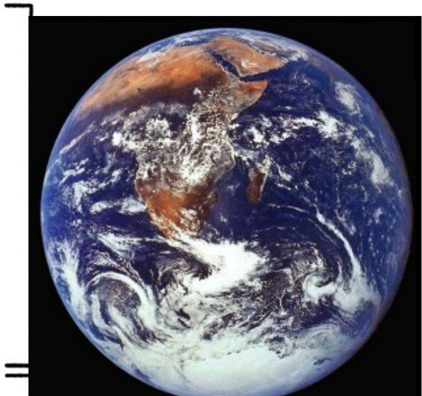
Rich Dekany, Ben Oppenheimer (Palm 3000, P1640)
Bruce Macintosh (Gemini Planet Imager)

Direct imaging of Exoplanets allows extensive characterization









Orbit

Atmosphere composition

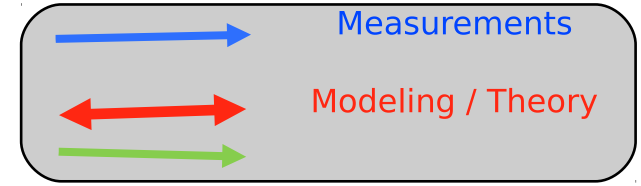
Continents vs. Oceans ?

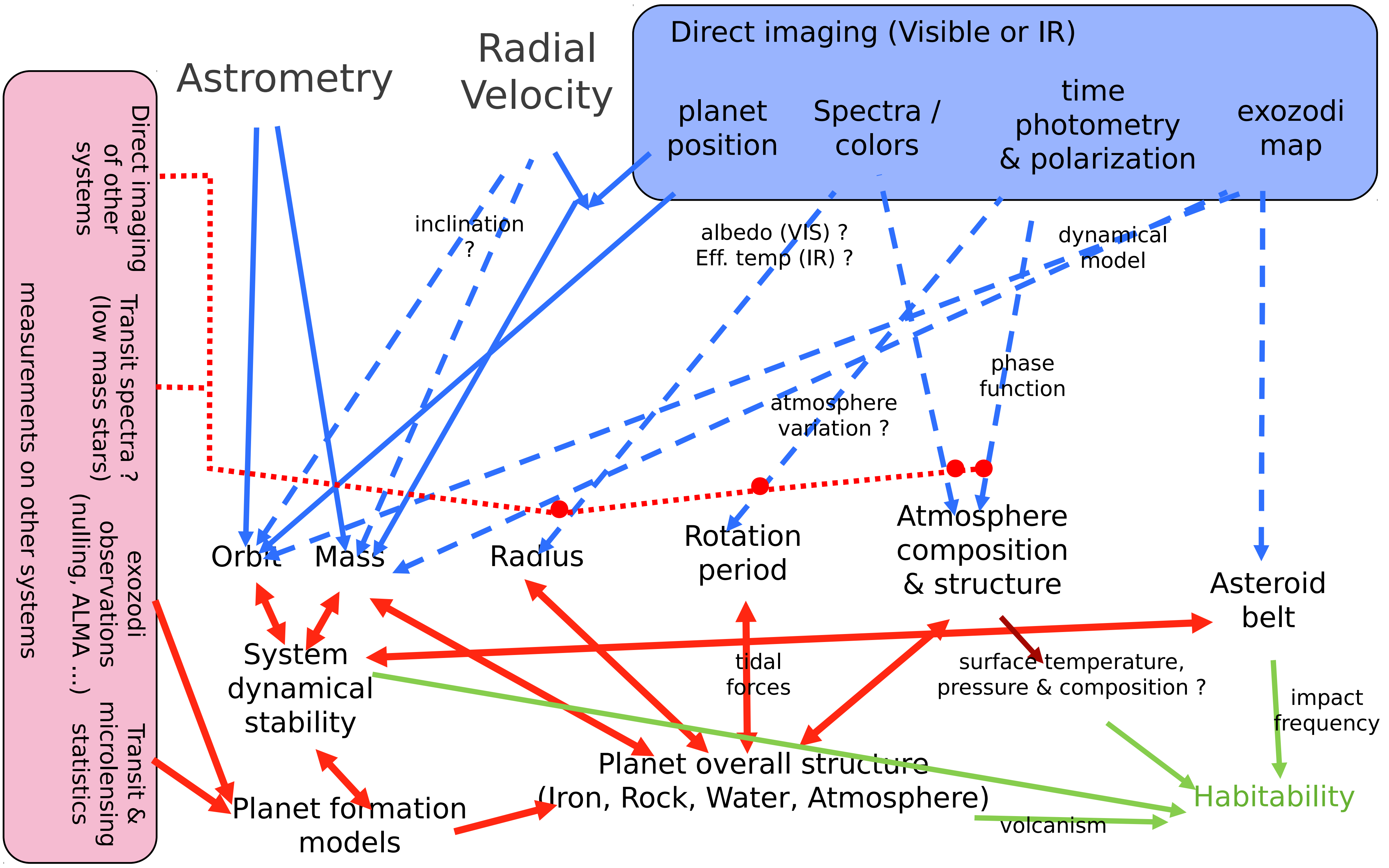
Rotation period

Weather patterns

Planetary environment :
Planets + dust

Exoplanet characterization with direct imaging

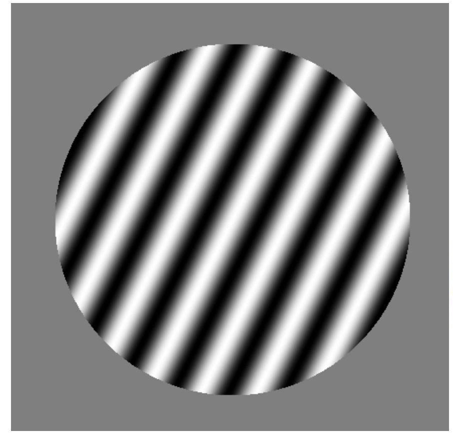




Why is it difficult ?
(even in space)

pupil plane complex amplitude

$$W(\vec{u}) = A(\vec{u})\, e^{i\phi(\vec{u})}$$



Cosine aberration in pupil phase

$$\phi(\vec{u}) = \frac{2\pi h}{\lambda} \cos\left(2\pi \vec{f}\vec{u} + \theta\right)$$

… creates 2 speckles

$$I(\vec{\alpha}) = PSF(\vec{\alpha}) + \left(\frac{\pi h}{\lambda}\right)^2 [PSF(\vec{\alpha}+\vec{f}\lambda) + PSF(\vec{\alpha}-\vec{f}\lambda)]$$

High Contrast

Small angular separation

Low Flux

What would it take to image an Earth-like planet around Sun-like star at 10pc ?

~1e-10 contrast in visible light, h=1.6e-12 m (0.0012 nm) = 1e-10 speckle

1e-10 speckle (or 1e-10 contrast planet) around Sun at 10pc = 0.1 ph/sec/m^2/um

On a 4-m telescope, with 10% efficiency and a 0.5 um spectral band:

Earth = 0.6 ph/sec

To measure phase and amplitude of speckle requires ~10 photon

10 photon = 16 sec

→ **This spatial frequency would need to be stable to 1/1000 nm over ~ minute**

(in near-IR, 1e-7 contrast → h = 0.16 nm)

Exoplanets: Contrast ratio, visible vs. infrared

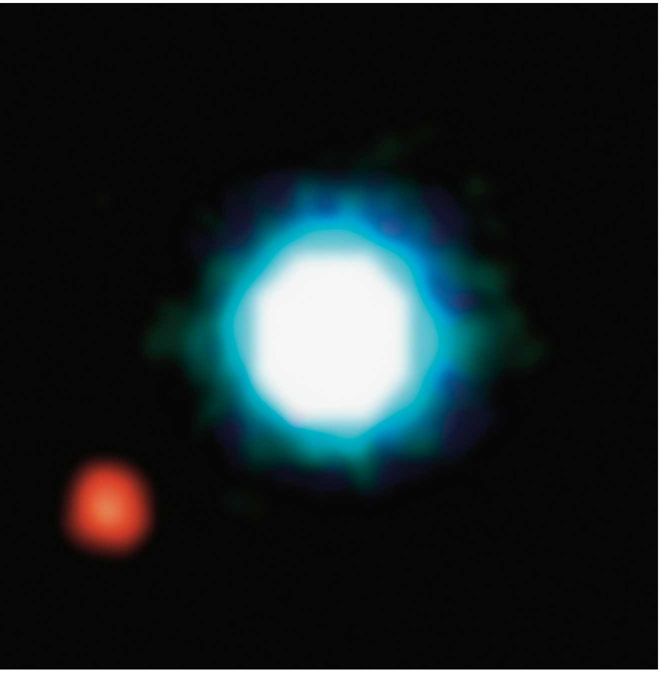
In the visible, the contrast is very challenging unless the
Exoplanet is very close to the star (luminosity goes as d^{-2})
Required wavefront quality cannot be achieved from the ground
Habitable planets cannot be imaged from the ground,
and will likely require dedicated space telescope+instrument.

In the near-IR (1-5 um), giant and young planets
•("young Jupiters") can be imaged:
- Adaptive Optics systems work well in the near-IR
- Giant planets emit their own light (thermal emission)
- Young planets are still very hot, and cool slowly after formation

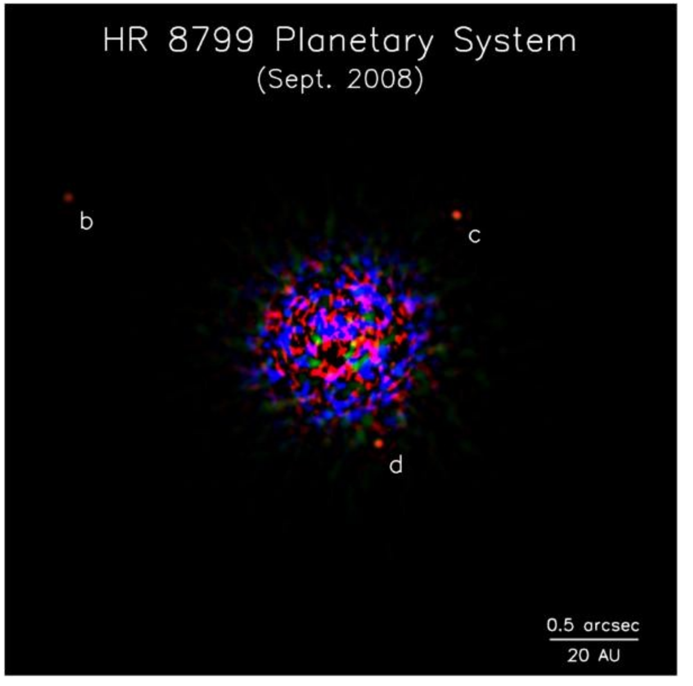
Best strategy for
Ground-based
imaging

In the Thermal IR (~10 um & longer), contrast
is even more favorable, and older giant planets
can be imaged (this is one of the key science goals of JWST)

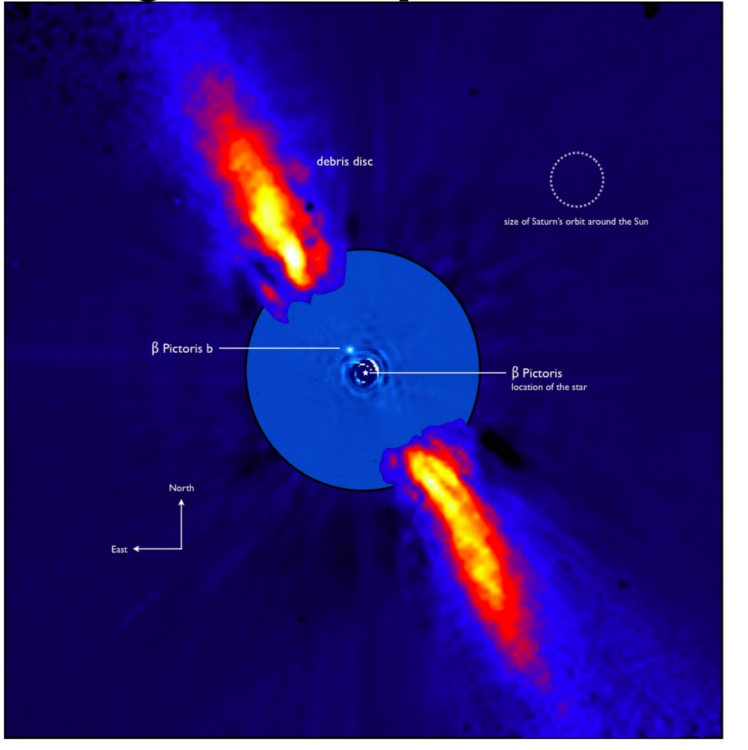
Exoplanets imaged with ground-based telescopes
… tip of the iceberg (Young, massive, and large orbits)



2M1207 exoplanet (Chauvin et al., ESO, 2004)
Possibly the first direct image of an exoplanet



HR8799: first image of exoplanetary sytem with multiple planets (Marois et al. 2009)



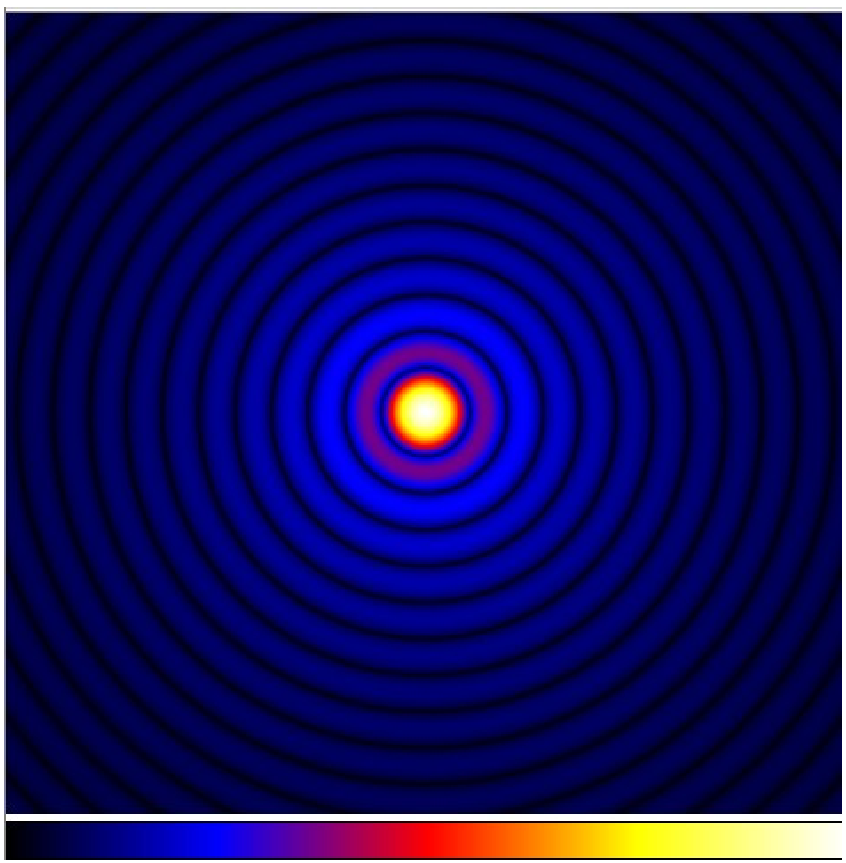
Beta Pic exoplanet and dust disk (Lagrange et al. 2009)

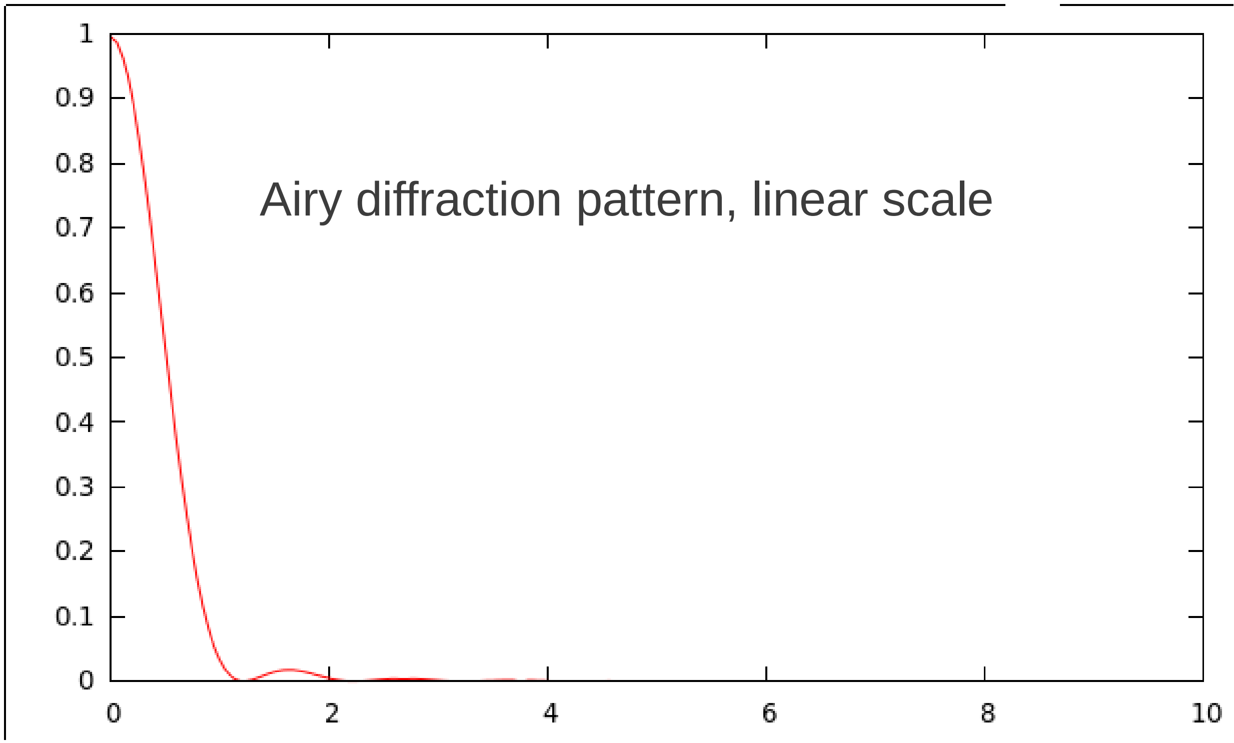
… using AO systems and camera which were not specifically designed for high contrast imaging (but often using optimized data acquisition and reduction techniques)

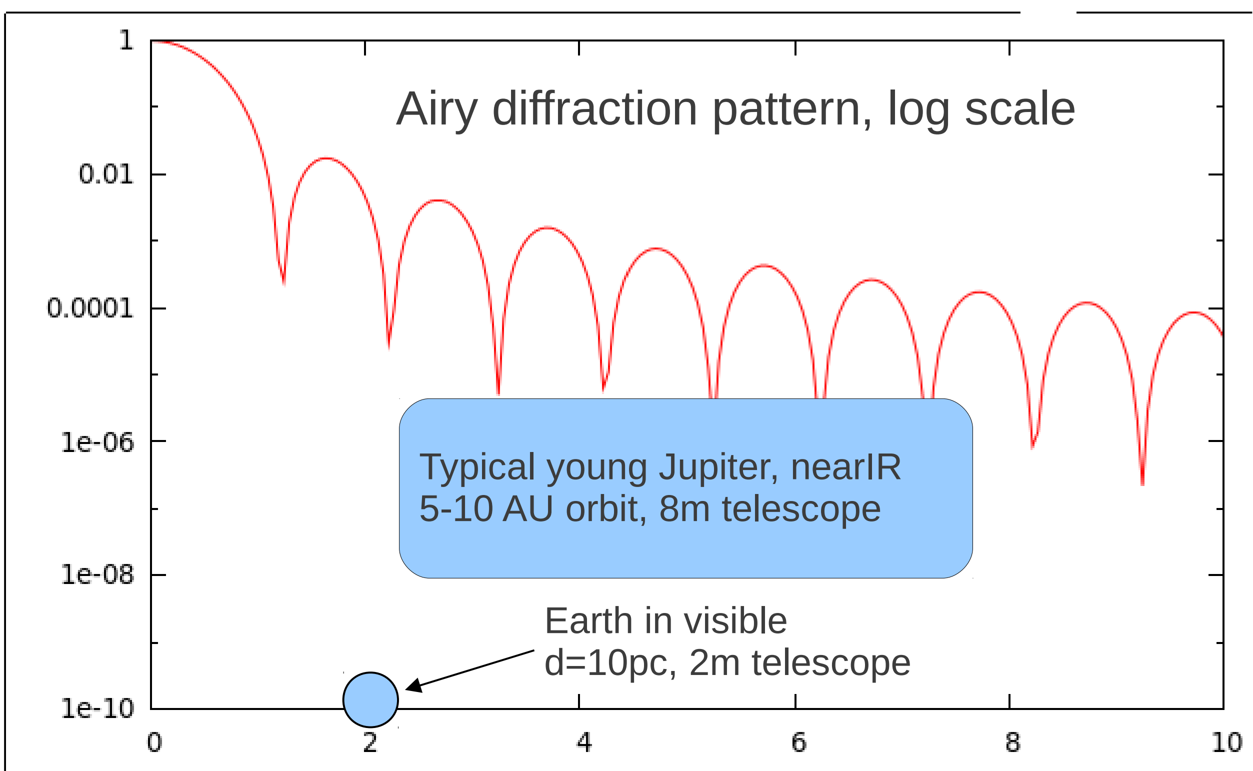
No coronagraph was used (in fact, coronagraphs are mostly useless on current AO systems)

Why coronagraphy ?

