

# Planet Occurrence Within 0.25 AU

Andrew W. Howard  
UC Berkeley

## California Planet Search Team:

Geoff Marcy, Debra Fischer, John Johnson, Jason Wright, Howard Isaacson, Julien Spronck, Jeff Valenti, Jay Anderson, Nikolai Piskunov, more!



Photo credit: Laurie Hatch

# Acknowledgements

## NASA: Kepler

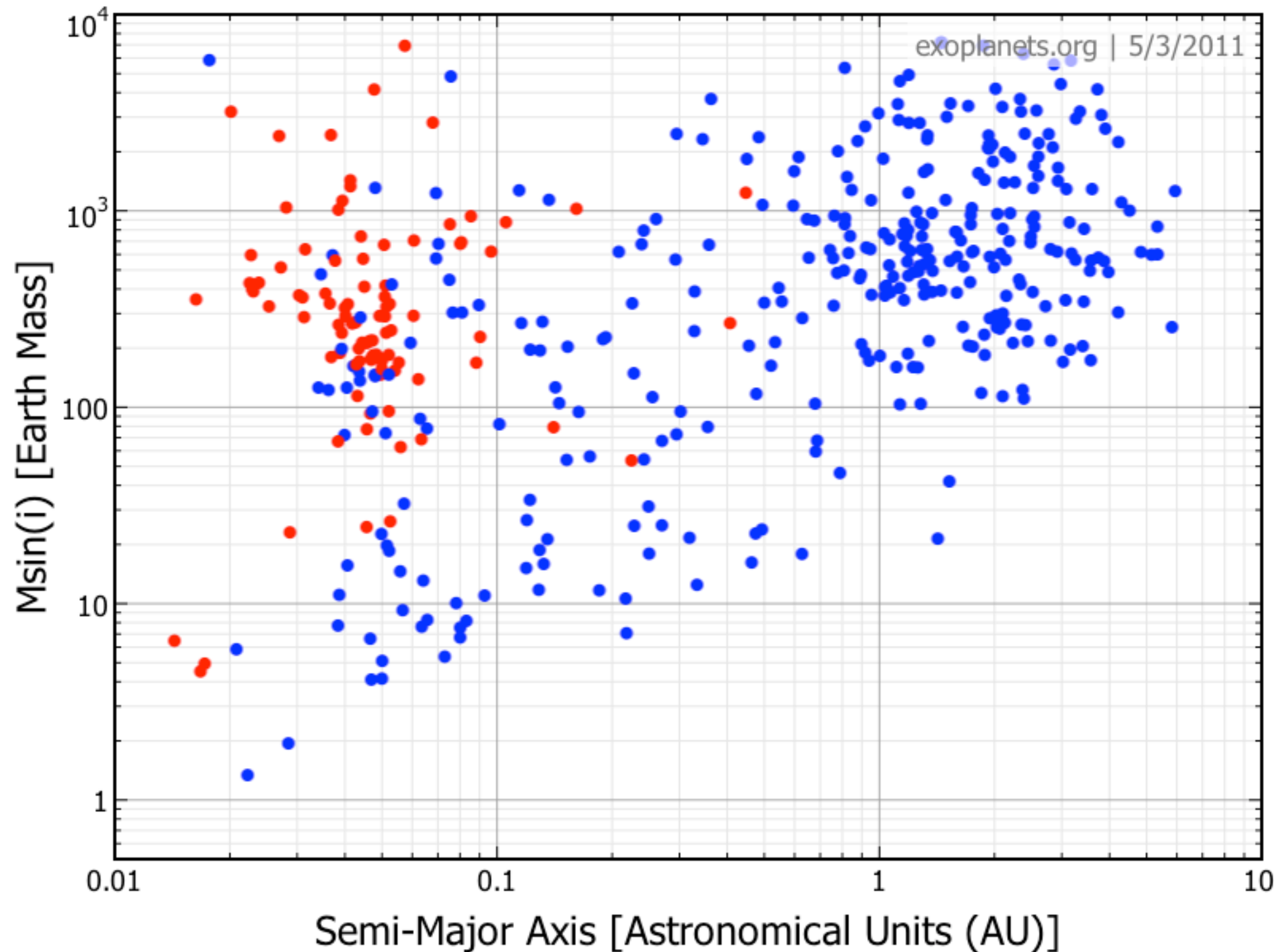
PI: Bill Borucki & Deputy PI: David Koch

## Ball Aerospace

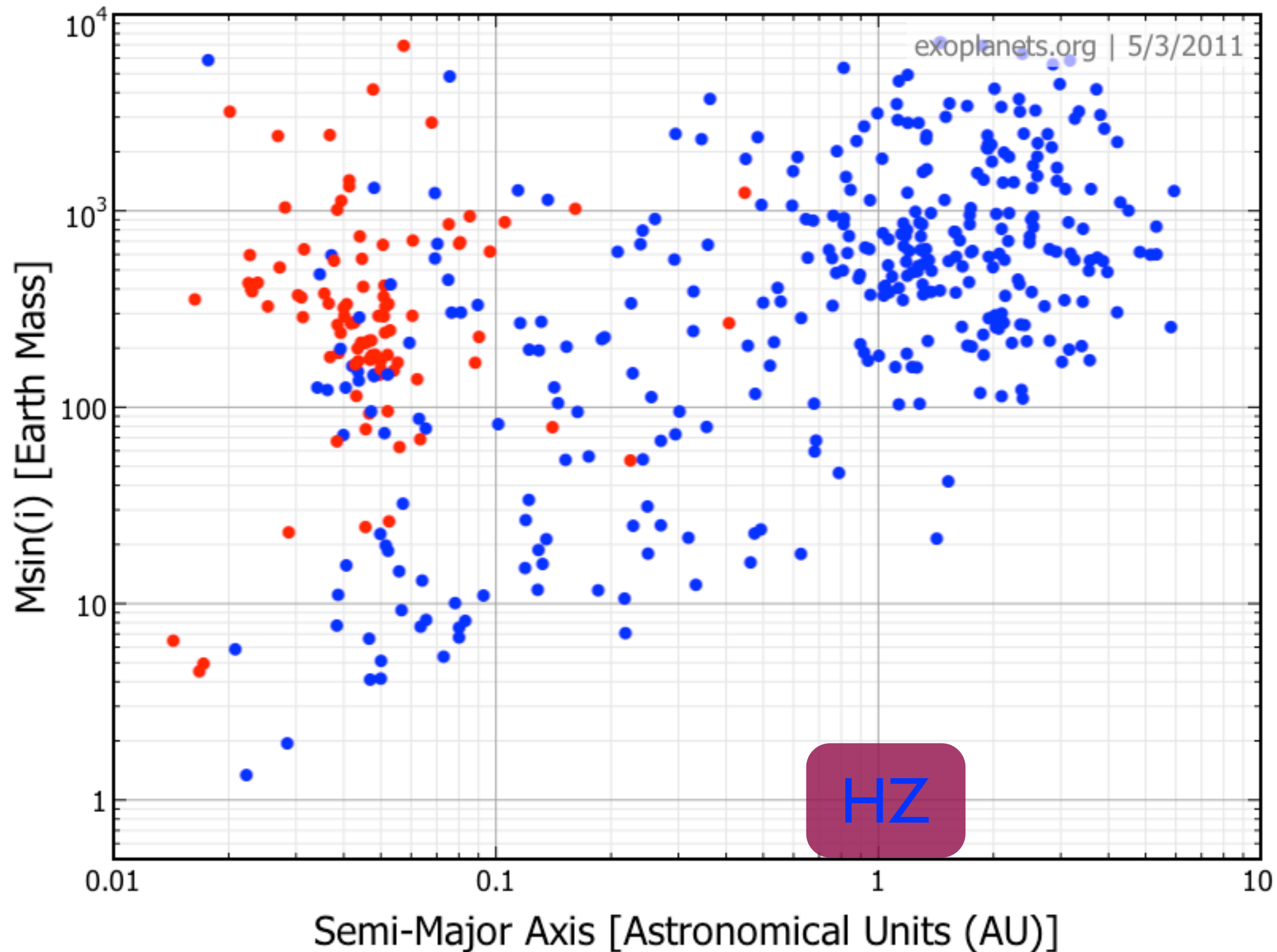
## The Kepler Team:

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Ronald L. Gilliland<sup>13</sup>, Steve B. Howell<sup>14</sup>, Howard Isaacson<sup>3</sup>, John Asher Johnson<sup>15</sup>,  
David W. Latham<sup>7</sup>, Jack J. Lissauer<sup>2</sup>, David G. Monet<sup>16</sup>, Jason F. Rowe<sup>2,17</sup>,  
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Bruce D. Clarke<sup>1</sup>, Jie Li<sup>1</sup>, Christopher Allen<sup>18</sup>, Peter Tenenbaum<sup>1</sup>, Hayley Wu<sup>1</sup>, Søren Meibom<sup>7</sup>,  
Todd C. Klaus<sup>18</sup>, Christopher K. Middelour<sup>18</sup>, Miles T. Cote<sup>2</sup>, Sean McCauliff<sup>18</sup>,  
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David L. Pletcher<sup>2</sup>, Jessie L. Dotson<sup>2</sup>, Michael R. Haas<sup>2</sup>