Conventional imaging systems are not suitable for high contrast (even if perfect) due to diffraction





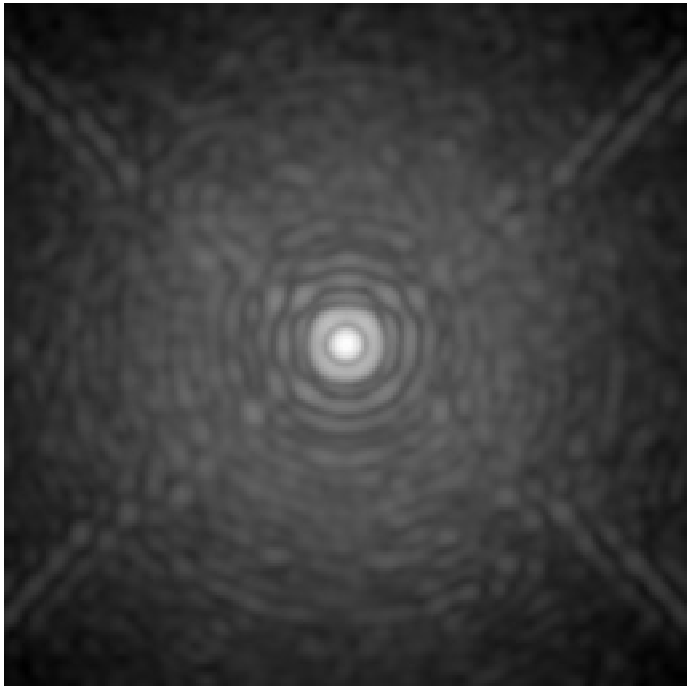


Why are coronagraphs (mostly) useless ?

A coronagraph can only remove a <u>known & static</u> diffraction pattern
Coronagraphs can't remove unknown (when the coronagraph is designed) diffraction (including speckles due to WF errors)
BUT static & known diffraction can be removed in the computer ?

Two fundamental reasons why a coronagraph is useful:
(1) Reduce photon noise from diffraction
(2) Avoid coherent amplification between speckles and diffraction pattern

Coronagraphs serve no purpose if the dynamic speckle halo (due to residual wavefront errors) is higher than the static diffraction (due to pupil shape)
… which is the regime under which most current AO systems operate.



Coherent amplification between speckles and diffraction pattern

Final image = PSF diffraction (Airy) + speckle halo

This equation is true in complex amplitude, not in intensity.

Intensity image will have product term → speckles are amplified by the PSF diffraction.

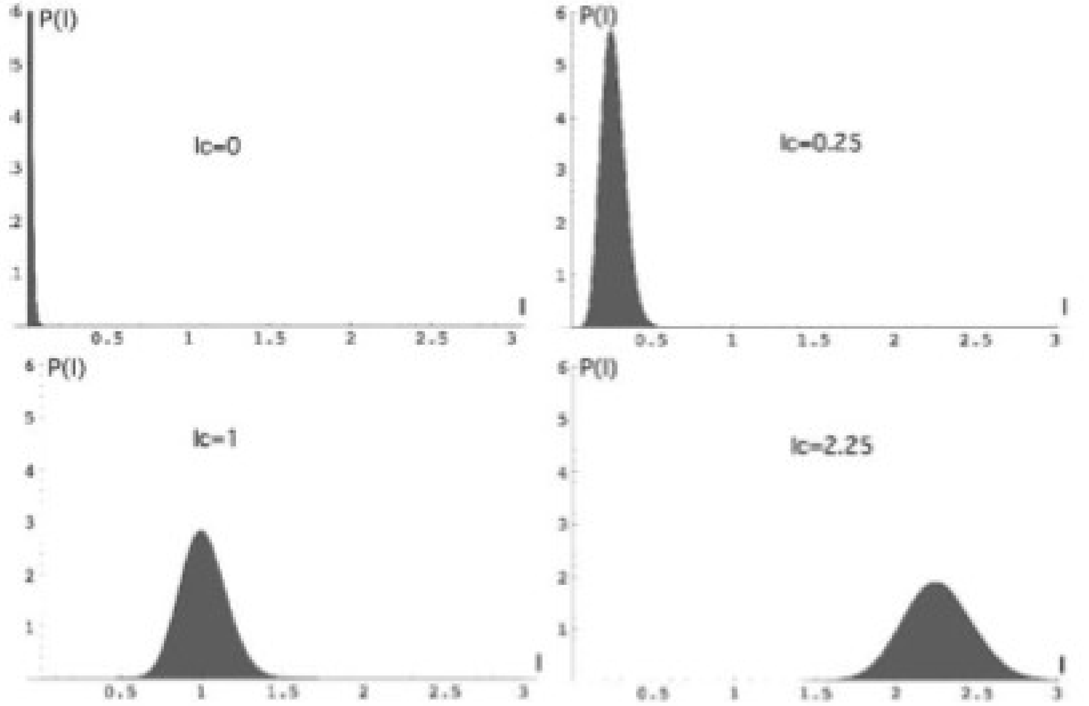
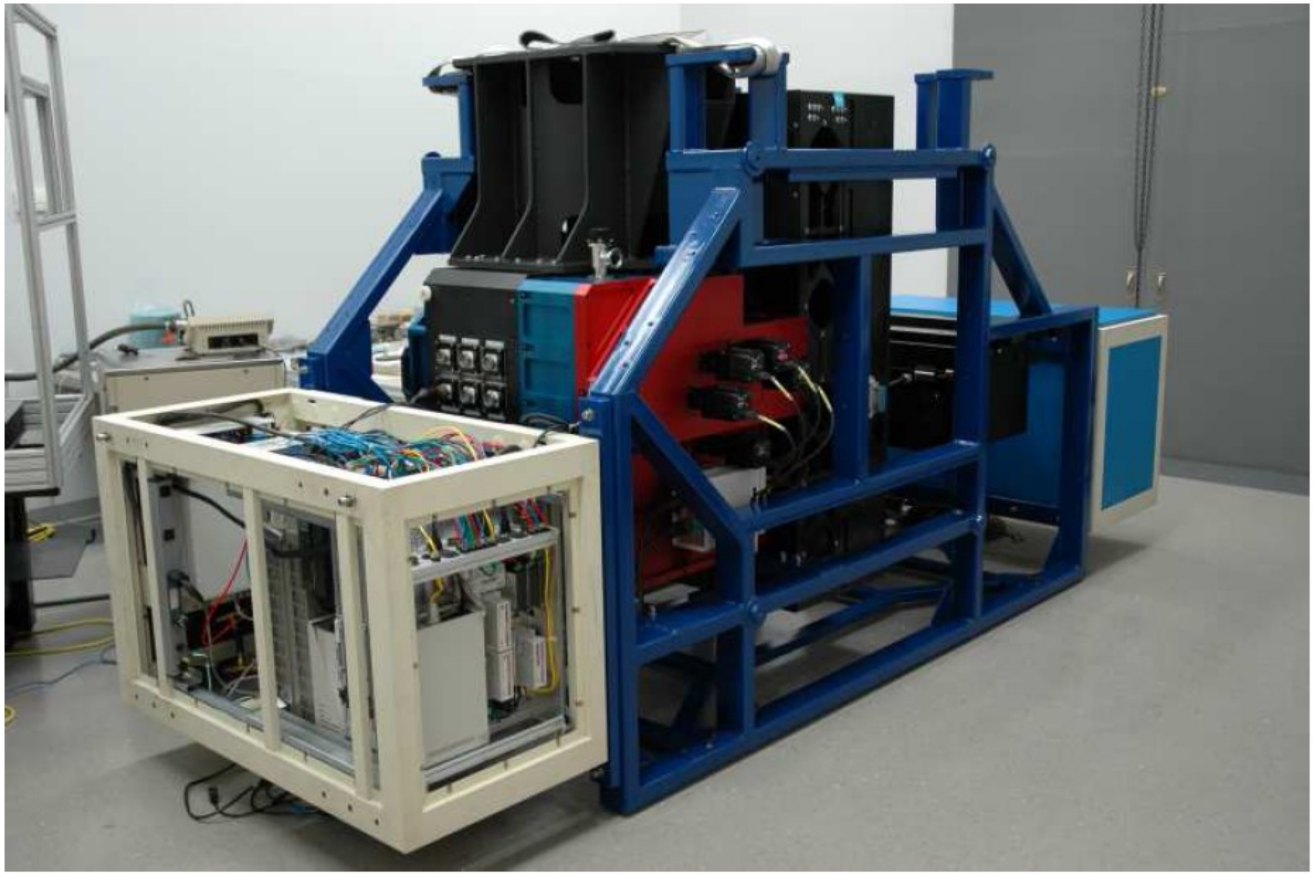
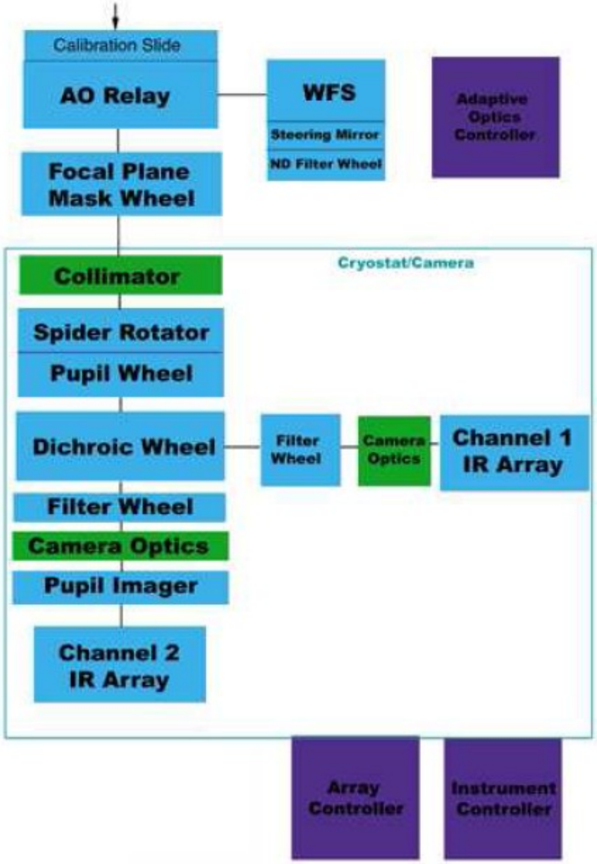


FIG. 3.—PDF of the light intensity at four different constant background intensity levels I_c and a single value of $I_s = 0.1$. High values of I_c correspond to locations near the perfect PSF maxima (rings), and low values of I_c correspond to locations near the zeros of the perfect PSF or far from the core. For $I_c = 0$ we have the pure speckle exponential statistics. The width of the distribution increases with an increase in the level of I_c. This explains speckle pinning; speckle fluctuations are amplified by the coherent addition of the perfect part of the wave.

Aime & Soummer 2004

NICI system (Gemini South)





Liu et al., 2010

85-element curvature AO system
Lyot type coronagraph
Simultaneous dual band imaging (inside / outside methane abs. Band)

Currently surveying a large sample of nearby young stars

NICI system (Gemini South)

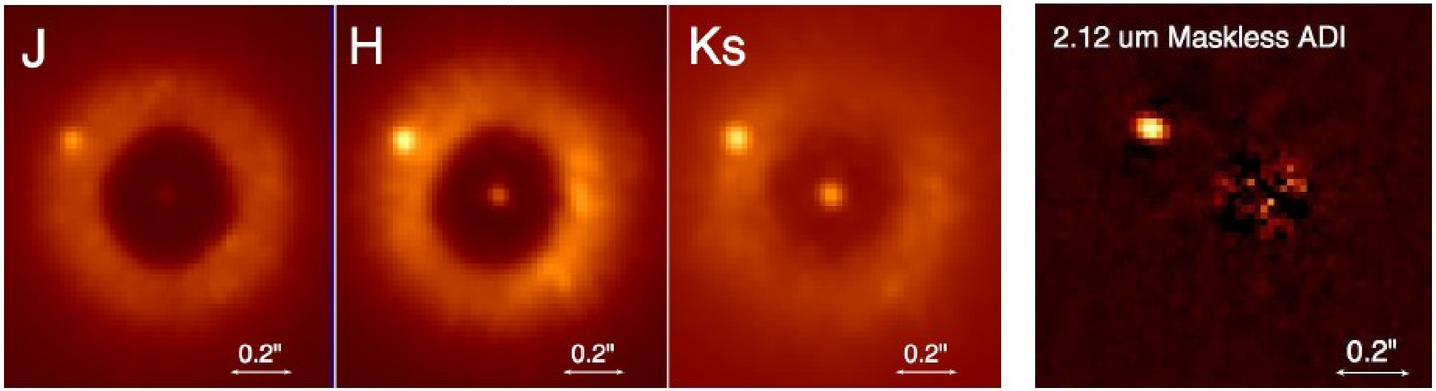


Fig. 1.— Left: J, H, and K_s-band images of the PZ Tel system obtained with NICI in direct imaging mode at the Gemini-South Telescope in May 2010. North is up, east is left. The primary resides at the center of the translucent 0.22" radius focal plane mask and is attenuated by a factor of 329 in J, 214 in H, and 131 in K_s-band. The confirmed companion is at 0.36" separation and PA=59.4° with flux ratios of $\Delta J = 5.40 \pm 0.13$, $\Delta H = 5.38 \pm 0.09$, and $\Delta K_s = 5.04 \pm 0.10$ mag. Right: Maskless ADI H_2 2-1 (2.12 μm) image from May 2010. Light from the primary has been removed by ADI processing.

Biller et al. 2010

NICI is the first system (coronagraph + AO) designed for a high contrast imaging survey of nearby young stars

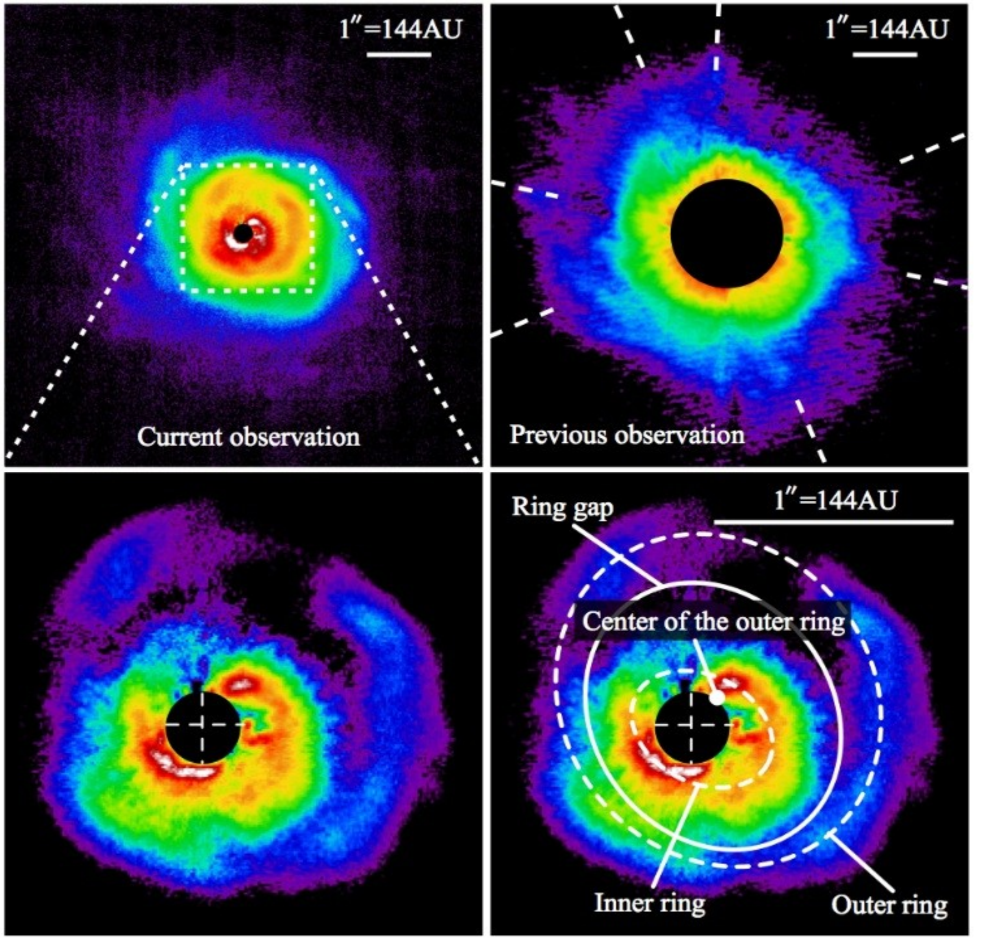
Combines together well-proven technologies:

- Curvature AO, 85 elements
- Lyot coronagraphy
- Spectral differential imaging

HiCIAO system (Subaru Telescope)

(see presentation by M. Tamura in this conference)

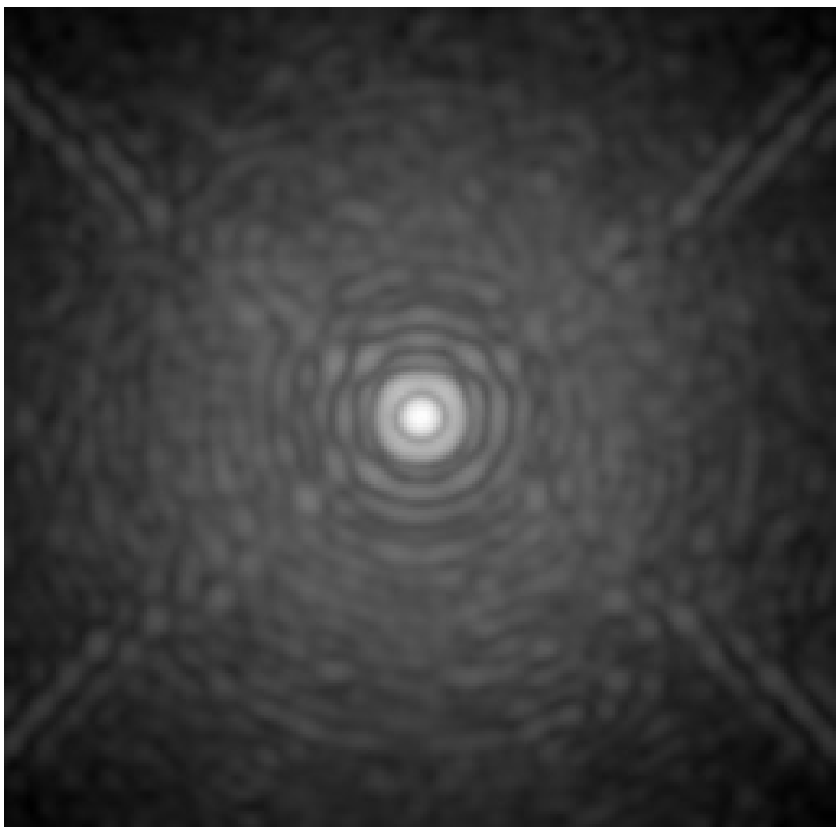
HiCIAO combines:
- Curvature AO, 188 elements
- Lyot coronagraphy
- Spectral differential imaging



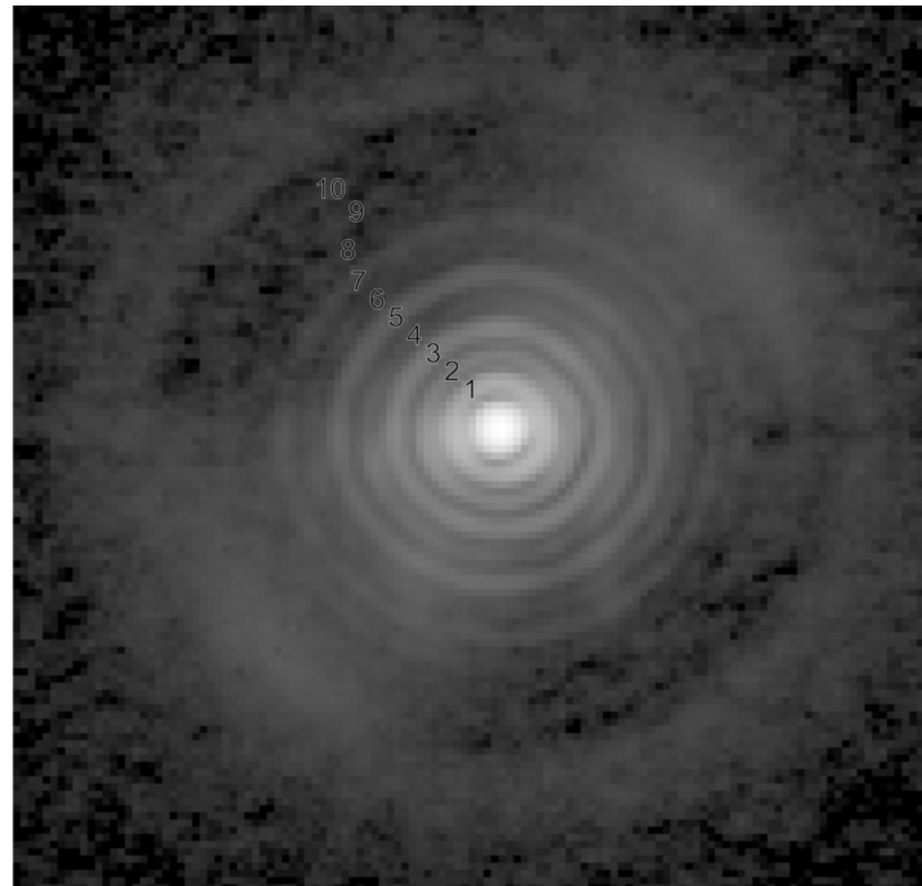
AB Aur disk imaged with HiCIAO

Higher quality AO correction (Extreme-AO) is becoming available, and will make efficient use of coronagraphs

(Coronagraphs will soon become very useful !!!)