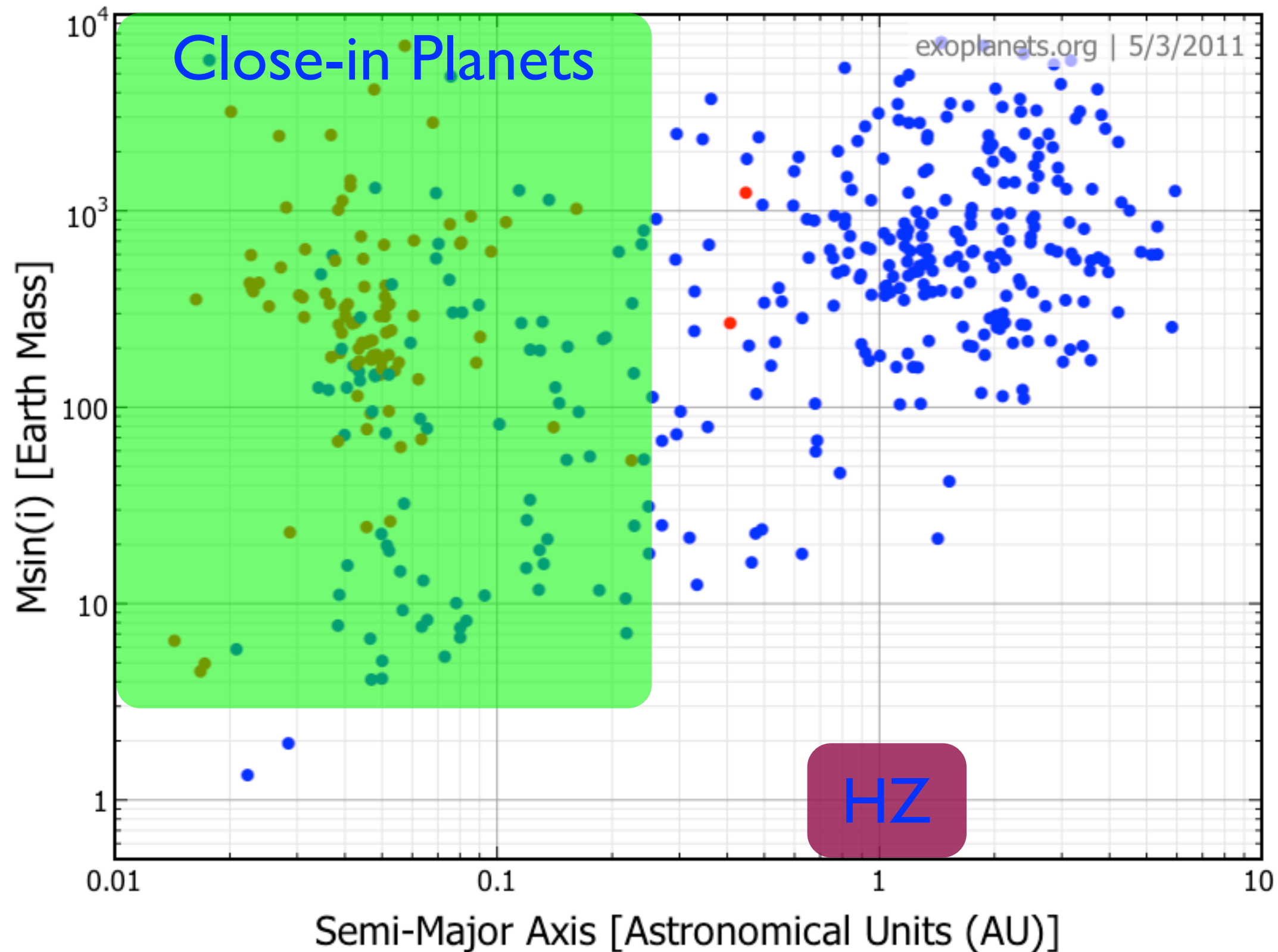
# Well-Characterized Planets



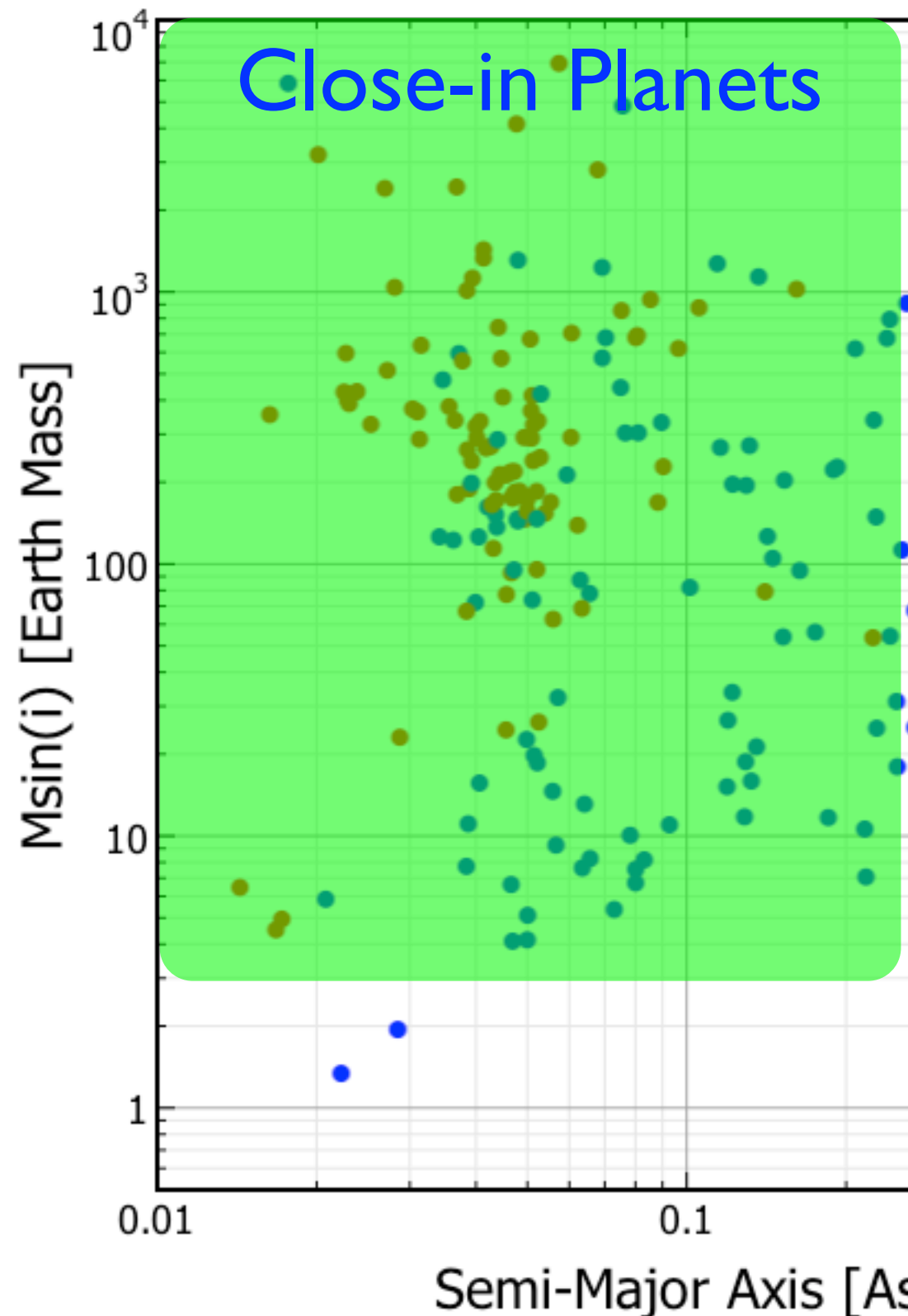
# Well-Characterized Planets



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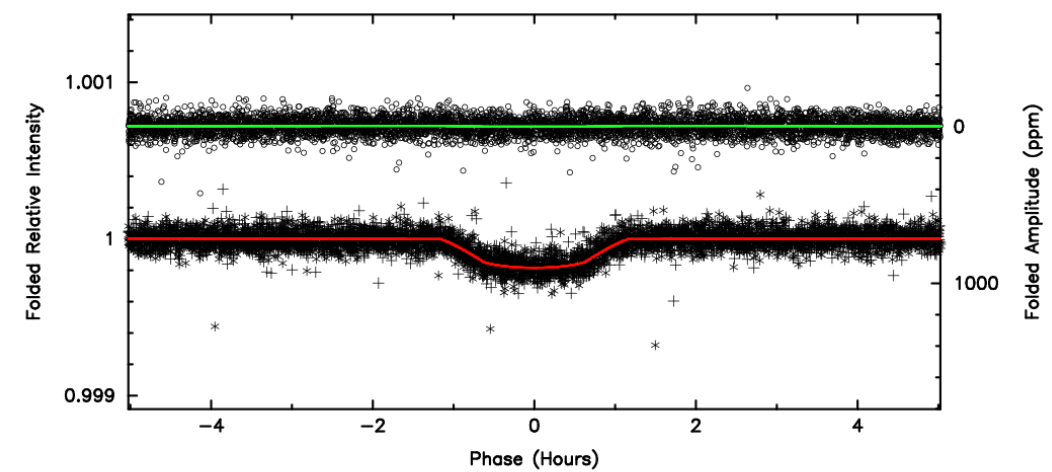