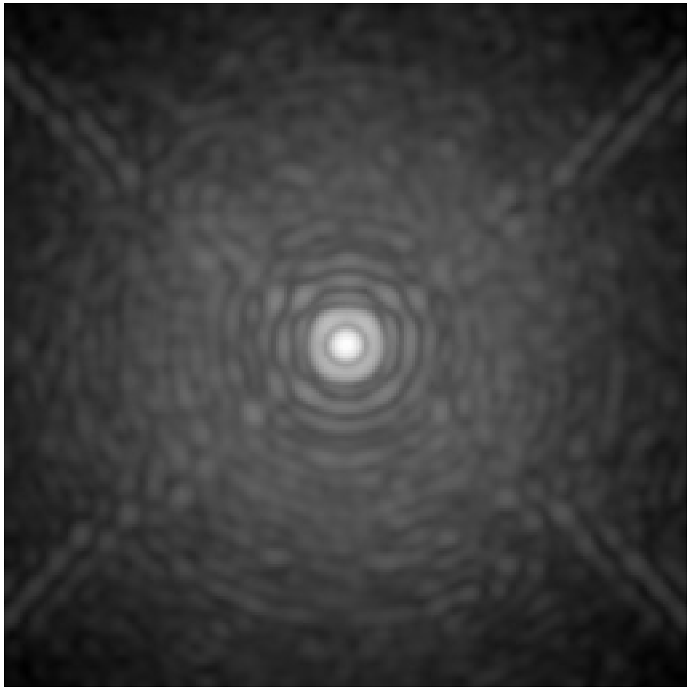


High quality PSF for HiCIAO (simulation)



LBT H-band PSF (~80% SR)

Higher quality AO correction (Extreme-AO) is not sufficient: **calibration and stability** are important



Removing bias in AO correction

It is essential to minimize slow and static speckles
that may look like planets.
There are the limiting factor in current systems
Fast atmospheric speckles are OK, as they average out very rapidly into a smooth PSF halo

Bias can come from non-common path error, bias in the wavefront sensor, uncorrected mode that is not detected by the WFS.

Calibrating residual speckle halo in the PSF

Using differential measurements, looking for:
- ways to change/modulate the speckle field without changing the planet image, or
- ways to change the planet image without changing the PSF halo

PSF calibration strategies

- **“classical” PSF subtraction: observe reference PSF, subtract**
- **Angular Differential Imaging: same-star subtraction thanks to sky rotation**
  - works well at large angular separations, where aberrations have large static component
  - poor performance close in to the star (not enough rotation)
- **Spectral differential imaging (dual band or IFS)**
  - Speckle calibration works well at large angular separations, where speckles are spectrally elongated, but does not work as well close in to the star
  - Science (spectra) + speckle calibration from same data
  - IFS is more powerful and efficient than simultaneous differential imaging
- **Coherent differential imaging: speckles are coherent with starlight, real sources are not**
  - Does not make any assumption about source (spectra, polarization)
  - combined wavefront sensing / PSF calibration
  - Works well very close to the star
  - works only within control radius of DM, requires good detector (fast readout favored)

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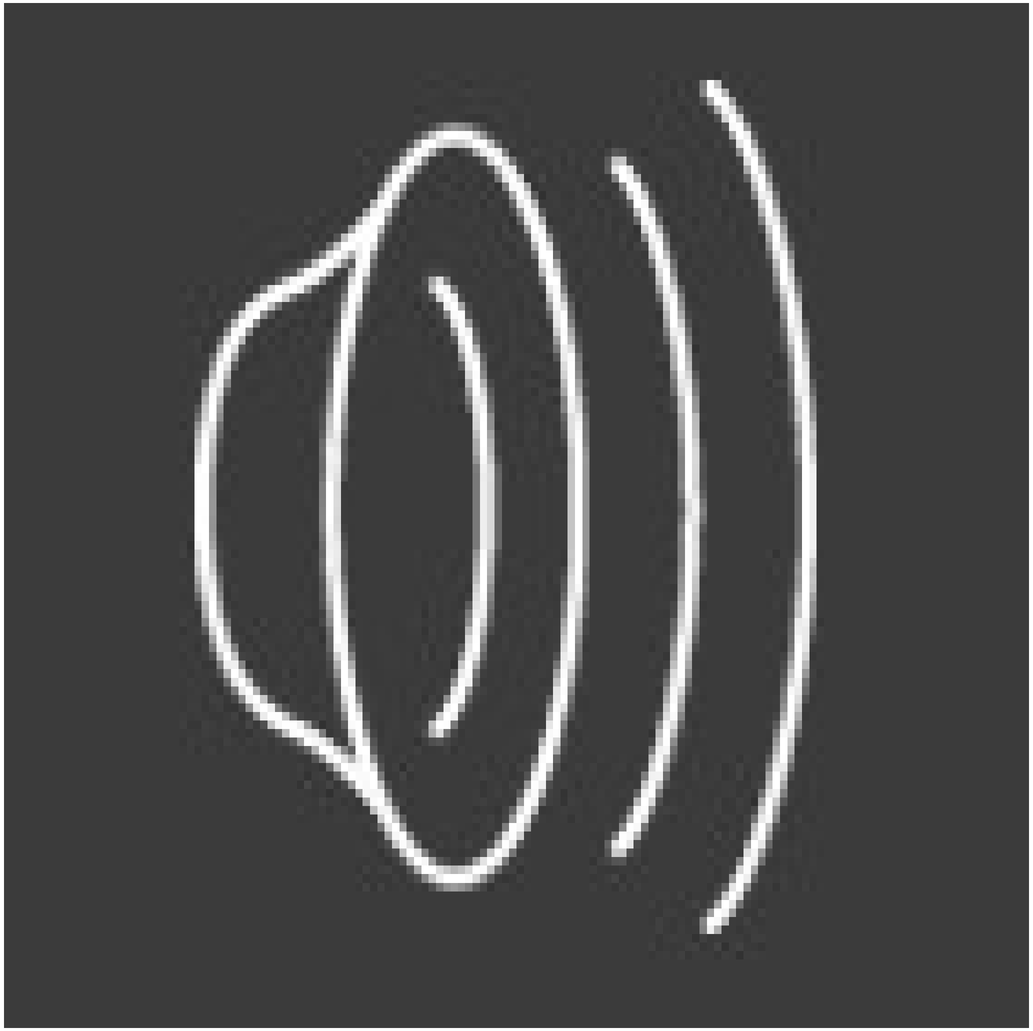
PSF calibration with Integral Field Spectrograph

Speckle spatial scale and contrast changes with lambda

IFS data can be used to fit speckle field image with a residual OPD error in the pupil plane

Real sources behave very differently :
    No change in position with lambda
    No interference with speckle field



Alcor IFS observation (courtesy of B. Oppenheimer)

From past/present to future instruments

“Conventional” AO imaging, ususally without coronagraph

First Systems designed for high contrast imaging, combining existing techniques

HiCIAO

NICI

Lyot coronagraph
Differential spectral imaging
Good AO system

Extreme-AO coronagraphic systems designed

PALM3000+P1640

GPI

SPHERE

SCExAO

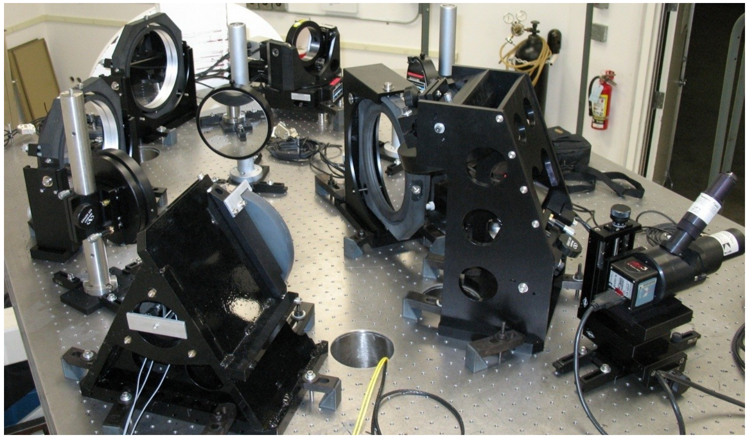
Better AO correction:
- More actuators
- Better WFSs

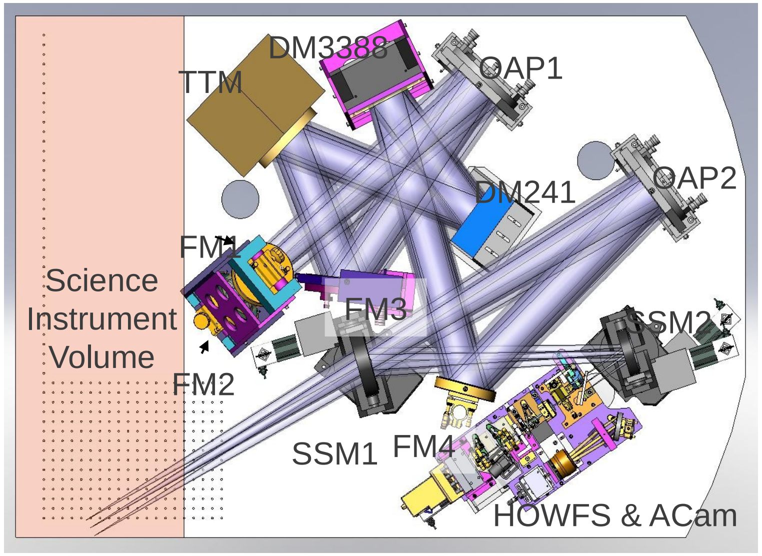
Improved PSF calibration:
- IFS, coherent PSF detection

Better coronagraphs:
- APLC, PIAA, 4 quadrant

The PALM-3000 upgrade to Palomar AO

- Science goals:
  - Comparative exoplanet and disk studies in the near-IR
  - Visible-light science at the diffraction-limit
- Deformable mirrors
  - 3388-actuator, 1µm stroke "tweeter"
  - 241-actuator, 5 µm stroke "woofer"
- Selectable WFS pupil sampling
  - 63x63, 32x32, 16x16, or 8x8
  - Up to 2kHz WFS frame rate
- Science instruments:
  - Project 1640 near-IR APLC coronagraphic IFS & CAL (AMNH & JPL)
  - PHARO near-IR Lyot and vortex coronagraph / spectrograph / imager (Cornell & JPL)
  - SWIFT visible-light IFS (Oxford)
  - Visible light imager and 'red' vortex coronagraph (Caltech & JPL)
  - Future option: Echelle fiber feed (Caltech)



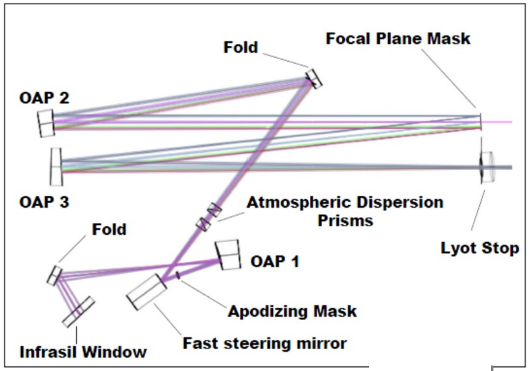


PI: Richard Dekany

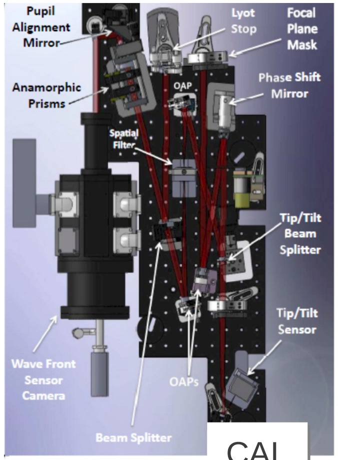
Project 1640

(AMNH, CIT, JPL: Ben Oppenheimer, PI)

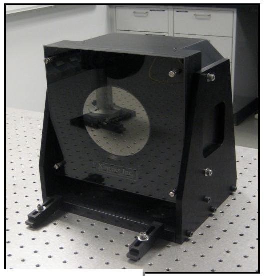
- Palomar 5m Telescope
- Apodized pupil Lyot coronagraph
- 200 x 200 spaxel JH IFS
  - R ~ 30; 20 mas / pixel
- < 5 nm RMS CAL
  - Currently 15 nm RMS
- Performance goals
  - 10^{-6} raw contrast (SNR = 1, 30s)
  - 10^{-7} post-LOCI (SNR = 5, 30m)
- Schedule
  - First high-order lock: July 2011
  - 99-night Survey 2011-2014