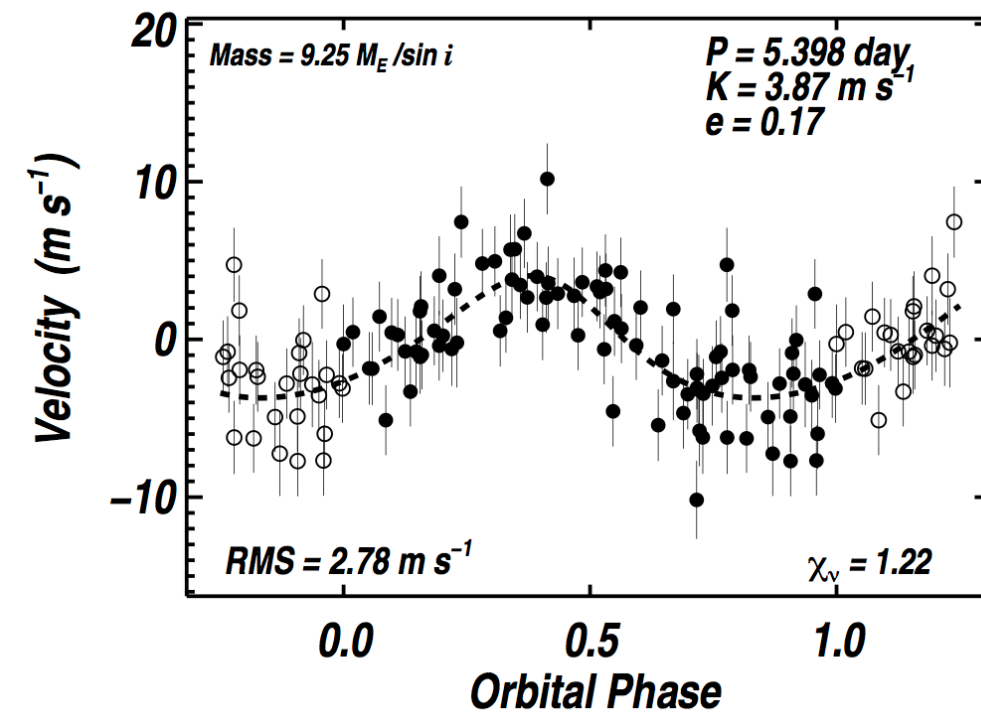
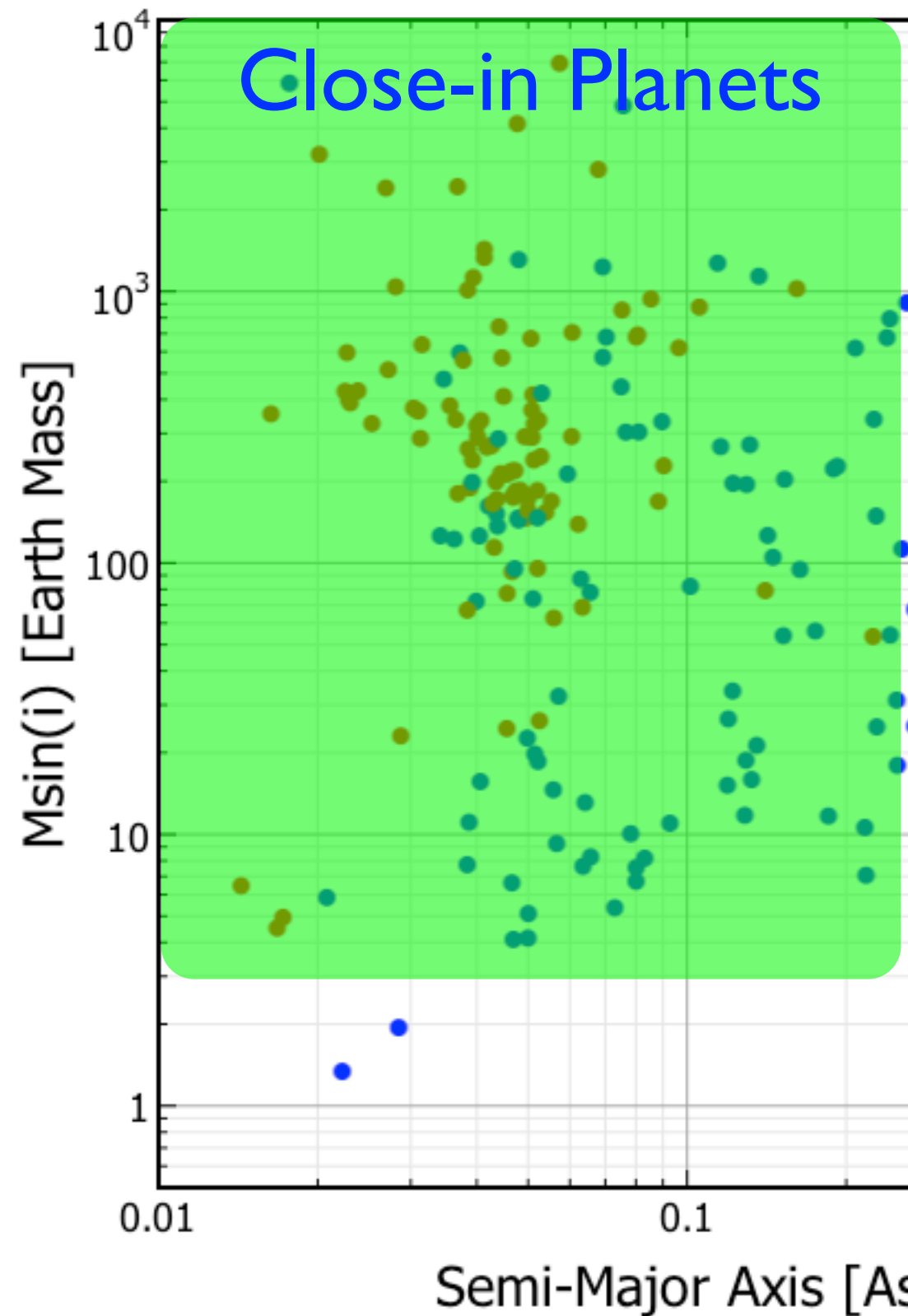


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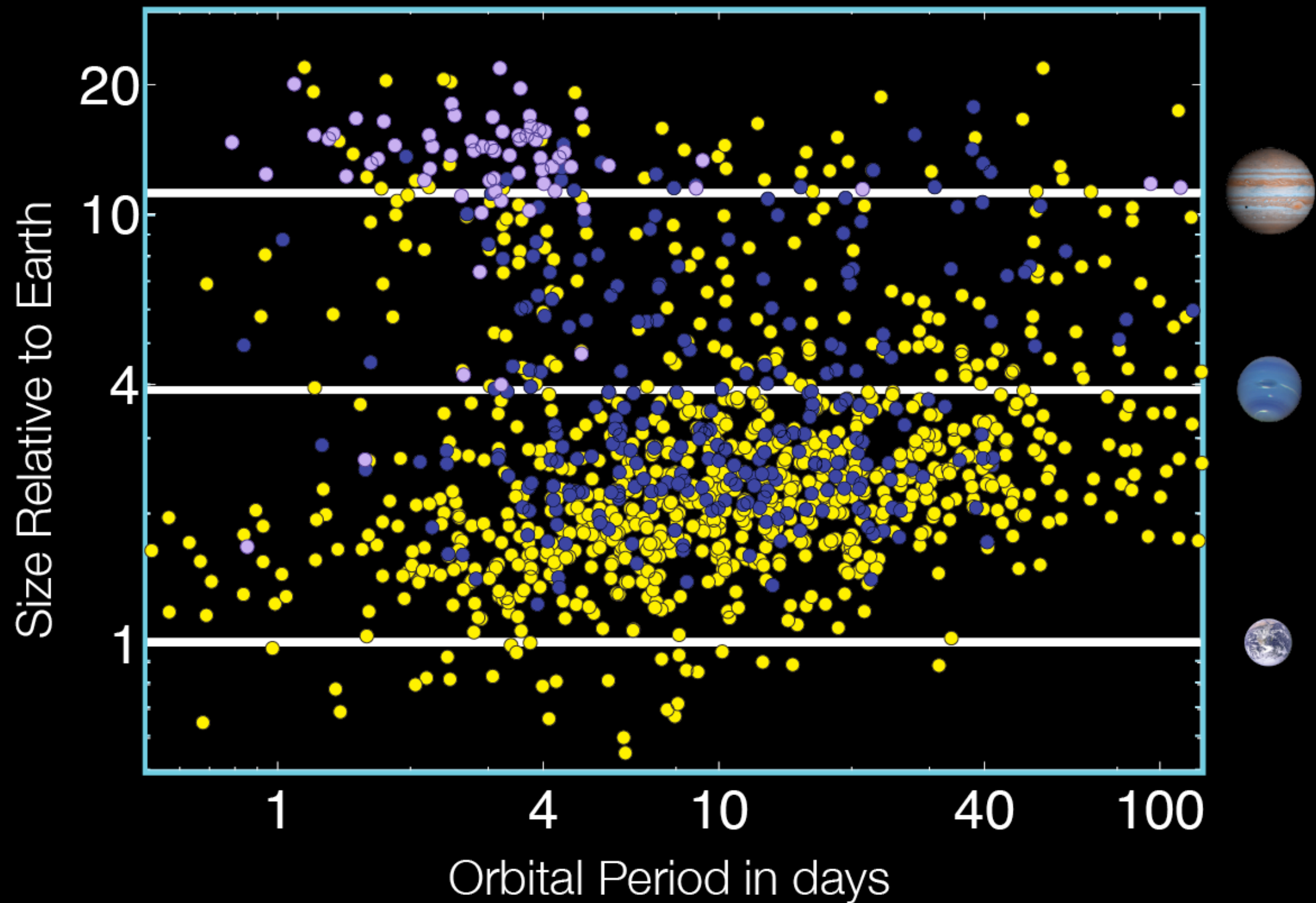