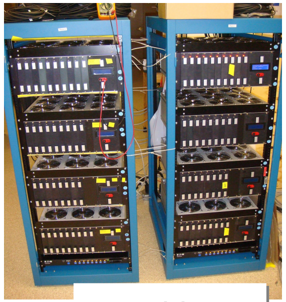
APLC



CAL

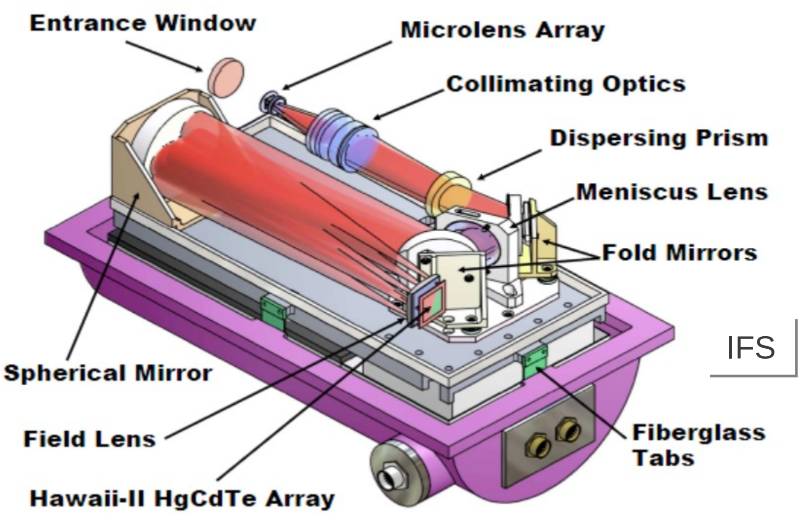


DM3388

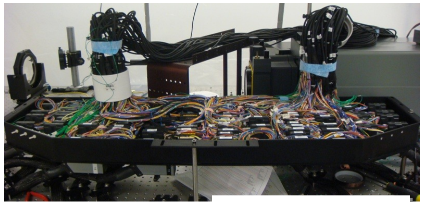


DM drivers

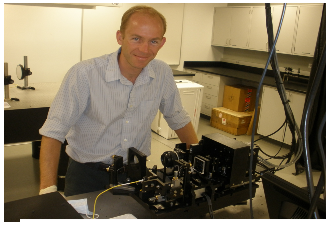
RTC



IFS



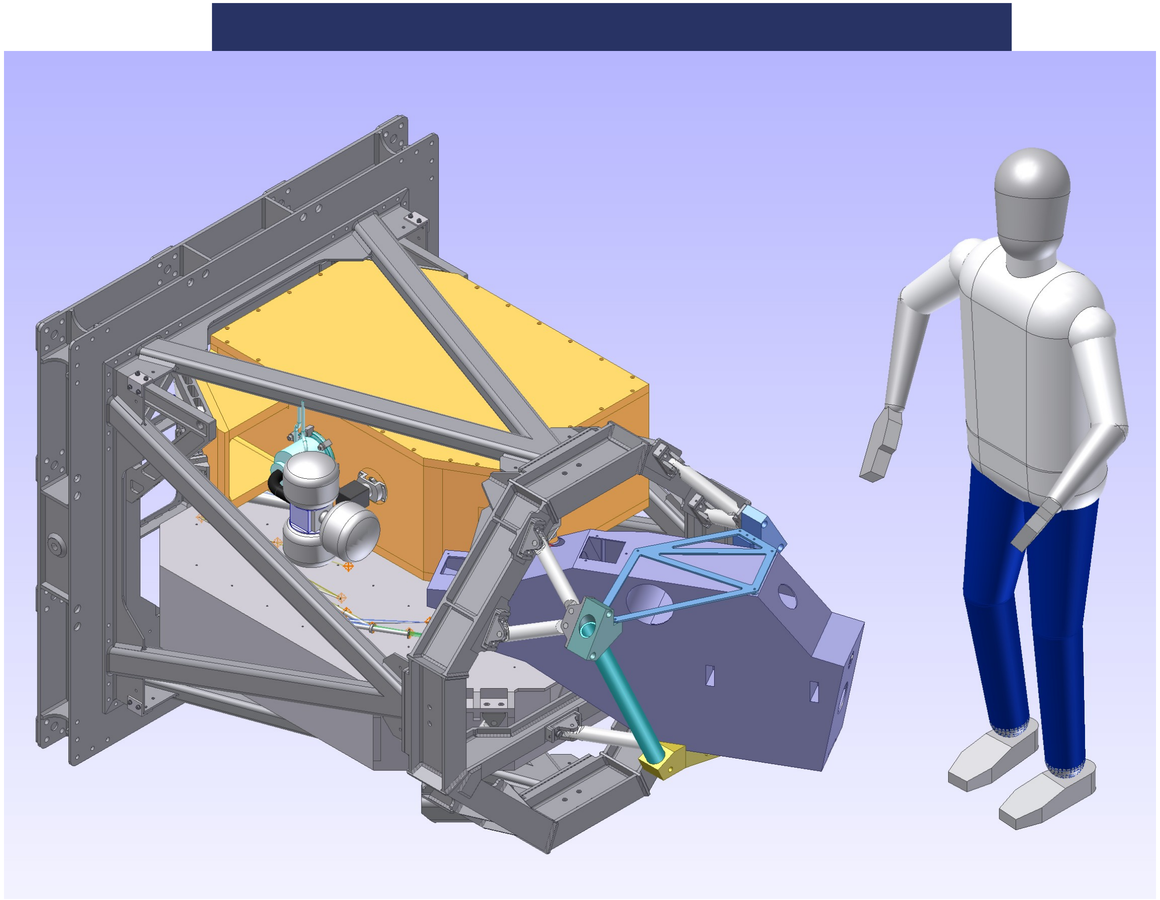
DM cabling

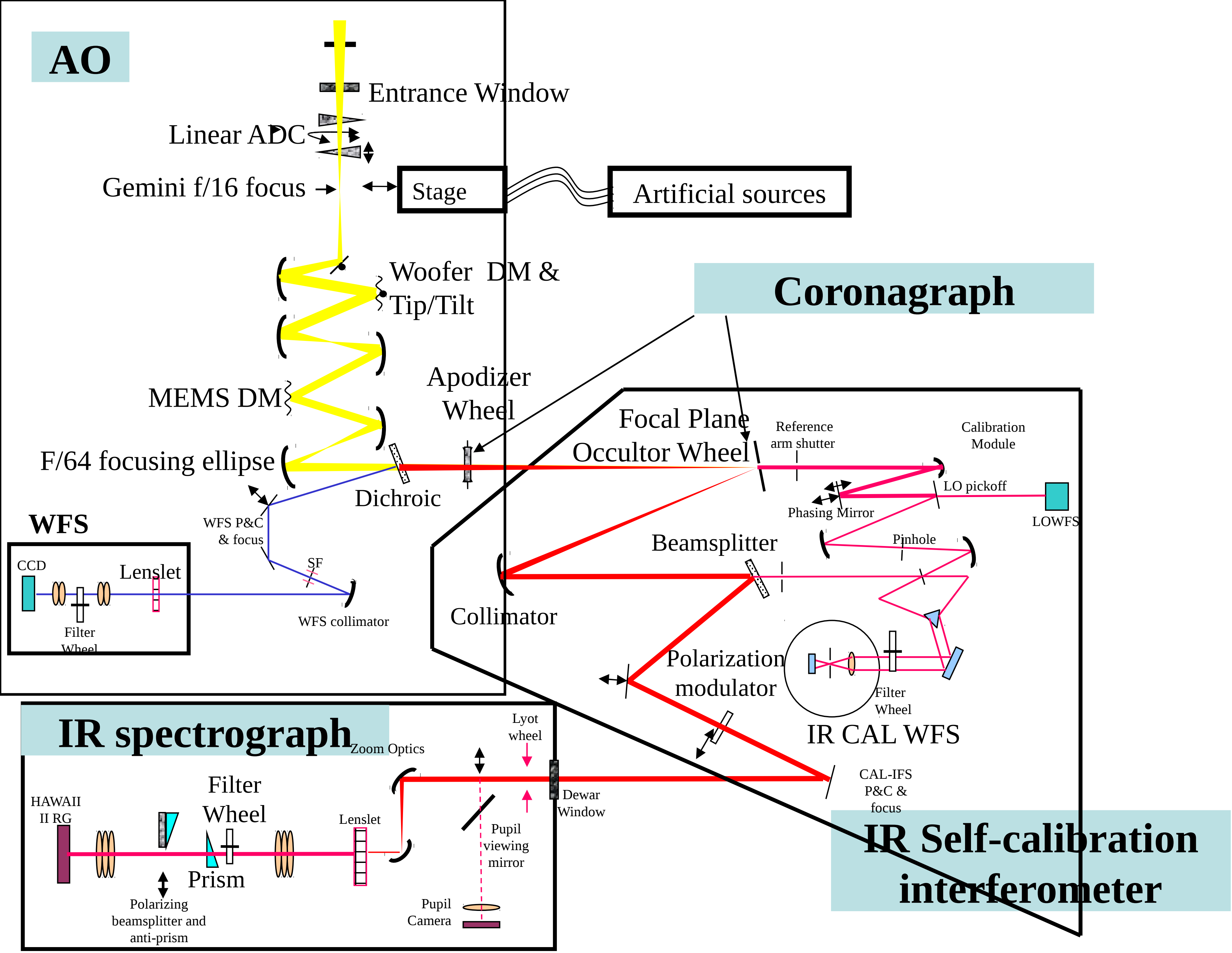


Antonin & the HOWFS

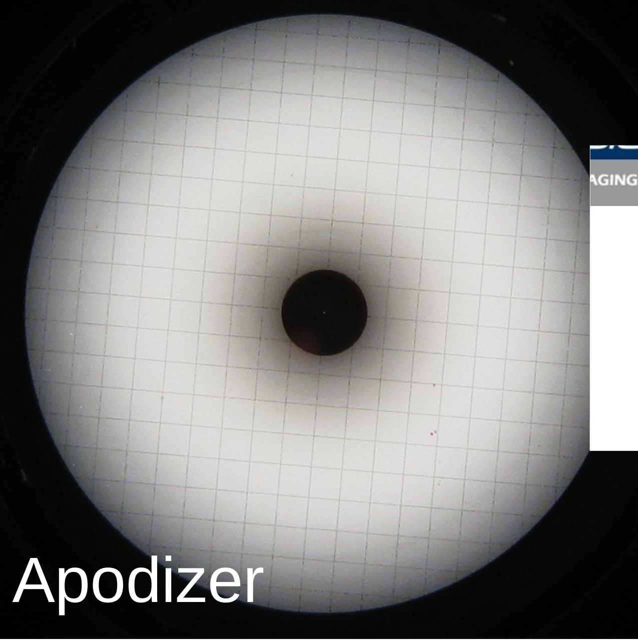
Gemini Planet Imager (GPI)

- 44x44 MEMS AO system
- Static speckles minimized
  - Precision optics
  - Daytime speckle nulling calibration
  - On-sky IR interferometer WFS
- APLC
- IR integral field spectrograph



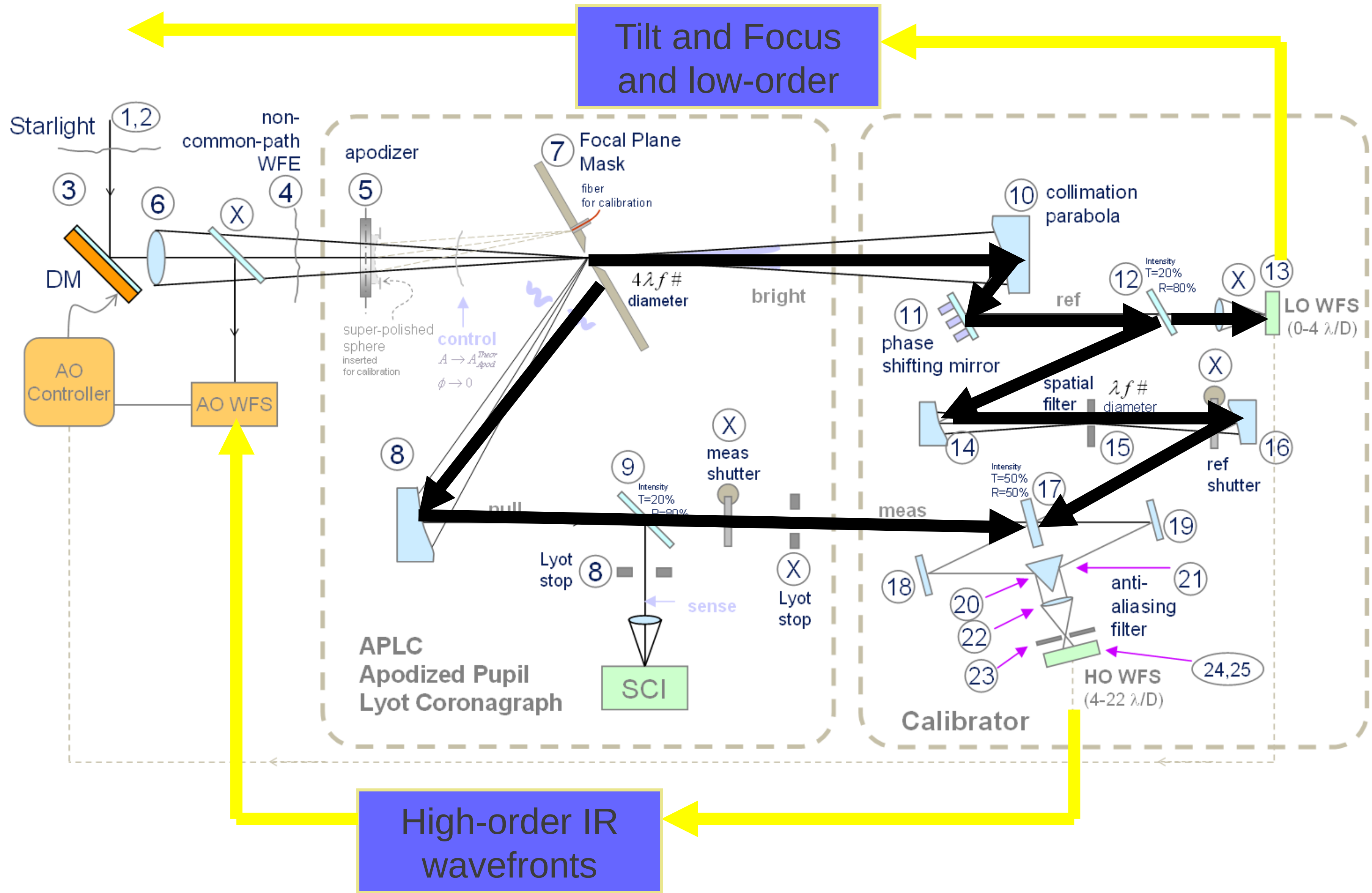

AO
Entrance Window
Linear ADC
Gemini f/16 focus
Stage
Artificial sources
Woofer DM &
Tip/Tilt
MEMS DM
Apodizer
Wheel
F/64 focusing ellipse
Dichroic
WFS
WFS P&C
& focus
SF
CCD
Lenslet
Filter
Wheel
WFS collimator
Coronagraph
Focal Plane
Occultor Wheel
Reference
arm shutter
Calibration
Module
LO pickoff
Phasing Mirror
LOWFS
Beamsplitter
Pinhole
Collimator
Polarization
modulator
Filter
Wheel
IR CAL WFS
IR spectrograph
Zoom Optics
Lyot
wheel
Dewar
Window
CAL-IFS
P&C &
focus
HAWAII
II RG
Filter
Wheel
Prism
Lenslet
Pupil
viewing
mirror
Polarizing
beamsplitter and
anti-prism
Pupil
Camera
IR Self-calibration
interferometer

APLC coronagraph masks

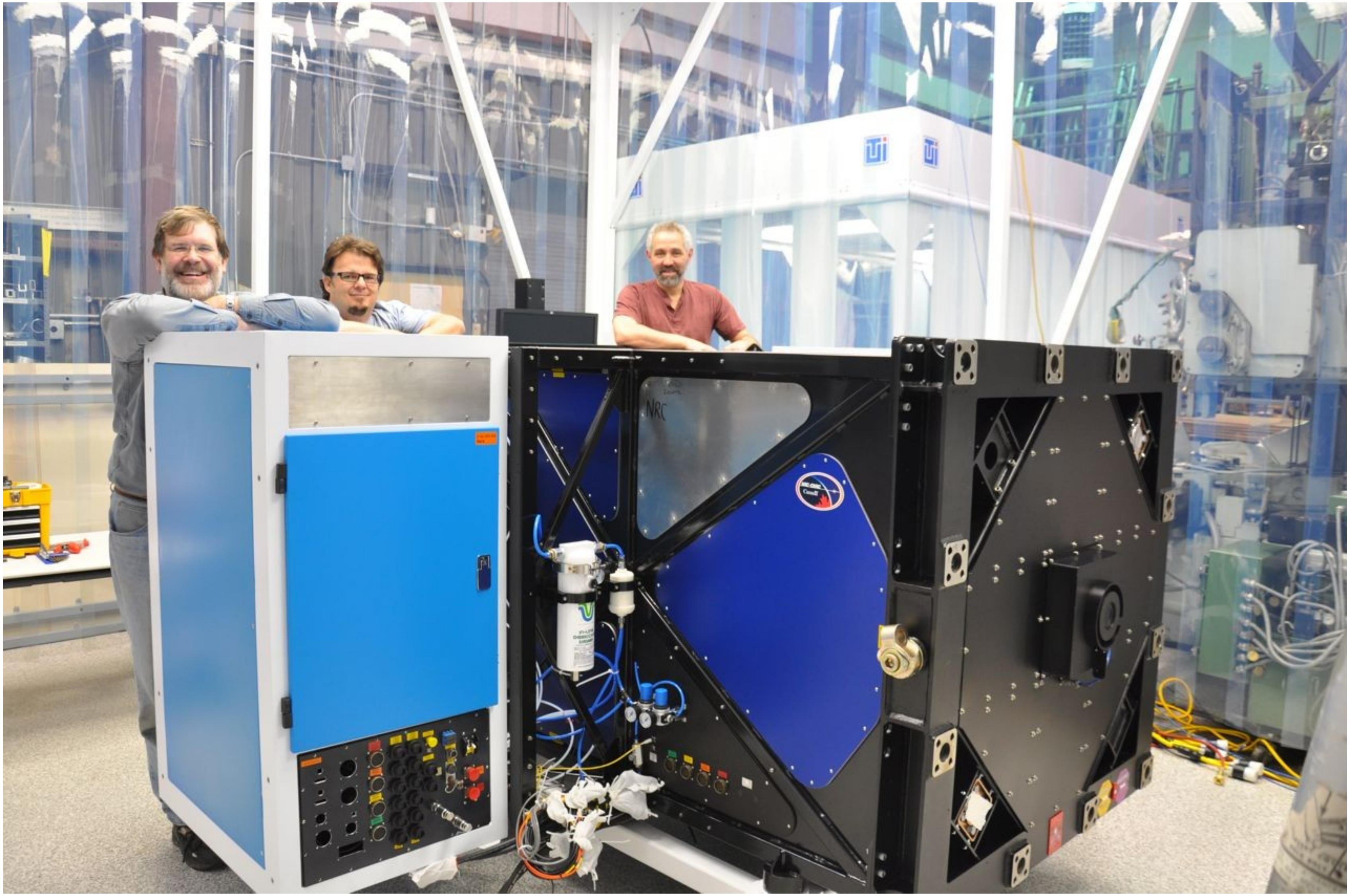


G3

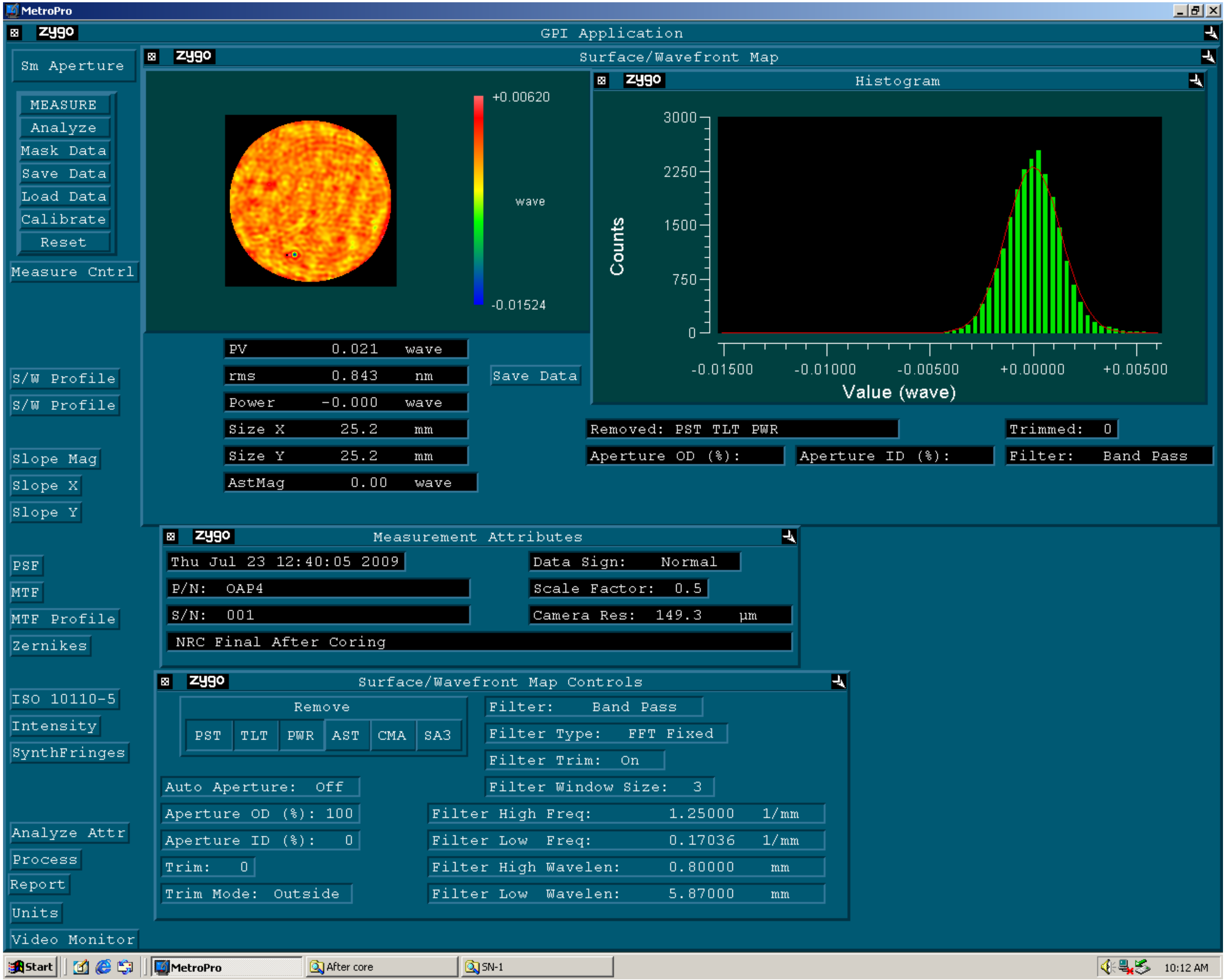
Calibration interferometer (JPL) measures slow aberrations to nanometer accuracy



UCSC clean room
December 18 2010

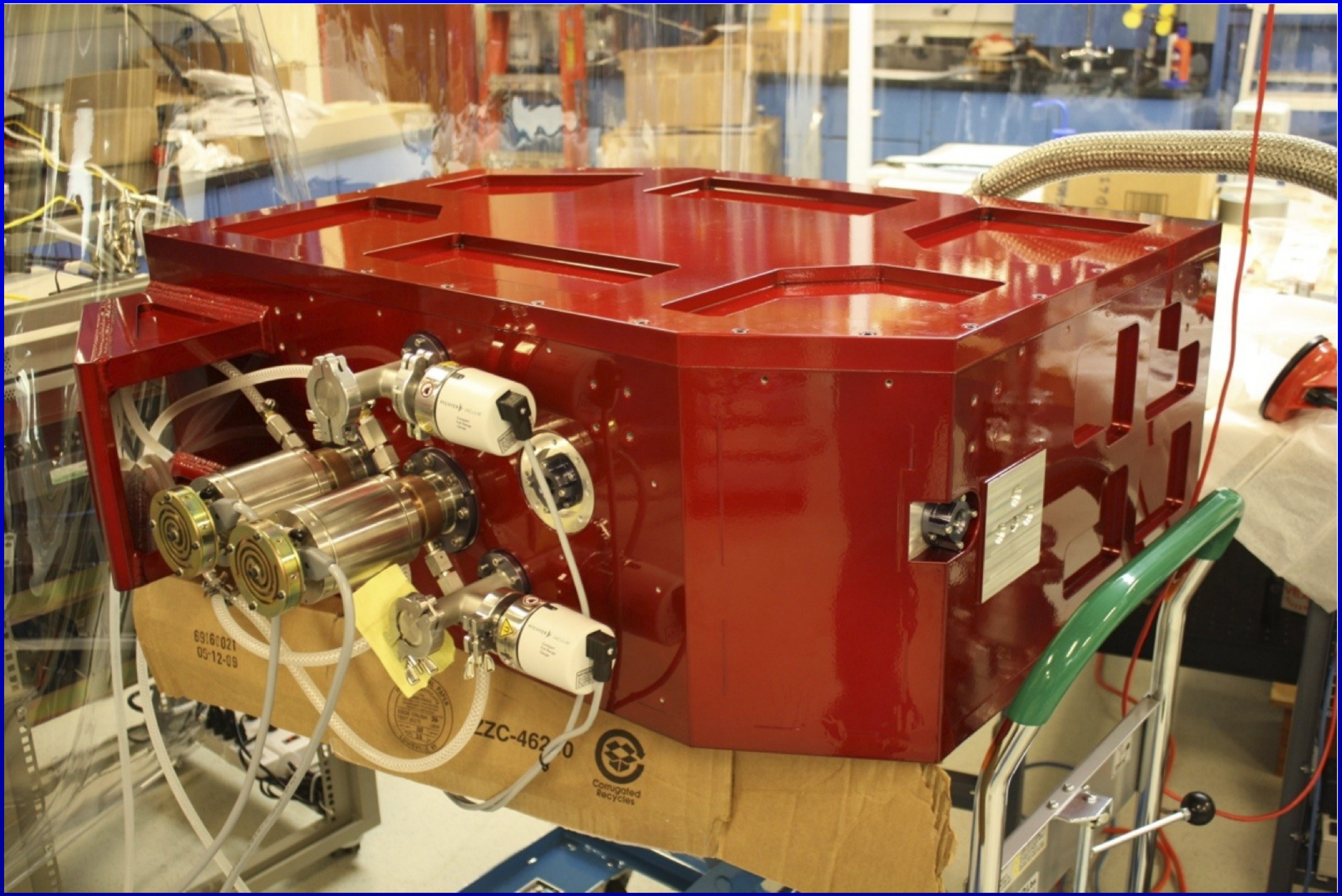


Individual optics < 1 nm

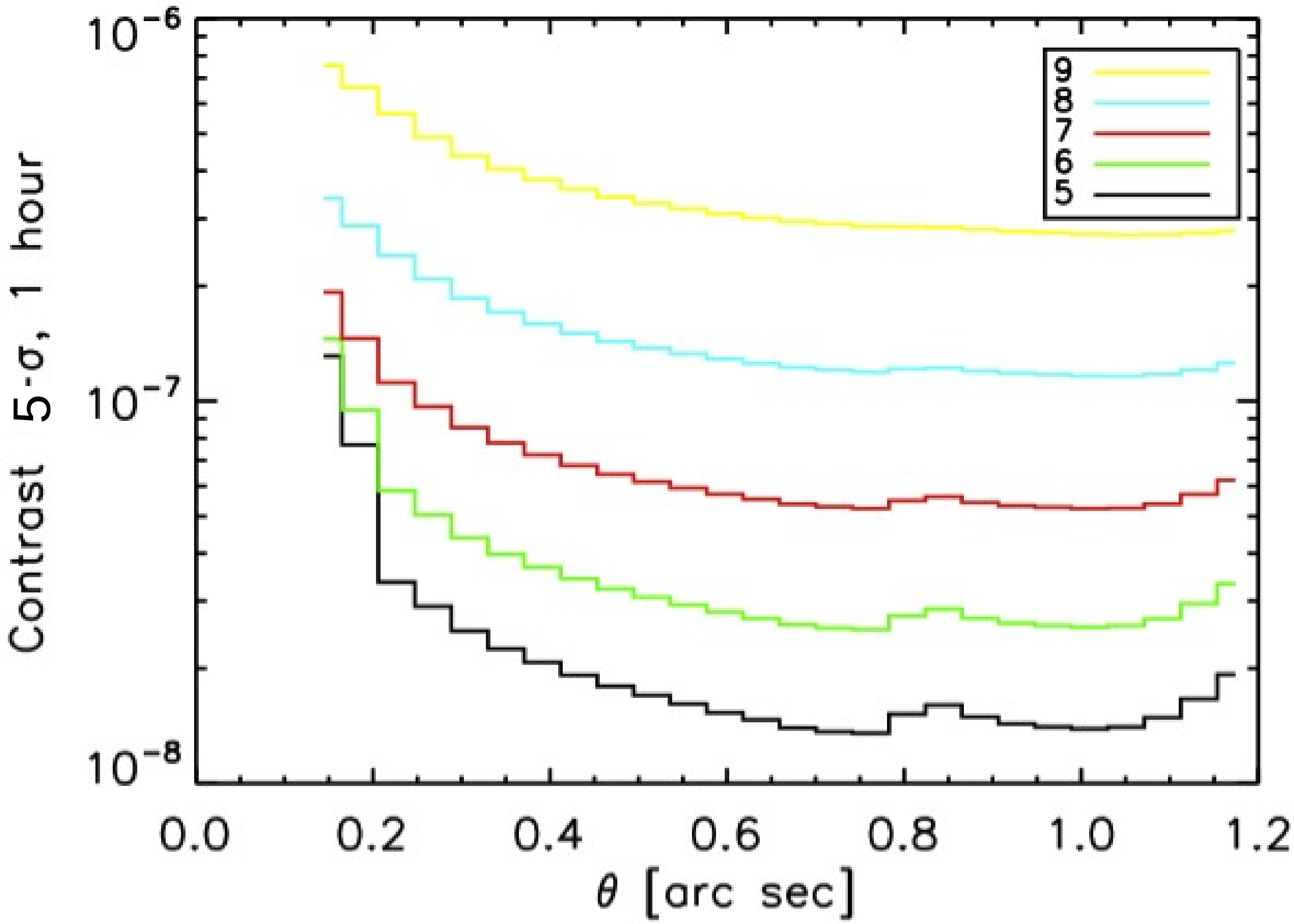


UCLA
U Montreal,
Immervision

Integral field spectrograph



Performance models now on http://planetimgaer.org



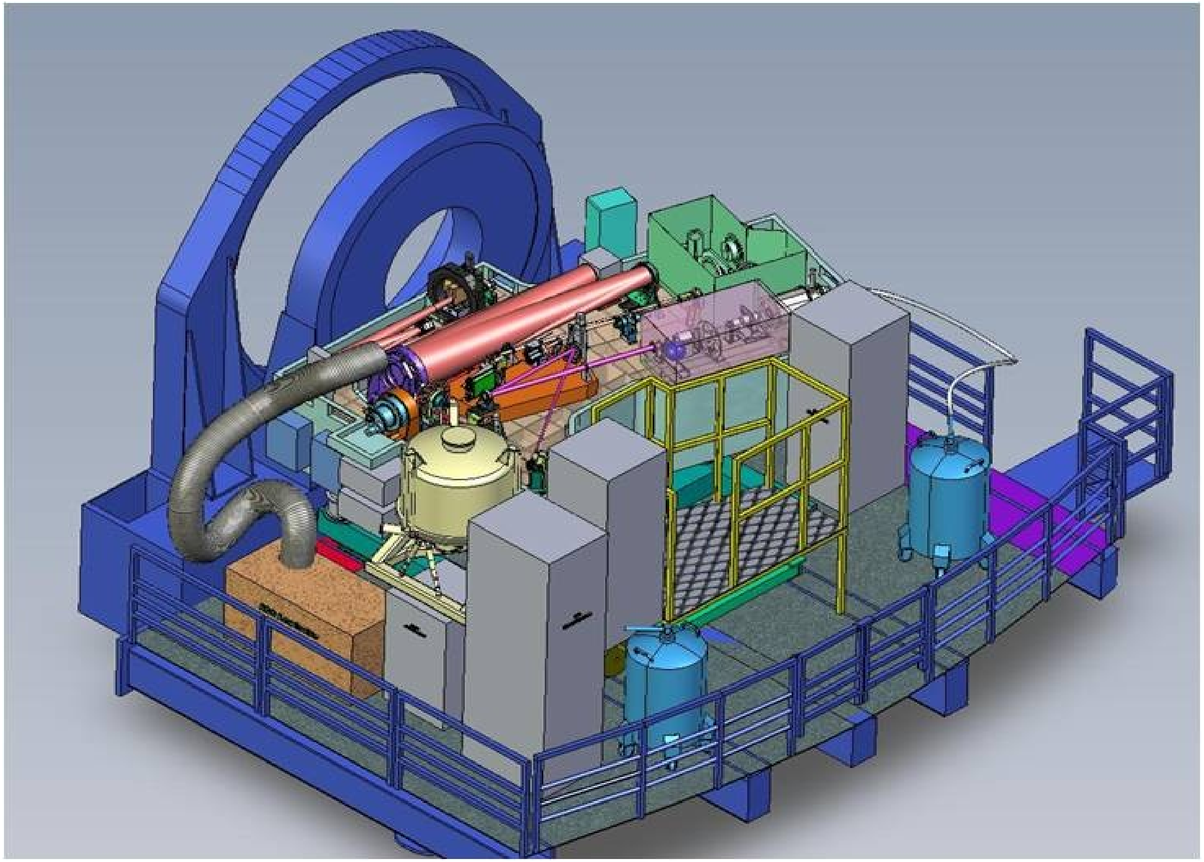
Spectro-Polarimetric High Contrast Exoplanet REsearch (SPHERE)

High order AO correction (SAXO)
 41X41 actuators DM
 ~90% Strehl in H band

Dual band near-IR imaging (IRDIS)
 <10 nm differential error

Integral Field Spectrograph (IFS)

Visible differential polarimetric imager (ZIMPOL)
 Direct imaging of disks and evolved planets (reflected light)



Large survey will start in ~2012
 Schedule and science goals similar to GPI

SPHERE & GPI:
high performance systems
+ large telescopes in good sites
+ large observing programs

Statistics: SPHERE & GPI will observe a large sample of stars to few MJ sensitivity

→ good planet statistics for >MJ planets at >5AU

Complementary to RV and transits, which favor close in planets

Will provide strong test for planetary formation theories (formation + migration)

Sample will be sufficiently large for comparative planetary system architecture studies

How do planetary systems change as function of stellar mass ?

Characterization: SPHERE & GPI will obtain spectra of exoplanets

→ contrain planetary atmosphere models, cooling rate → better mass estimates

The Subaru Coronagraphic Extreme-AO (SCExAO) system

O. Guyon, F. Martinache, V. Garrel, C. Clergeon, J. Morino, T. Kudo, P. Stewart, T. Groff

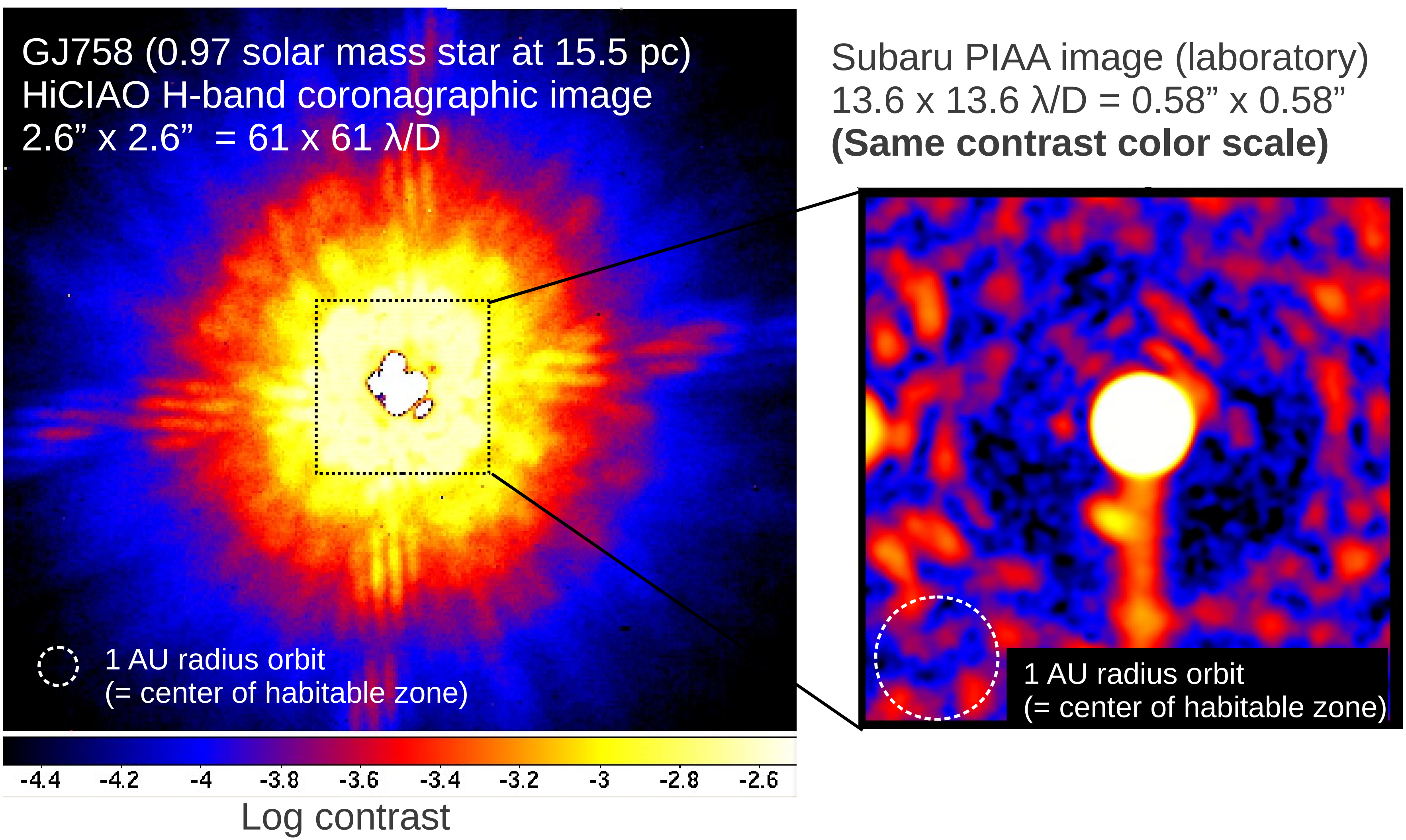
1 λ/D PIAA coronagraph
- NIR focal plane WF control/calibration
- Coronagraphic low order WFS
- ExAO-optimized visible WFS

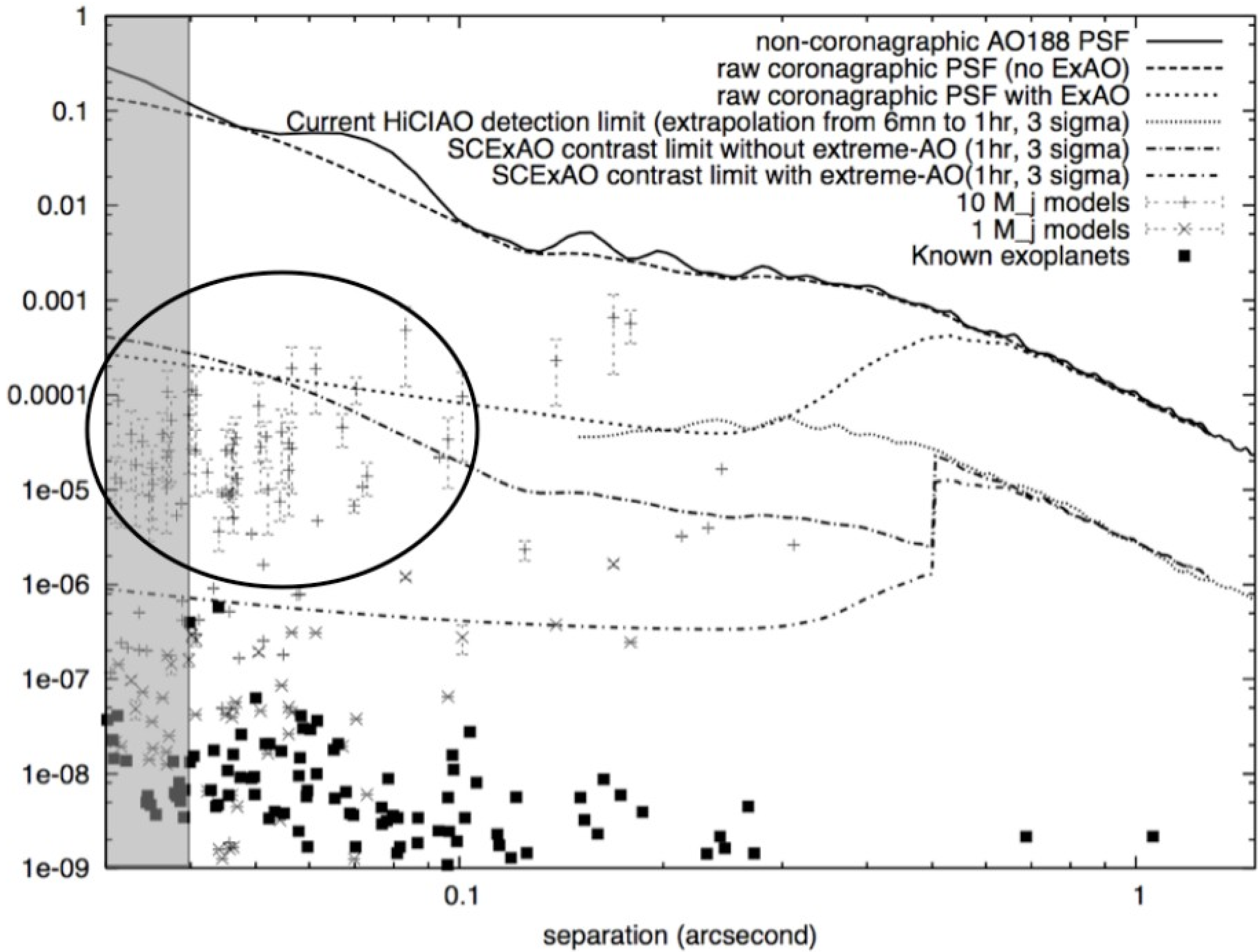
High contrast imaging at small angular separation is scientifically extremely valuable:
- allows sytem to probe **inner parts of young planetary systems** (<10 AU)
- constrain planet formation in the **habitable zone** of stars
- **direct imaging** of reflected light planets may be possible (reflected flux goes as a^{-2})

Designed as a highly flexible, evolvable platform
→ uses new (more risky) techniques than GPI and SPHERE
Efficient use of AO188 system & HiCIAO camera
Ongoing telescope engineering runs (Feb 2011, Jul 2011)

High contrast imaging in lab reaches much higher performance than what is currently achieved on-sky: newer technologies, more stable environment, better calibrations

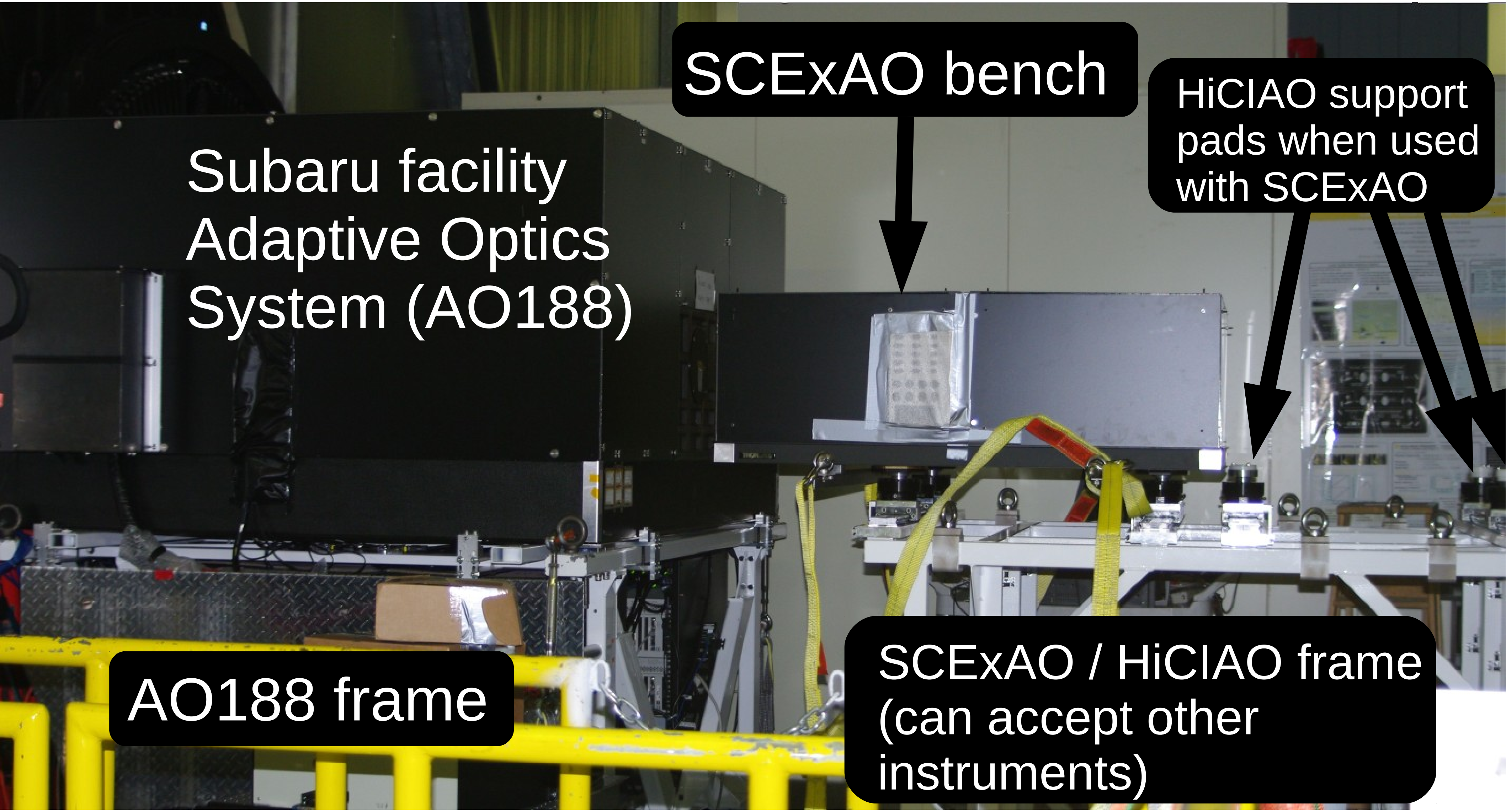
SCExAO's goal is to deploy on the telescope new techniques which have been demonstrated in the lab to offer high performance, and to create the conditions necessary to achieve this high performance