- Measure planet mass & radius functions
- Constrain planet densities
- Probe planet formation and migration
- Test population synthesis models
- Guide future searches

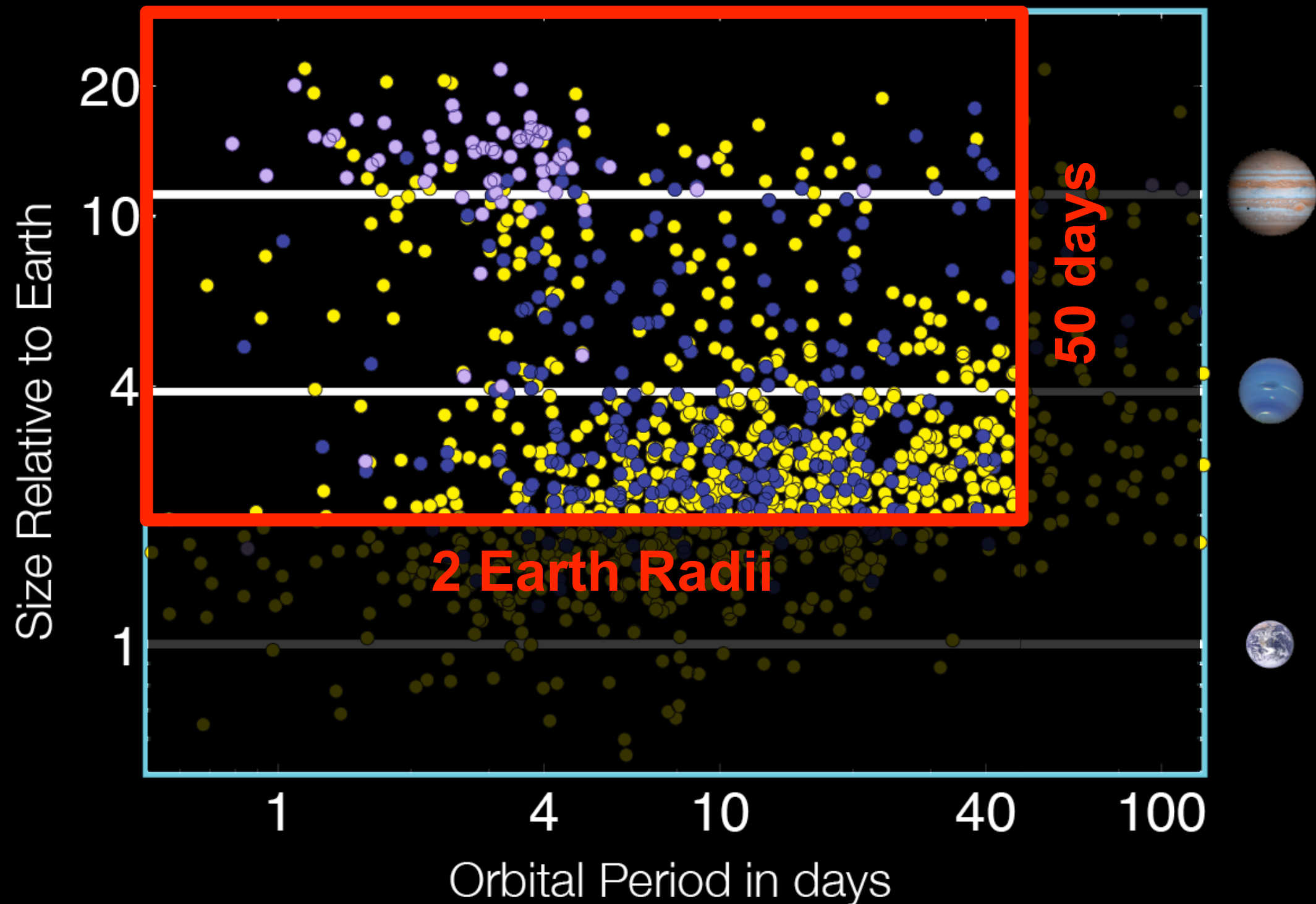
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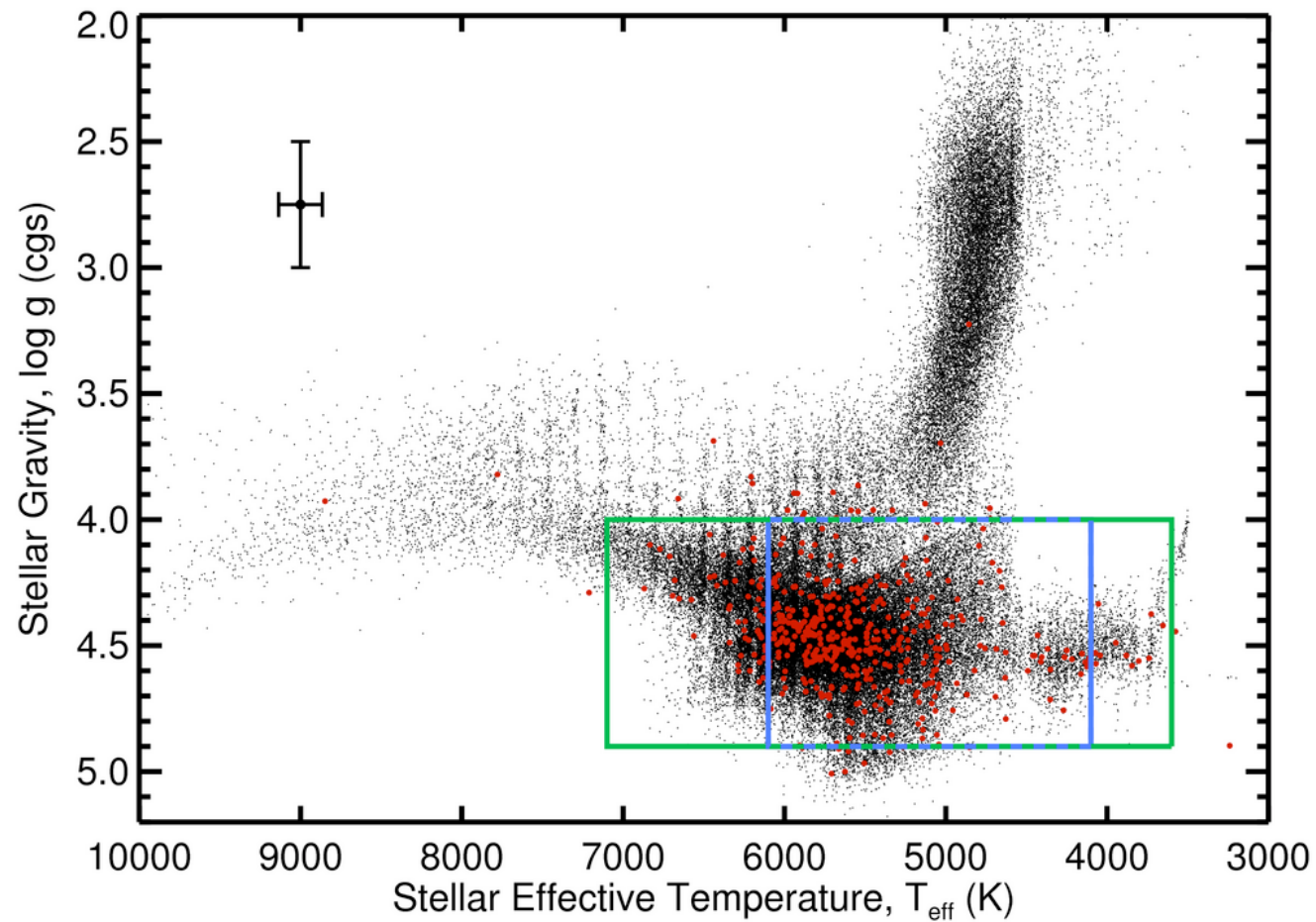
# Kepler Candidates as of February 1, 2011



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# Choose Subset of Planets/Stars for High Detectability



- Only GK dwarfs ( $T_{\text{eff}}/\log g$ )
- Only bright stars ( $Kp < 15$ )
- Only high SNR transits (SNR  $> 10$  in 90 days)