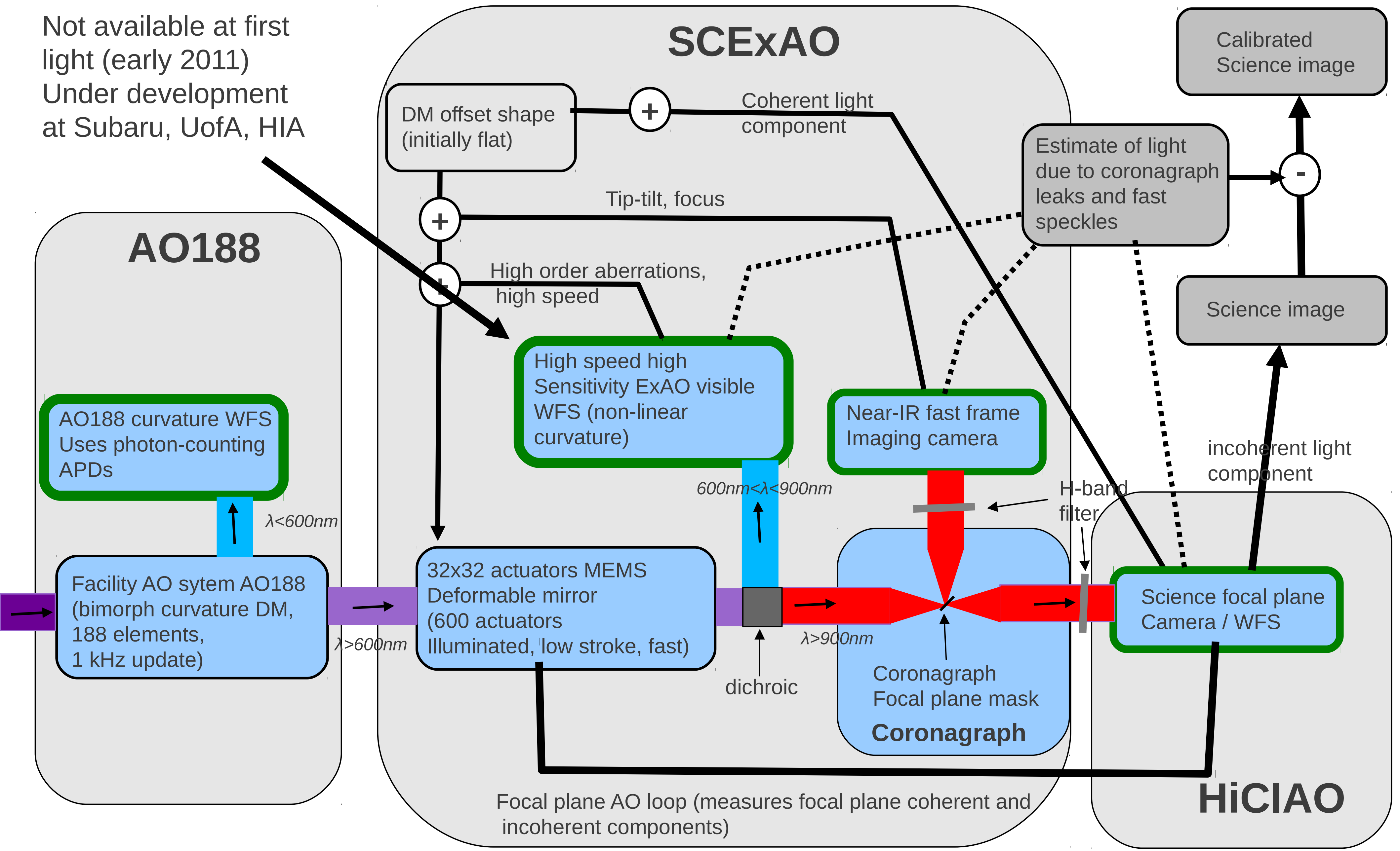


non-coronagraphic AO188 PSF
raw coronagraphic PSF (no ExAO)
raw coronagraphic PSF with ExAO
Current HiCIAO detection limit (extrapolation from 6mn to 1hr, 3 sigma)
SCExAO contrast limit without extreme-AO (1hr, 3 sigma)
SCExAO contrast limit with extreme-AO(1hr, 3 sigma)
10 M_j models
1 M_j models
Known exoplanets
1
0.1
0.01
0.001
0.0001
1e-05
1e-06
1e-07
1e-08
1e-09
0.1
1
separation (arcsecond)

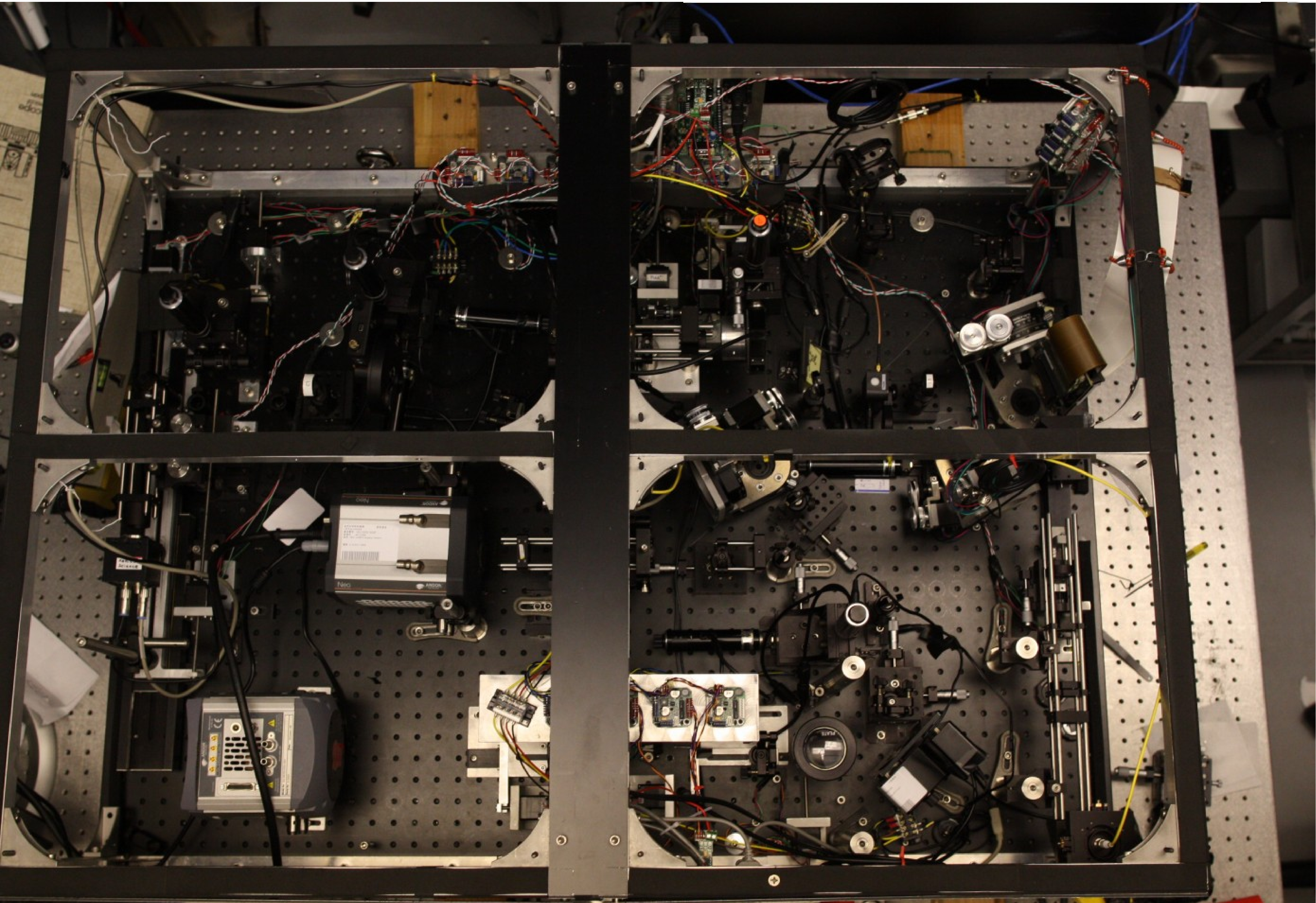
SCExAO at Subaru Telescope (Aug 2010)



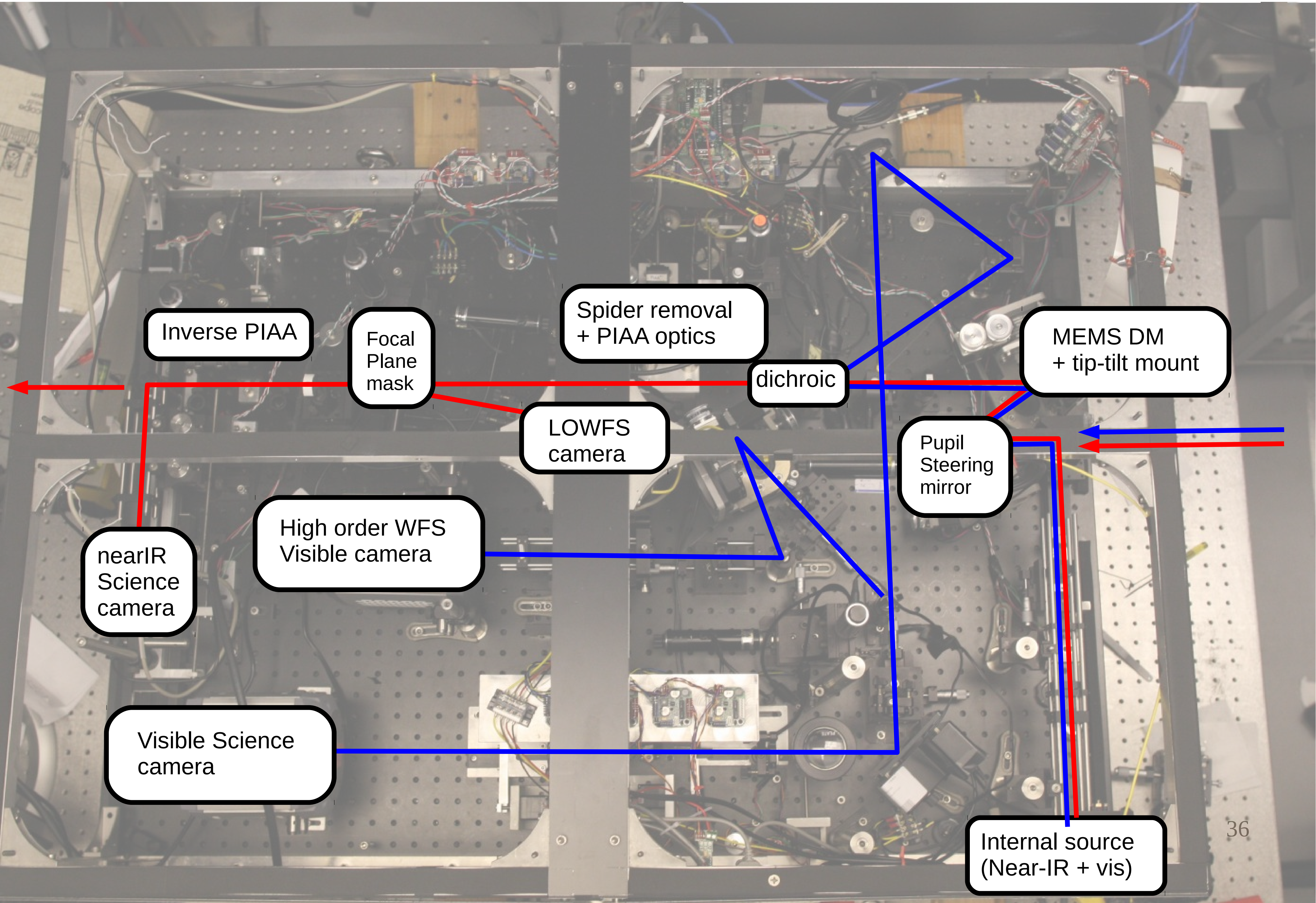
SCExAO Wavefront Control architecture and speckle calibration



The Subaru Coronagraphic Extreme-AO (SCExAO) system



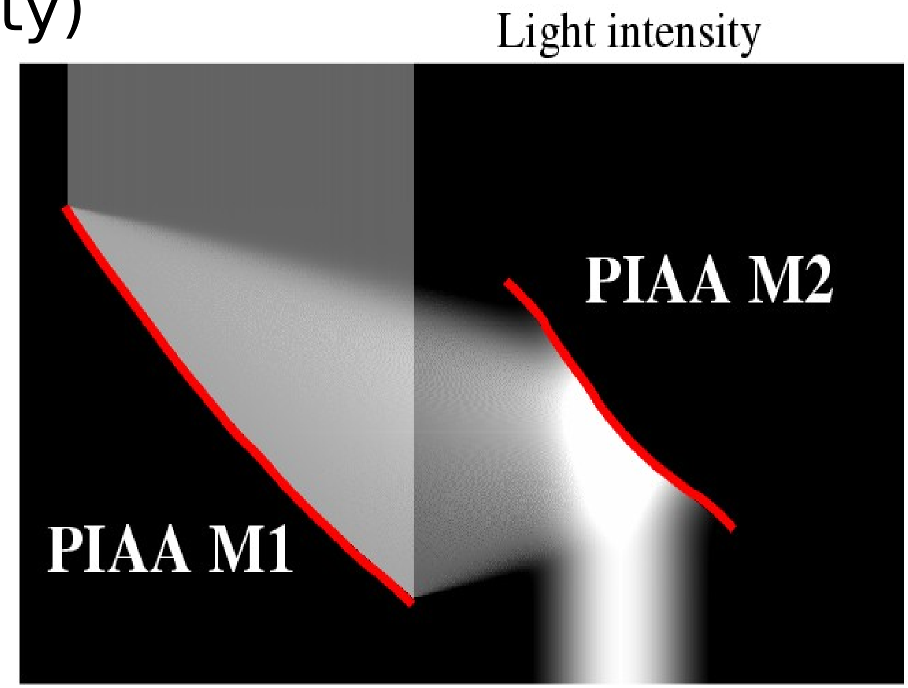
The Subaru Coronagraphic Extreme-AO (SCExAO) system



Phase-Induced Amplitude Apodization (PIAA) coronagraph

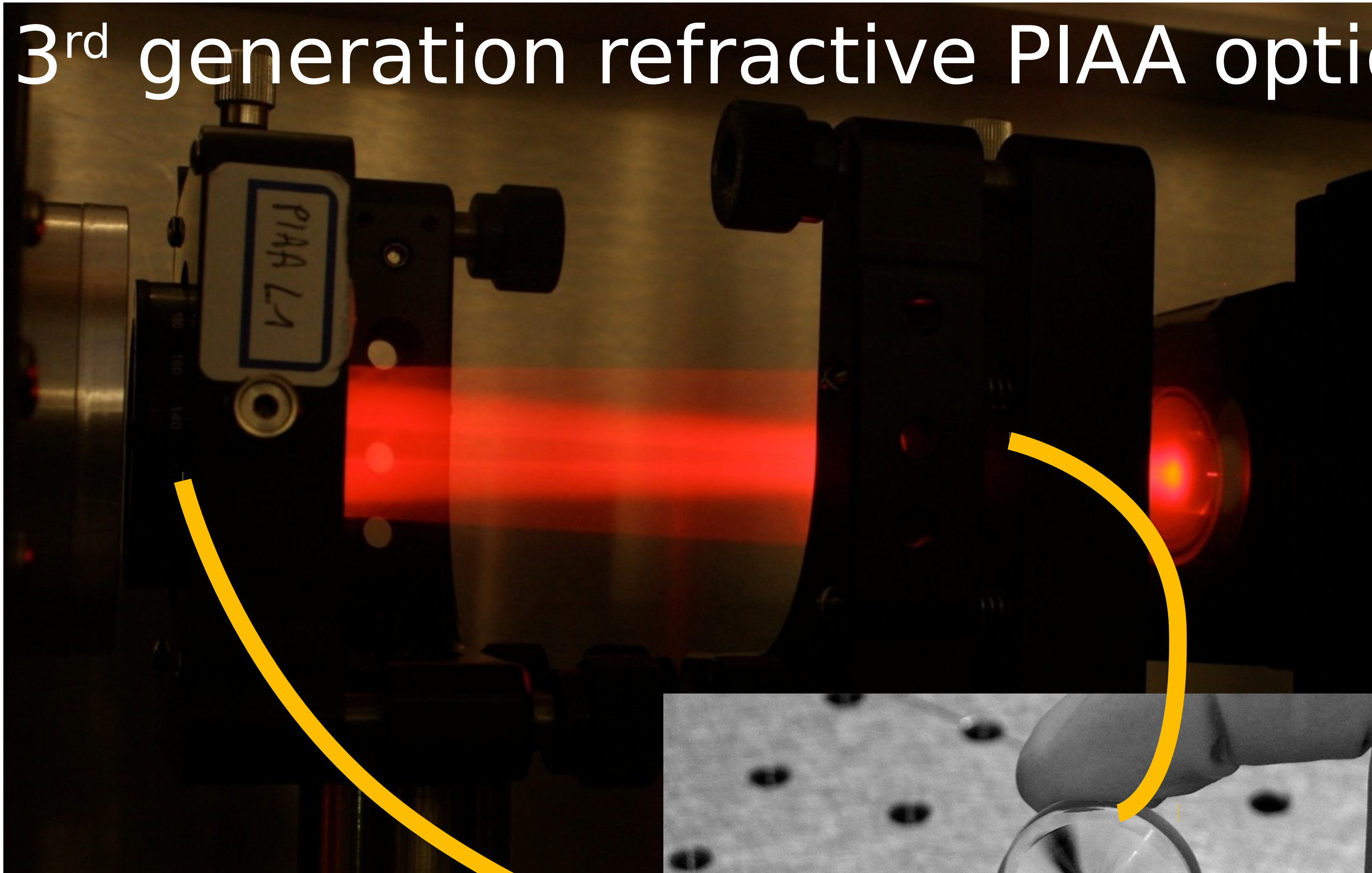
Utilizes lossless beam apodization with aspheric optics (mirrors or lenses) to concentrate starlight in single diffraction peak (no Airy rings).

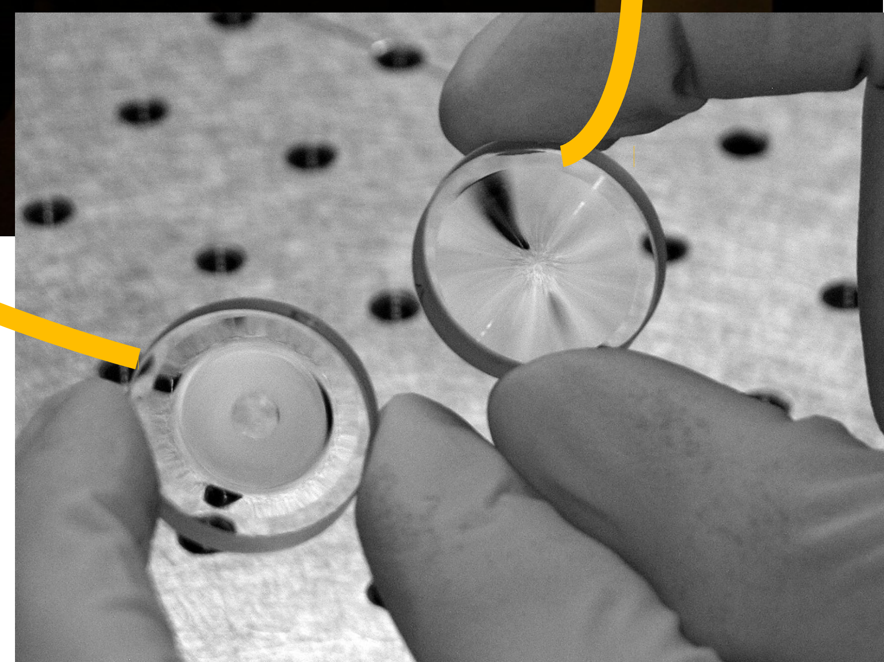
- high contrast (limited by WF quality)
- Nearly 100% throughput
- IWA 0.64 λ/D to 2 λ/D
- 100% search area
- no loss in angular resol.
- can remove central obsc. and spiders
- achromatic (with mirrors)



Refs: Guyon, Pluzhnik, Vanderbei, Traub, Martinache ... 2003-present

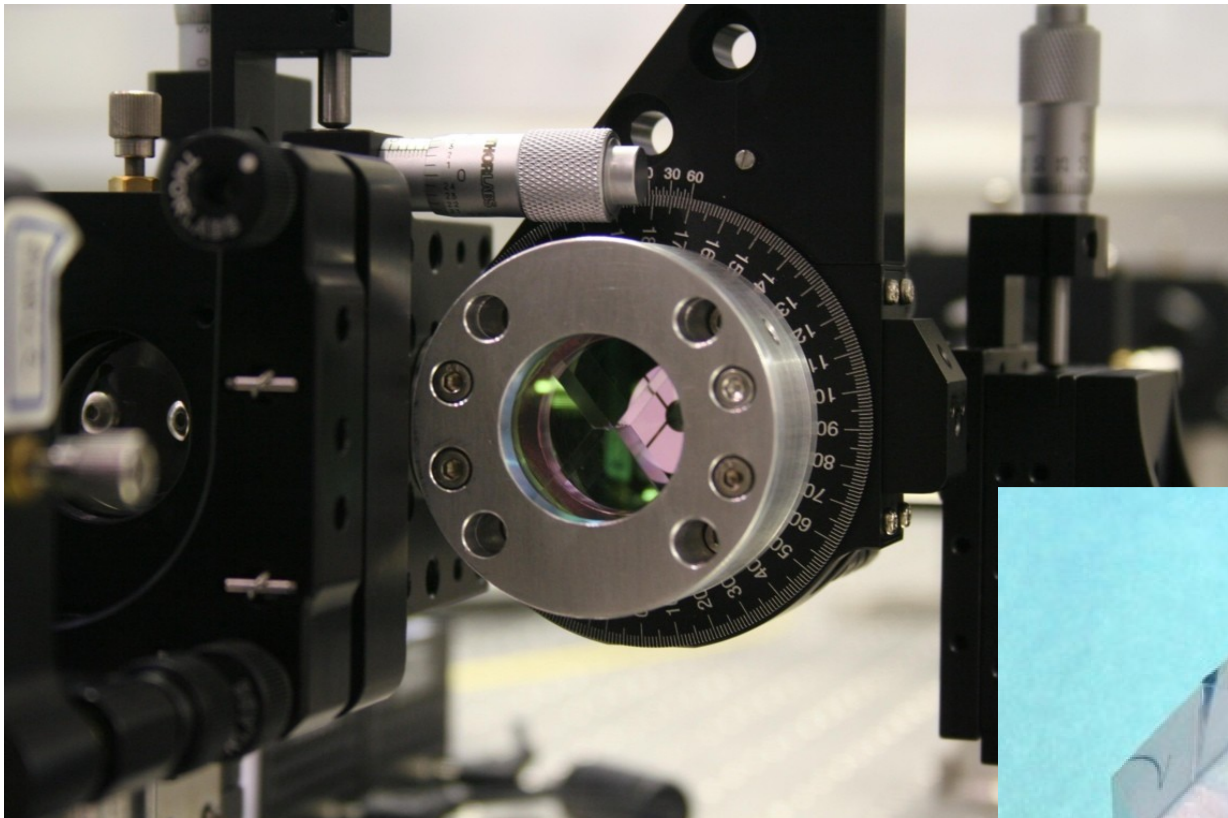
3rd generation refractive PIAA optic

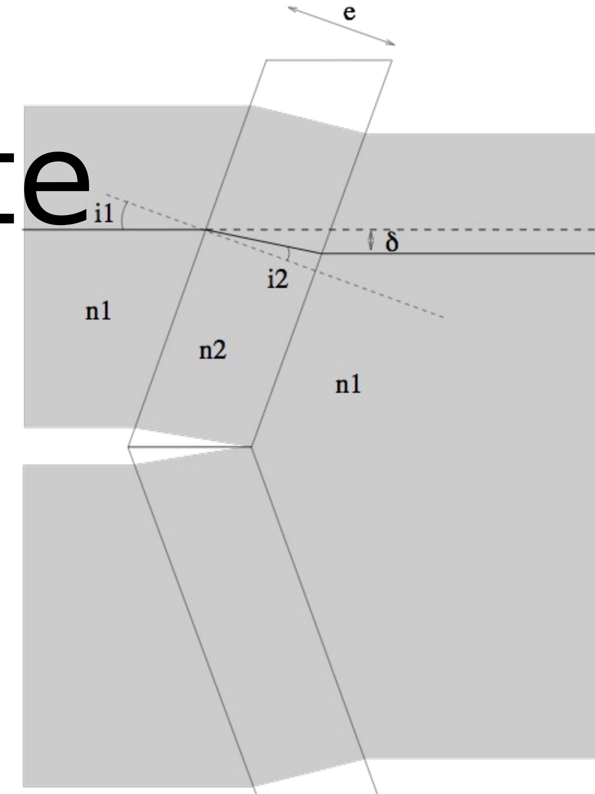


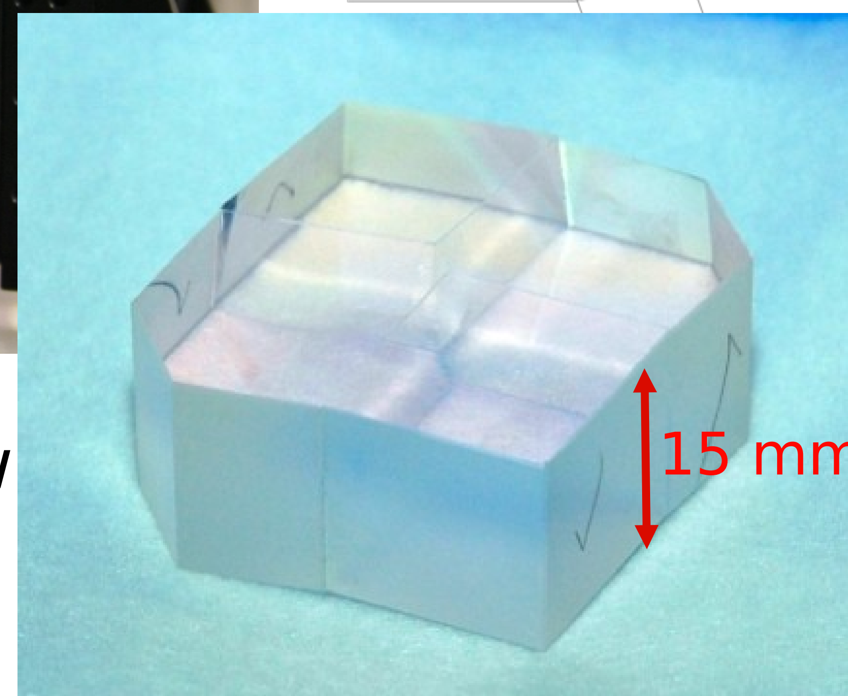


- On-axis lenses
- Lenses are 96 mm apart
- Apodize the beam
- Remove the central obscuration

Spider Removal Plate

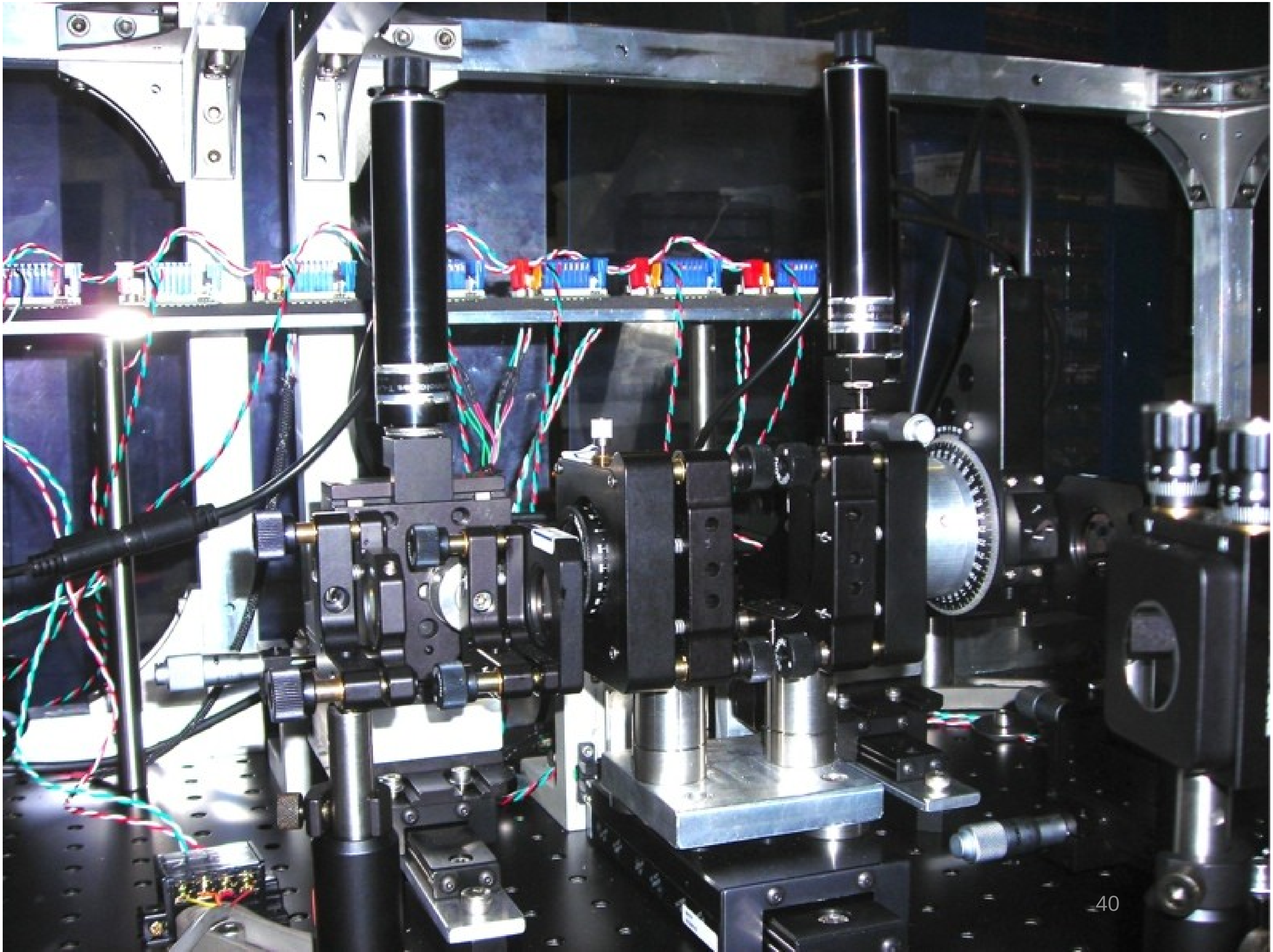


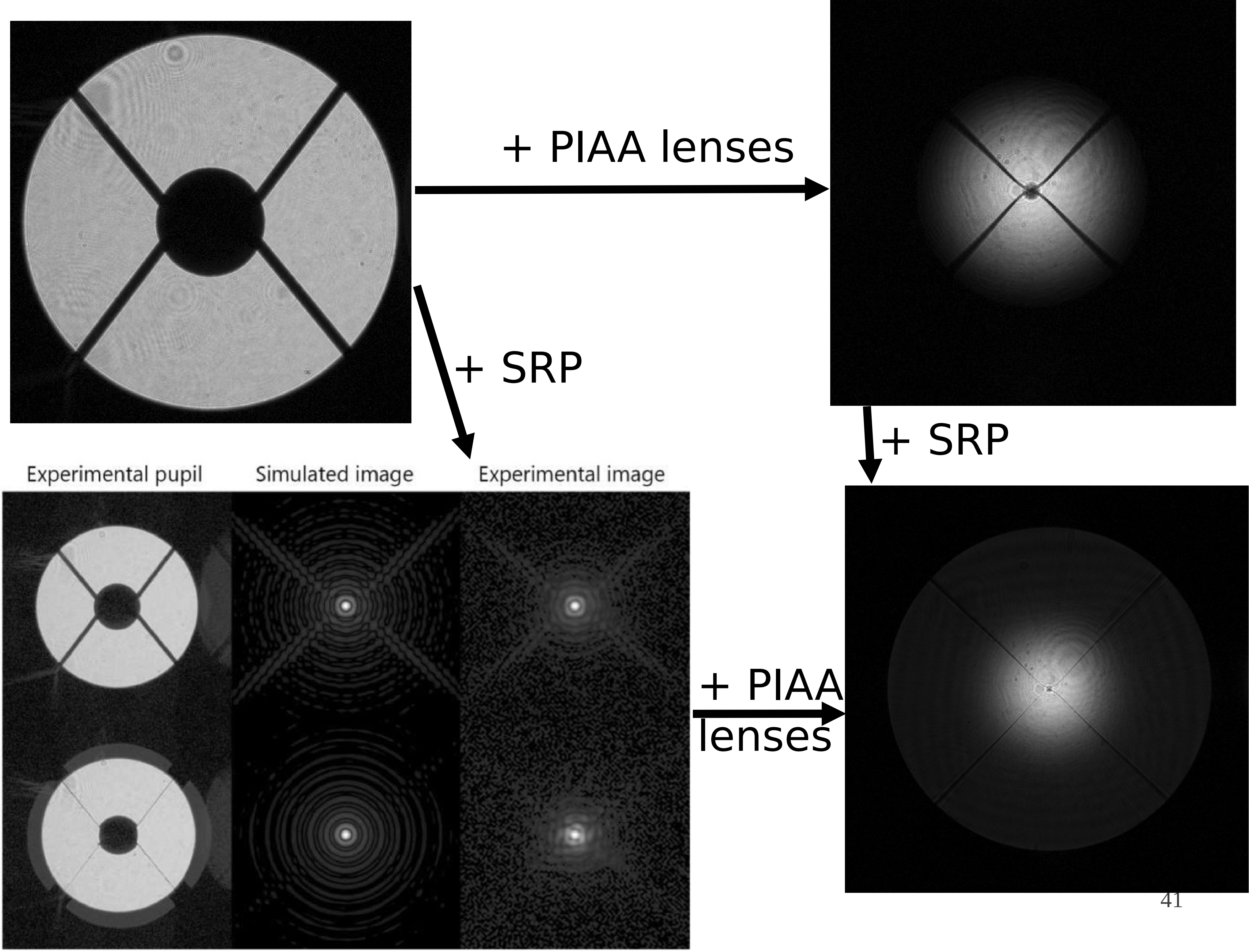




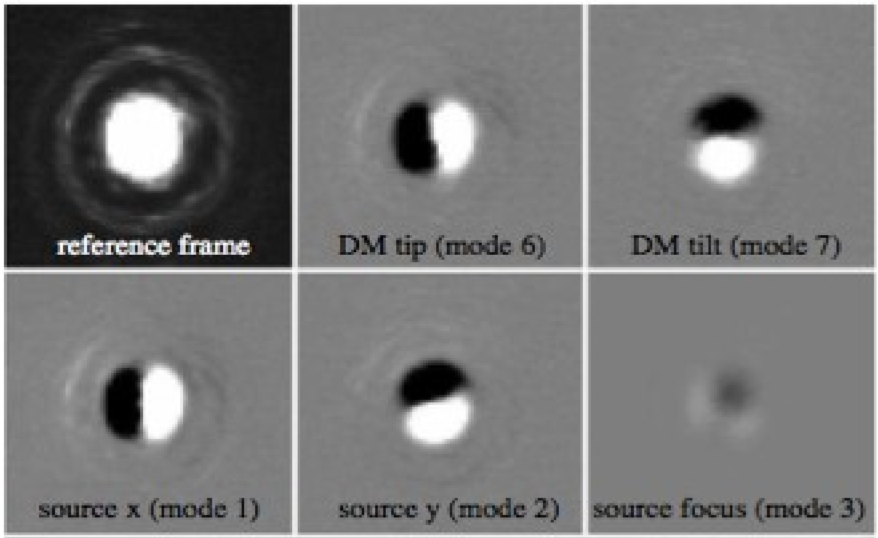
- 15 mm thick precision window
- Fused Silica
- Tilt angle: 5 +/- 0.02°

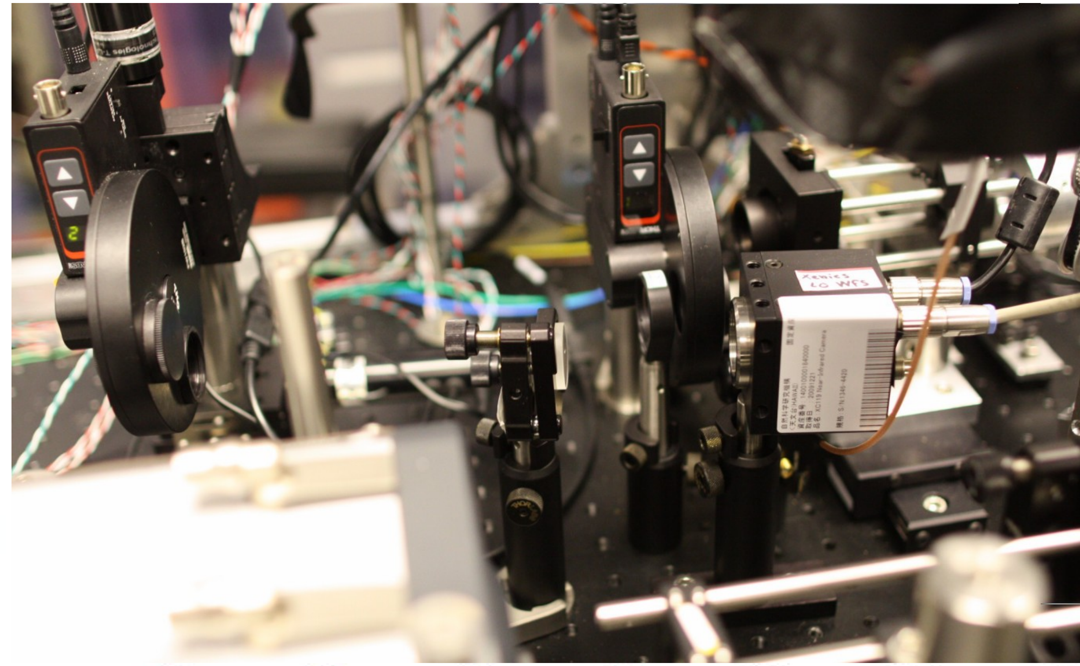
Beam shaping hardware



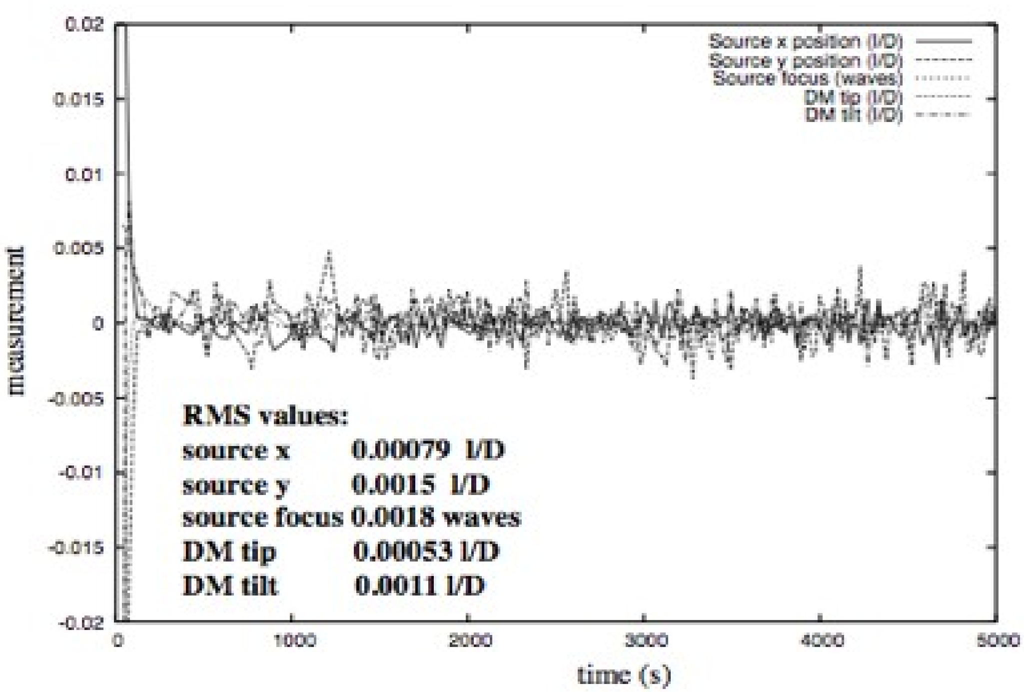

+ PIAA lenses
+ SRP
+ SRP
Experimental pupil
Simulated image
Experimental image
Without SRP
With SRP
+ PIAA lenses