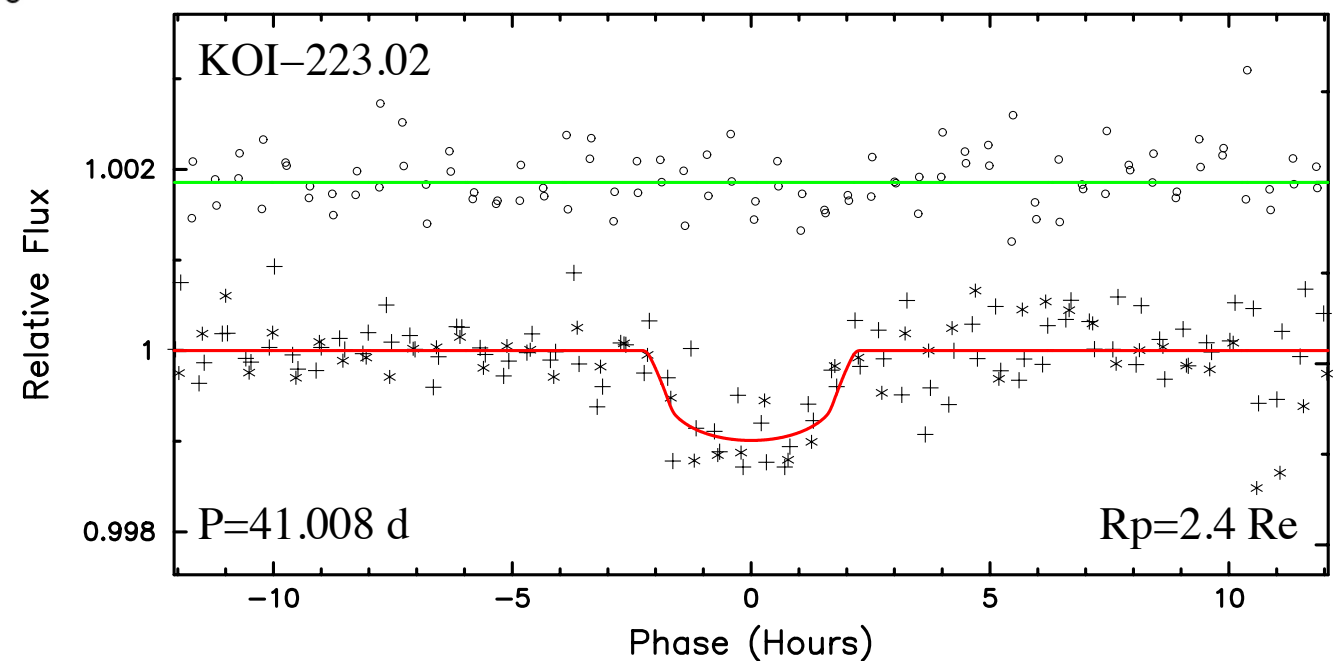
## Full sample

~156,000 stars  
1,230 planets

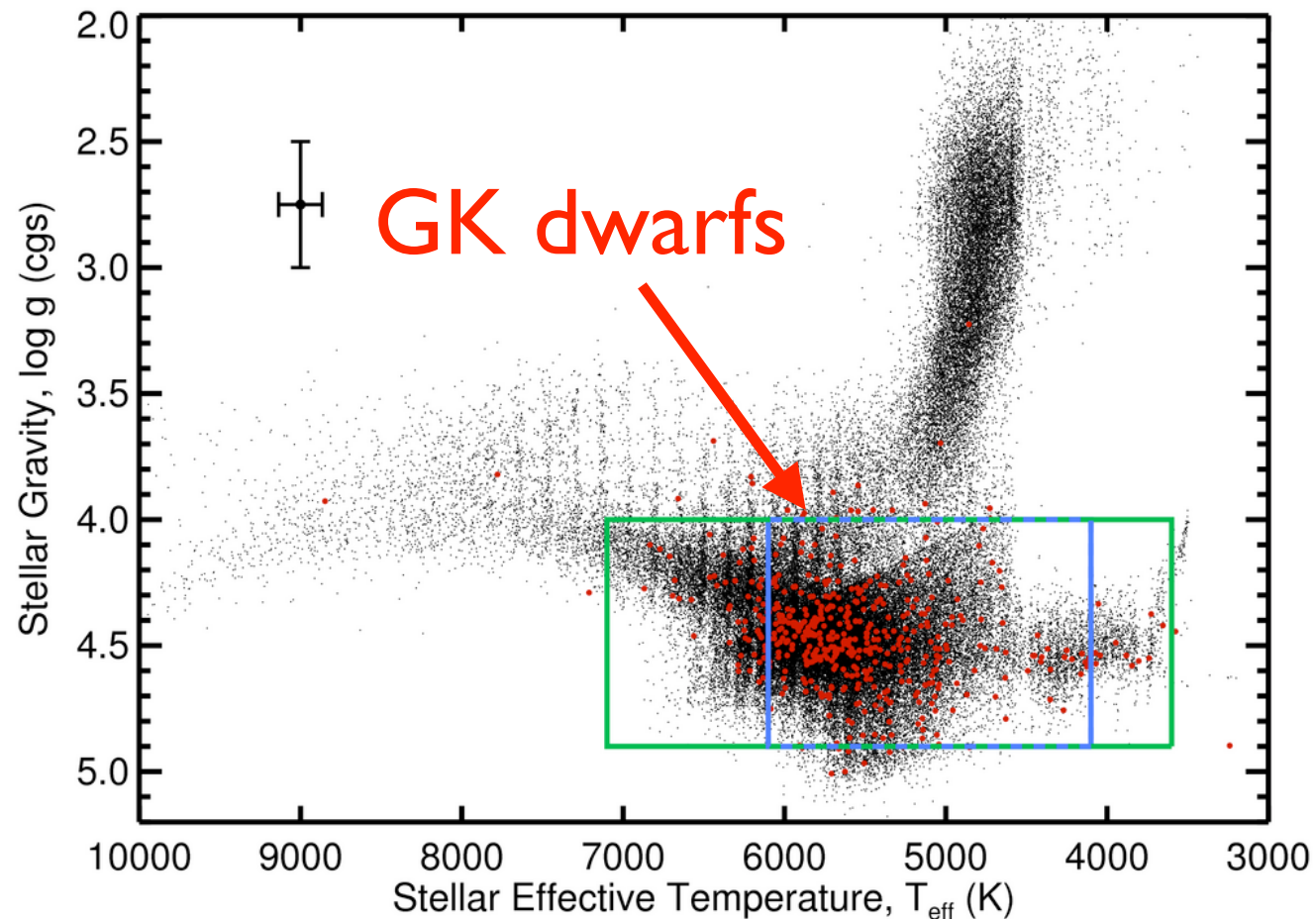


## Restricted sample

~58,000 stars  
438 planets



# Choose Subset of Planets/Stars for High Detectability



- Only GK dwarfs ( $T_{\text{eff}}/\log g$ )
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- Only high SNR transits (SNR > 10 in 90 days)

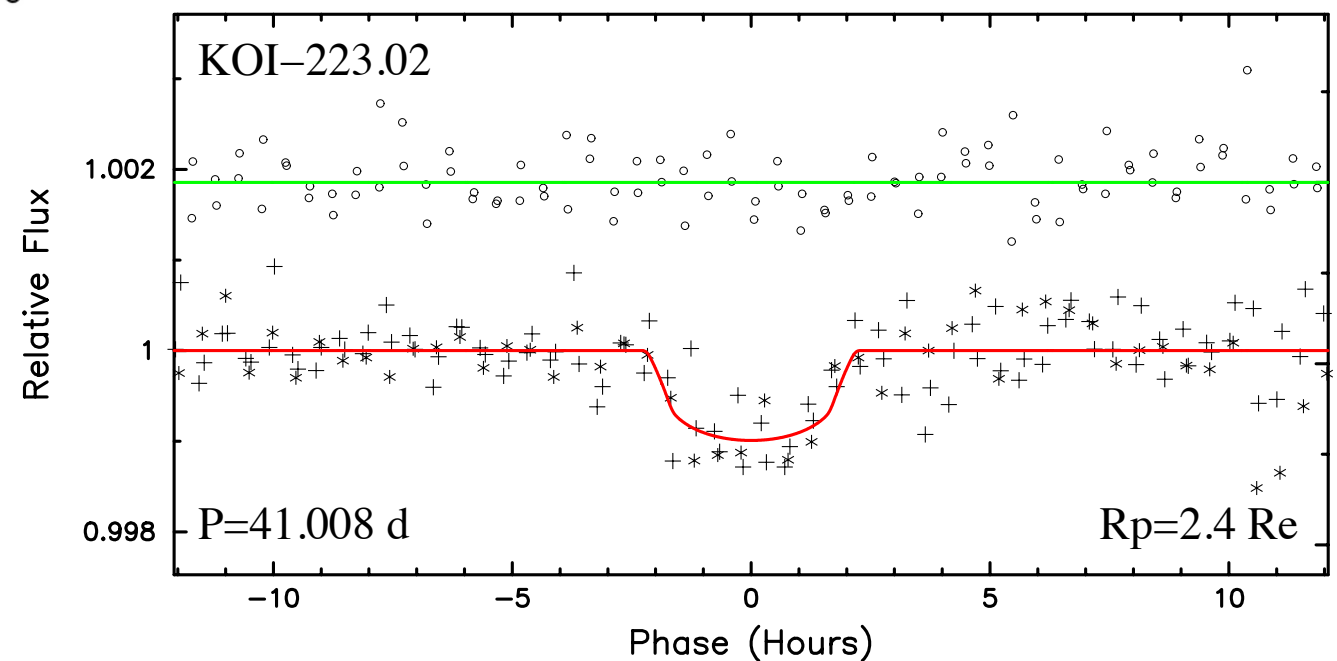
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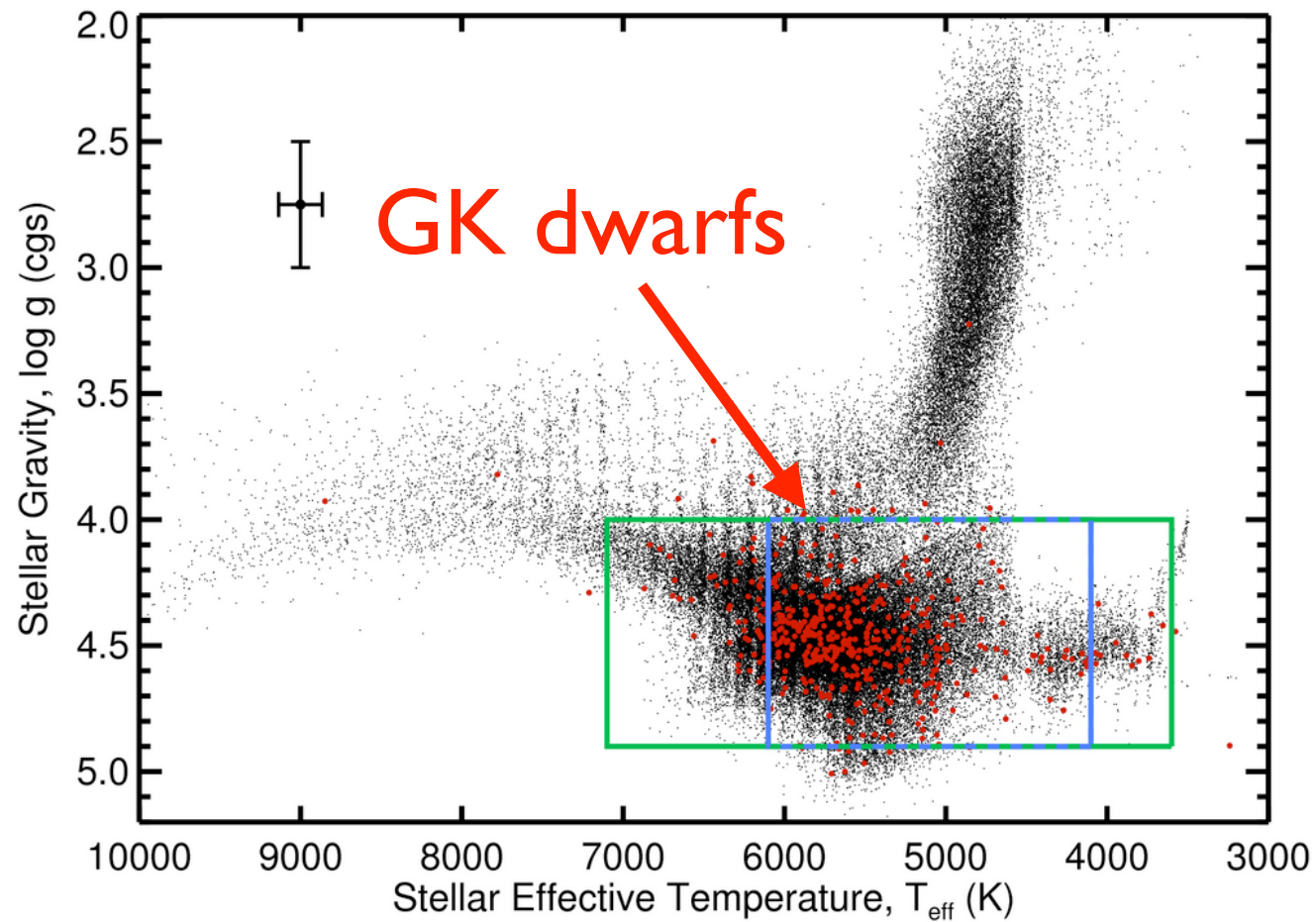


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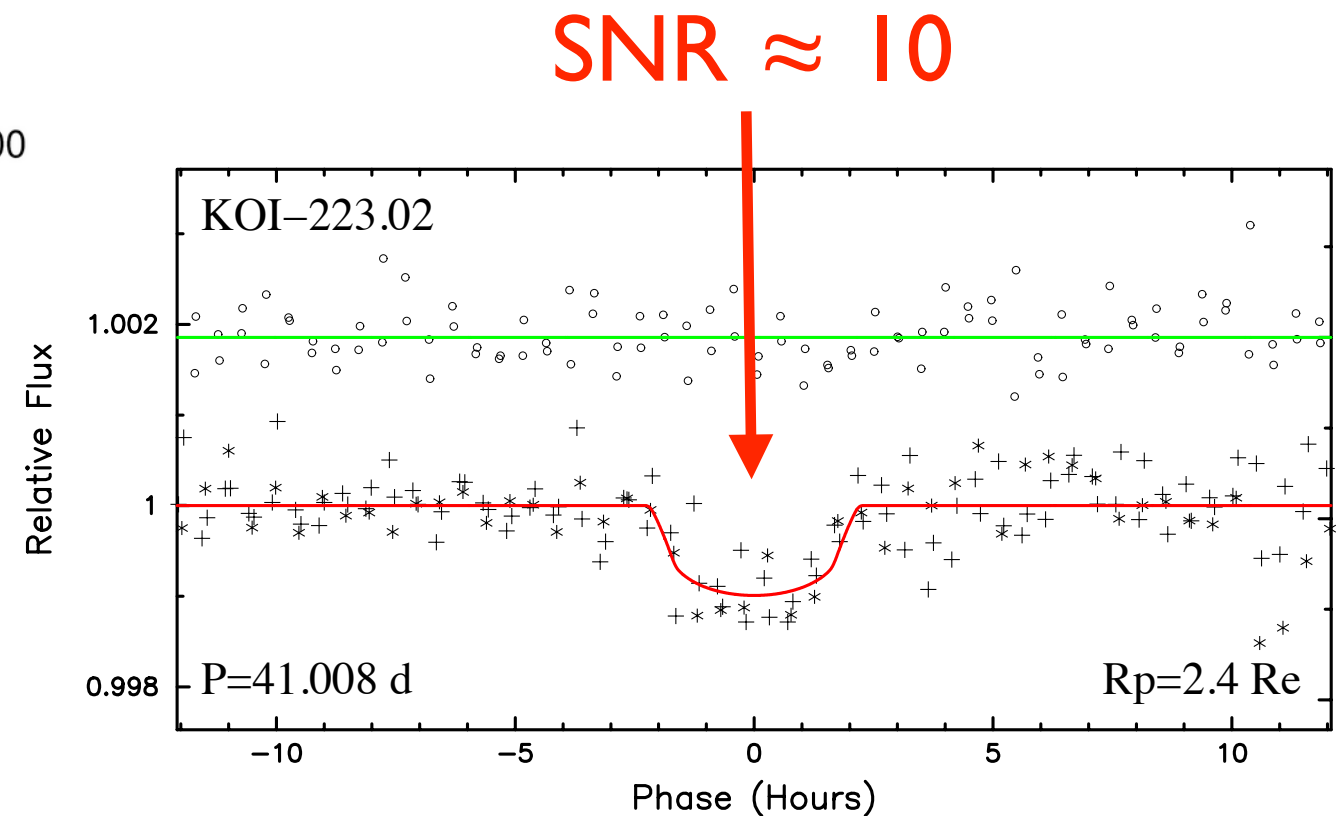
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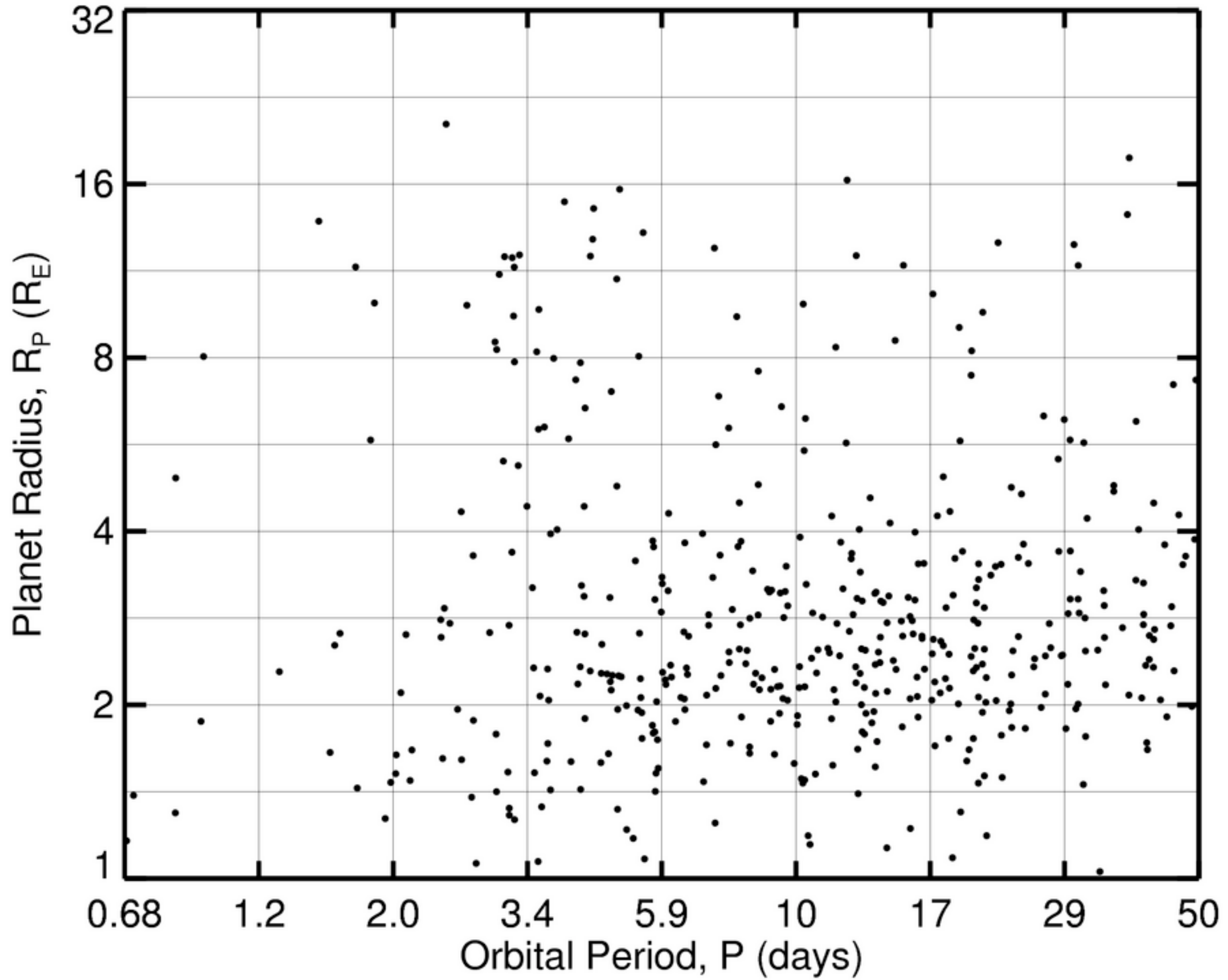
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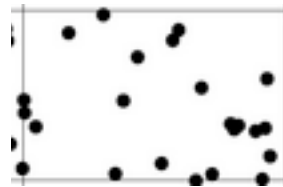