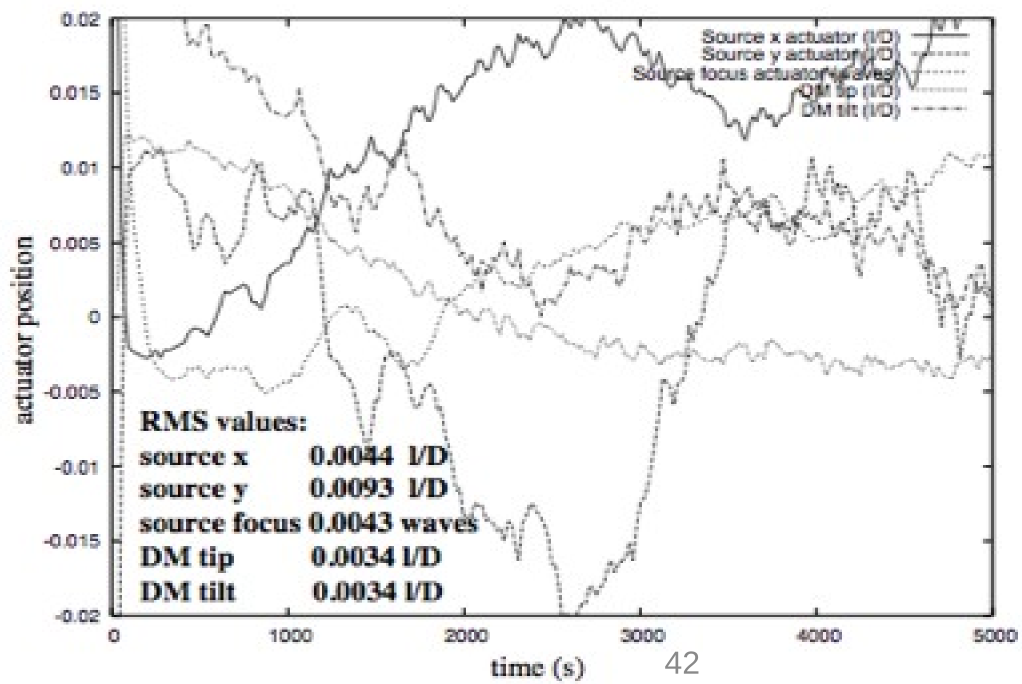
CLOWFS pointing control demonstrated to 1e-3 λ/D in visible





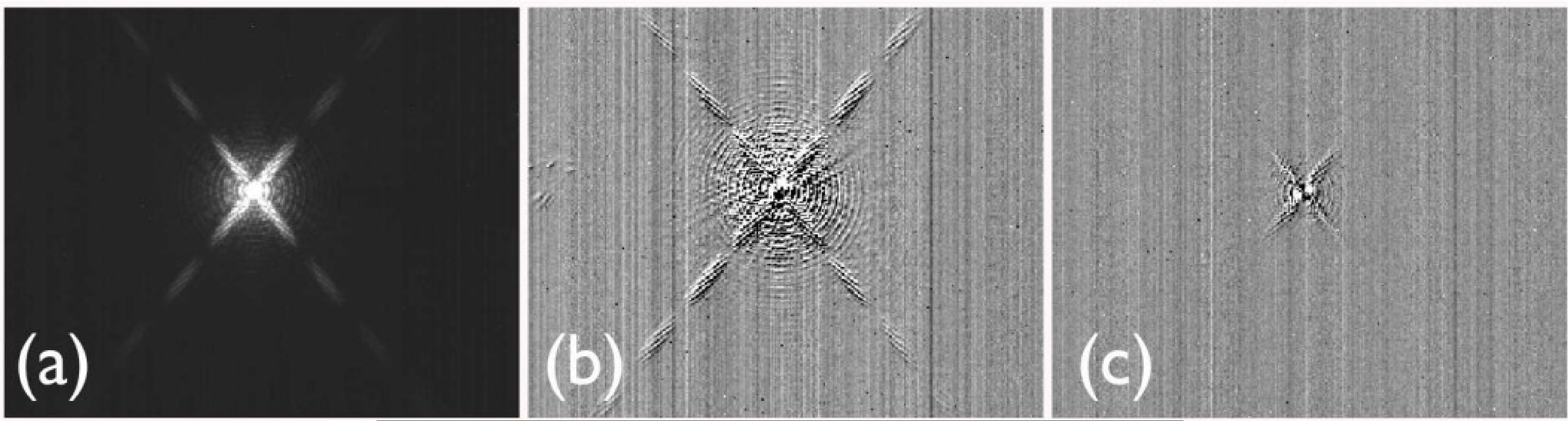
time (s)

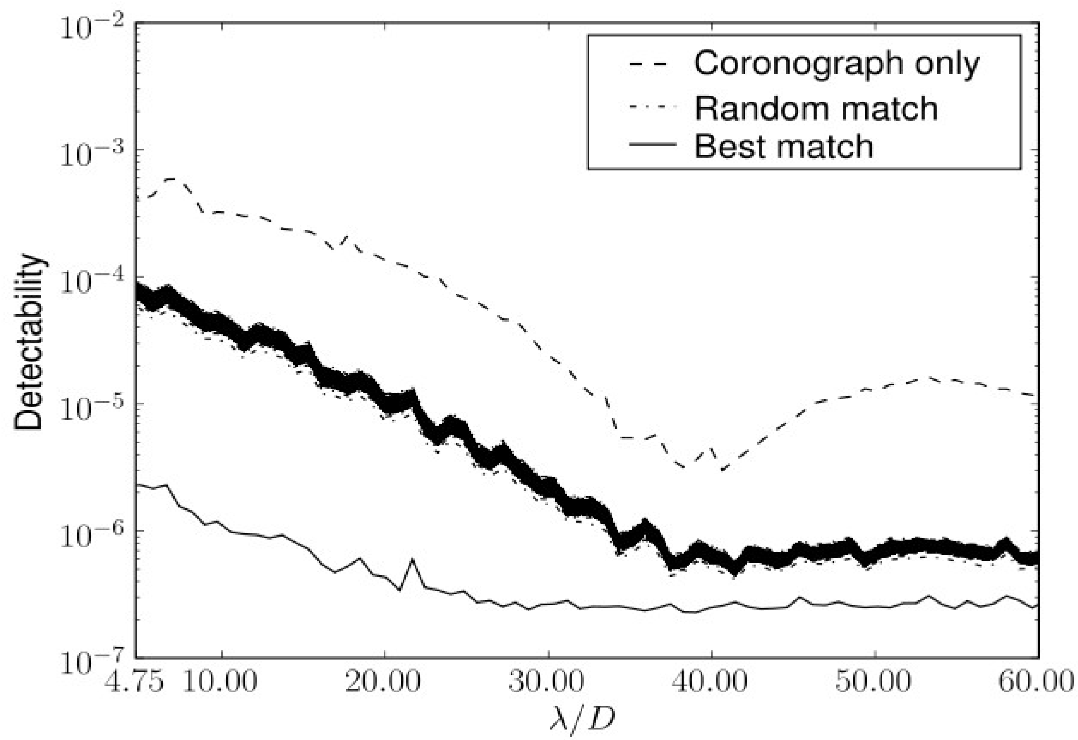




Coronagraph leaks calibrated to 1% in SCExAO (Vogt et al. 2011, submitted)

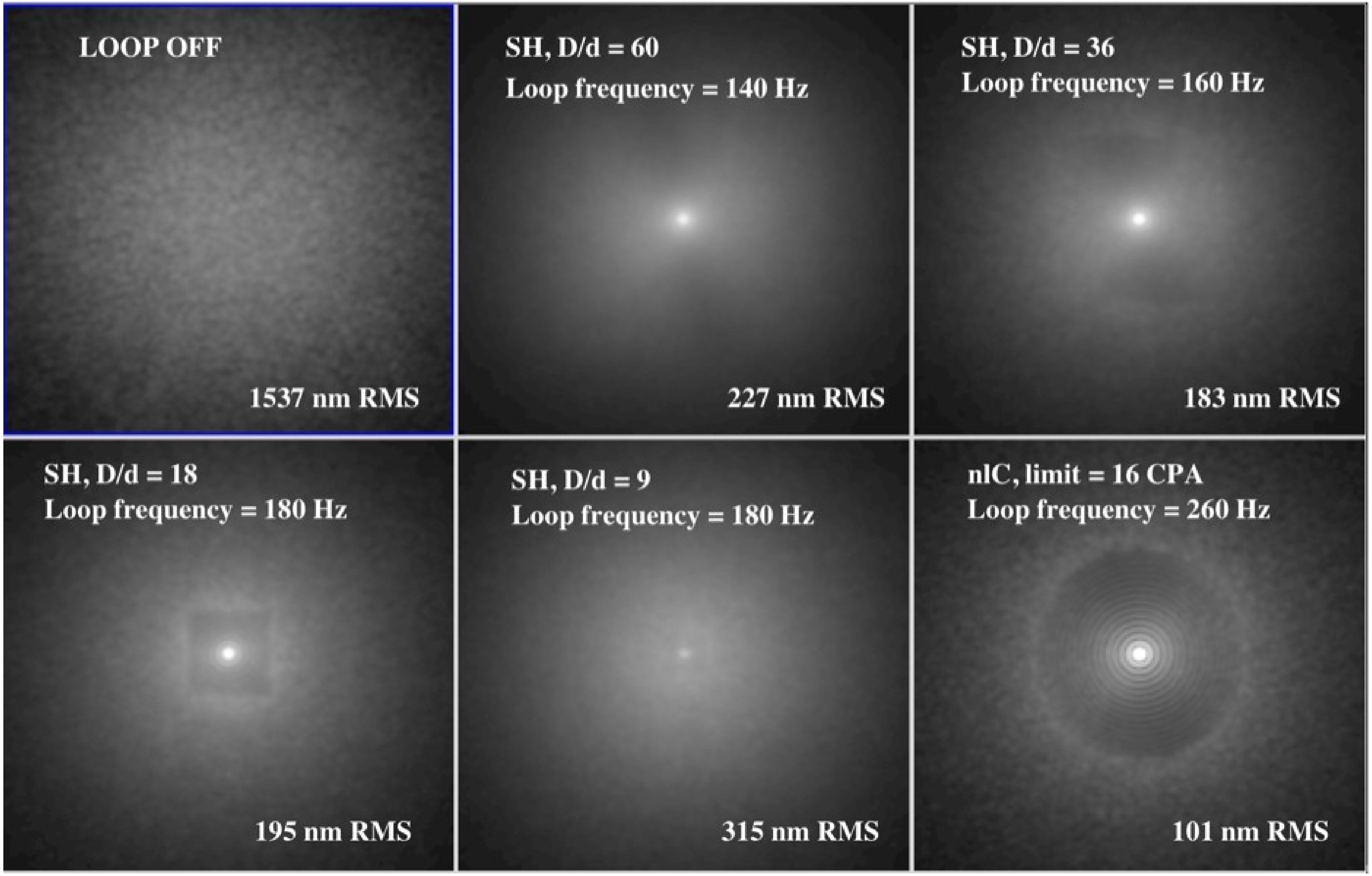
Co-added science image | Standard PSF subtraction | MMA





Computer Simulations showing contrast gain with high sensitivity WFS (non-linear curvature)

m ~ 13

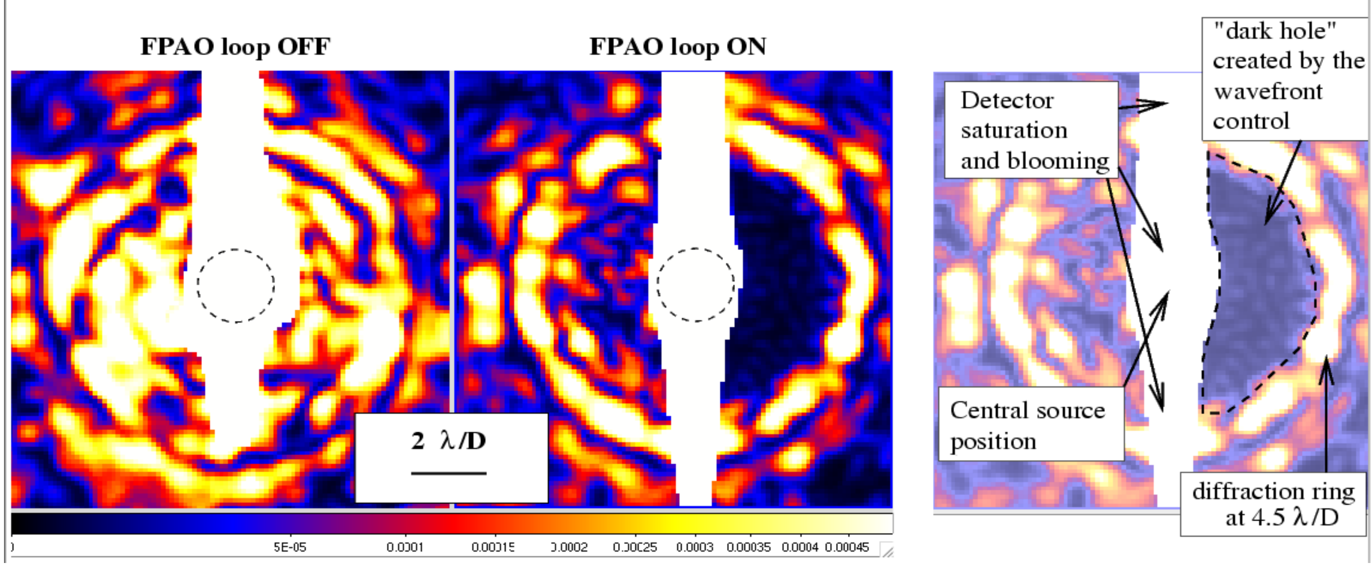


WFS	Loop frequ	RMS	SR @ 0.85 um	SR @ 1.6 um
nlCurv	260 Hz	101 nm	57%	85%
SH - D/9	180 Hz	315 nm	~4%	22%
SH - D/18	180 Hz	195 nm	~13%	56%
SH - D/36	160 Hz	183 nm	~16%	60%
SH - D/60	140 Hz	227 nm	~6%	45%

44

Focal plane AO and speckle calibration

Use Deformable Mirror (DM) to add speckles



SENSING: Put “test speckles” to measure speckles in the image, watch how they interfere

CORRECTION: Put “anti speckles” on top of “speckles” to have destructive interference between the two (Electric Field Conjugation, Give’on et al 2007)

CALIBRATION: If there is a real planet (and not a speckle) it will not interfere with the test speckles

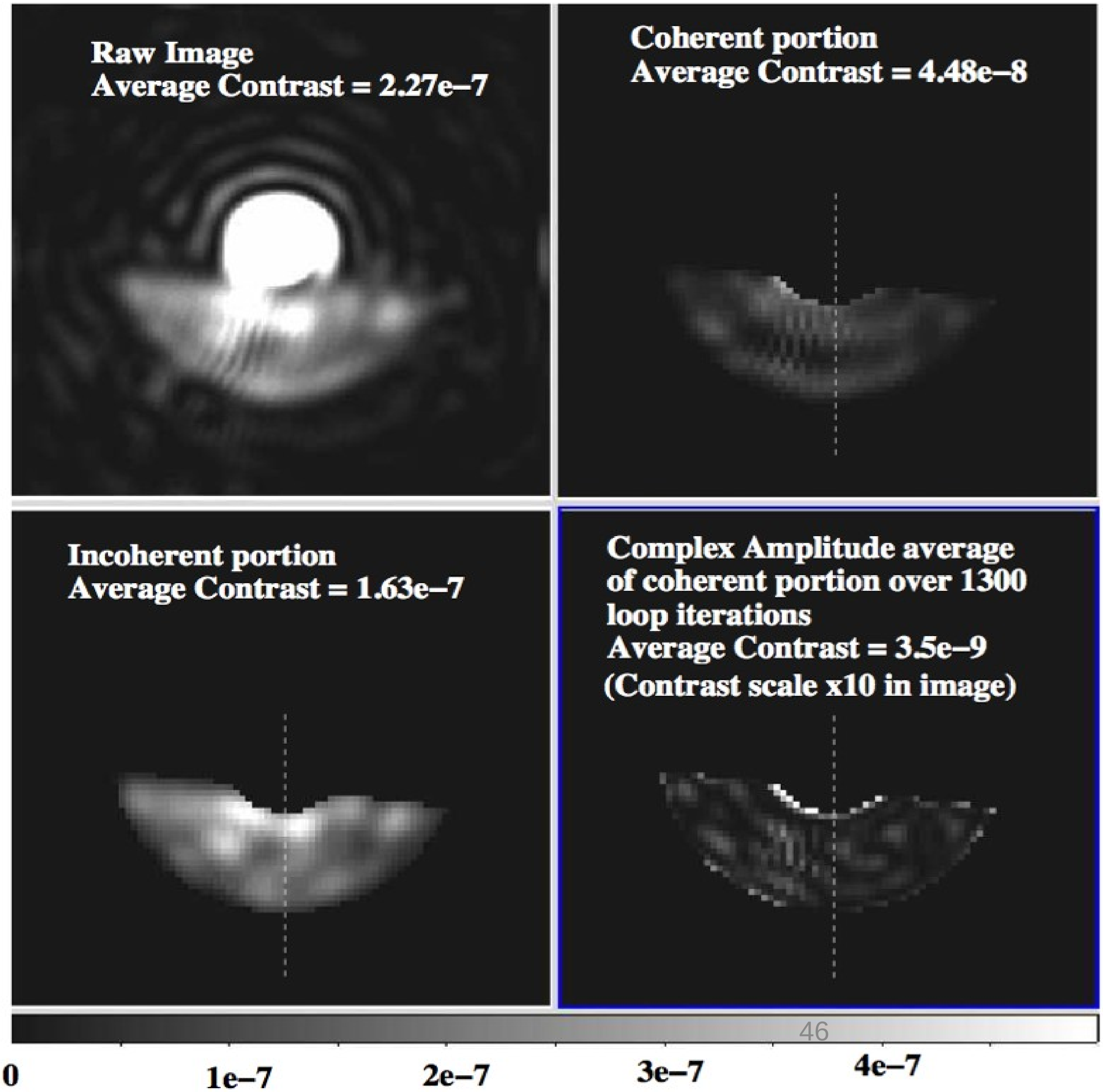
Fundamental advantage:
Uses science detector for wavefront sensing:
“What you see is EXACTLY what needs to be removed / calibrated”

Focal plane WFS based correction and speckle calibration: lab demo

2e-7 raw contrast obtained at 2 λ/D

Incoherent light at 1e-7
Coherent fast light at 5e-8
Coherent bias <3.5e-9

Test demonstrates:
- ability to separate light into coherent/incoherent fast/slow components
- ability to slow and static remove speckles well below the dynamic speckle halo



Reflected light imaging

Initially very hard – requires ~1e-7 contrast and small inner working angle

Contrast is not a steep function of planet mass → small gains in contrast performance can lead to large science gain

Easiest targets are the closest stars (few pc) – stars can be old

High performance system on 8-m telescope can detect reflected light from a few of the known RV planets
This will be an important science case for ELTs

Upcoming ExAO systems will validate technologies for reflected light imaging on larger telescopes