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# Compute Planet Occurrence in Small Domain

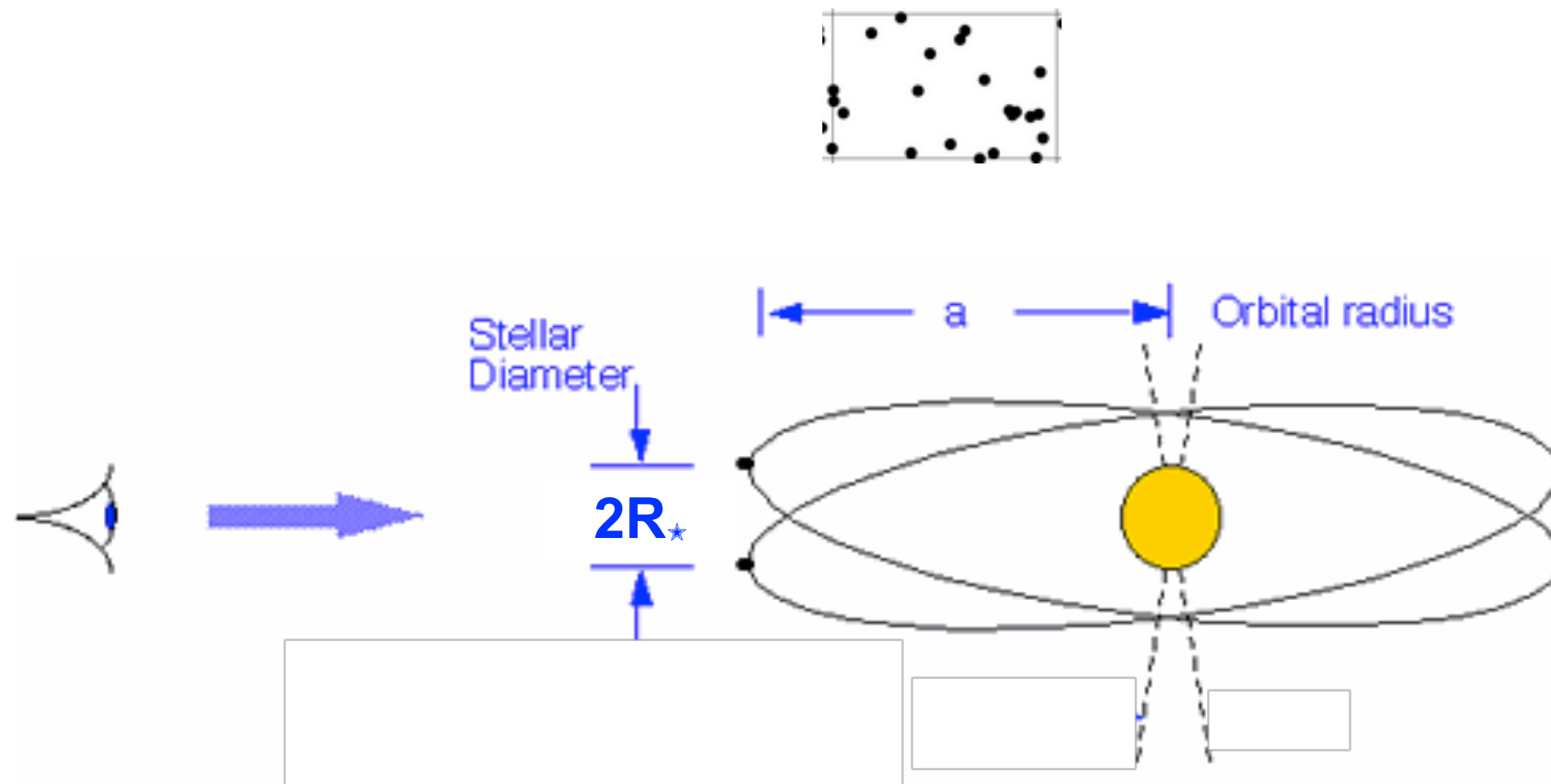


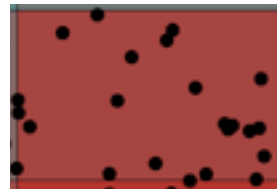
# Compute Planet Occurrence in Small Domain

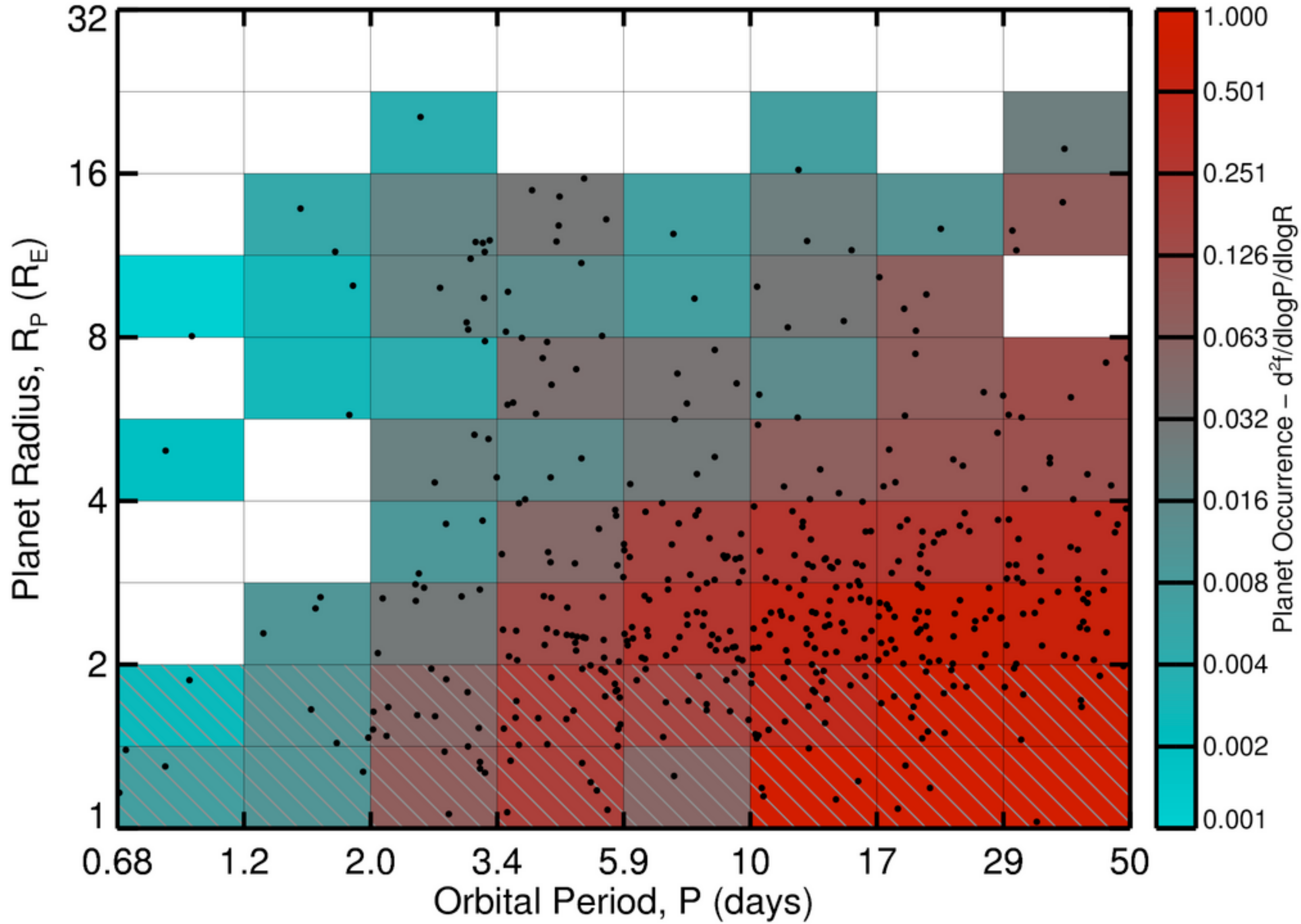
Correct for:

Transit probabilities -  $p_{\text{transit}} = R_{\star}/a$

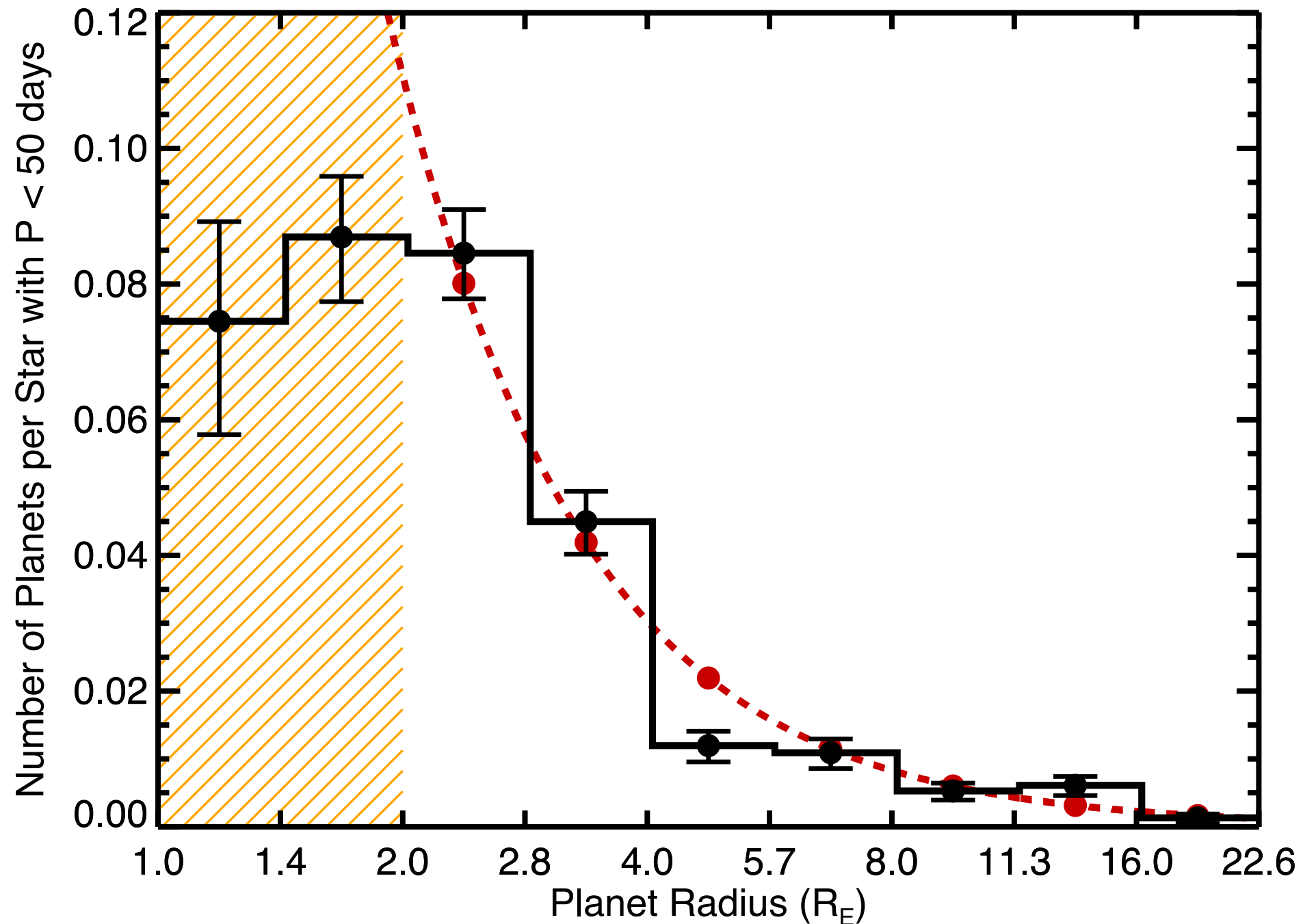
Incompleteness - some stars lack the  
photometric sensitivity







# Planet Radius Distribution

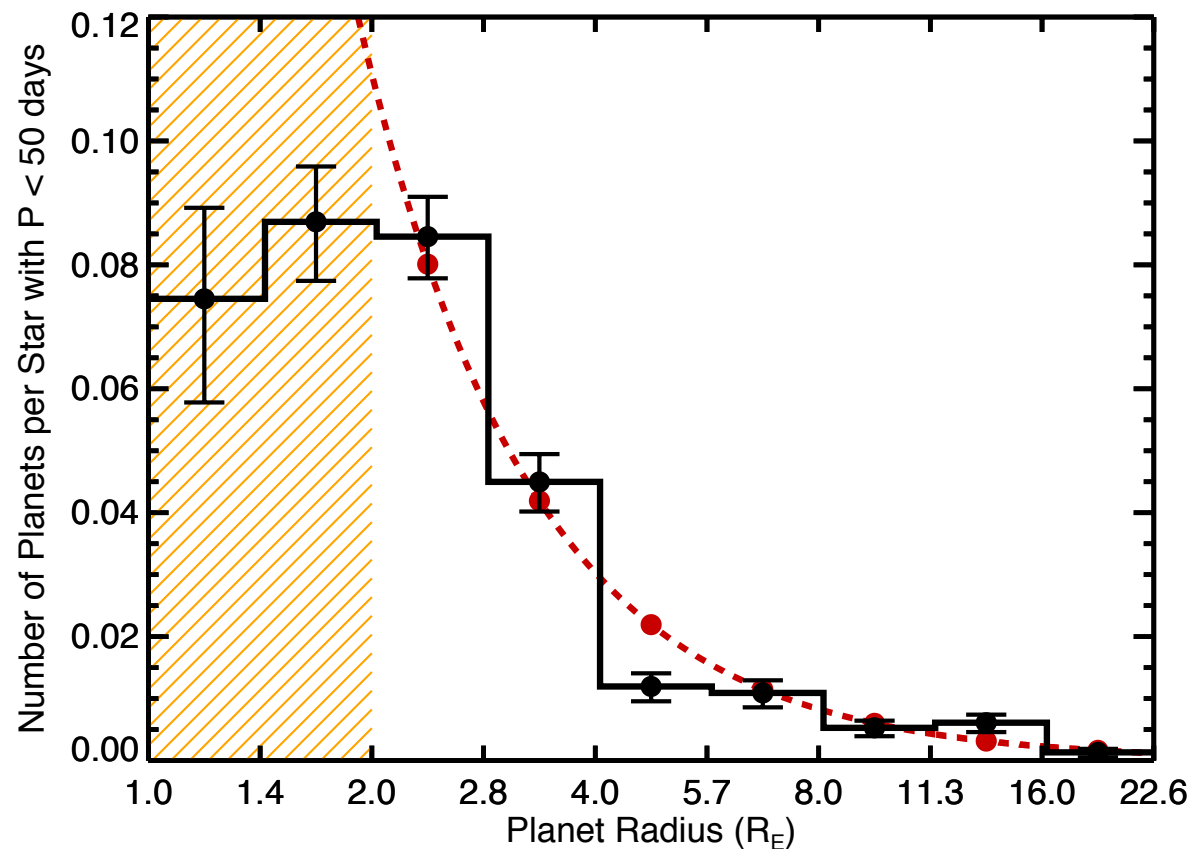


Power law:  
 $dN/d\log R = kR^\alpha$

$$k = 2.9 \pm 0.5$$
$$\alpha = -1.92 \pm 0.11$$

# Planet Radius Distribution *Kepler*

Howard et al. 2011c, ApJ submitted, arXiv:1103.2541



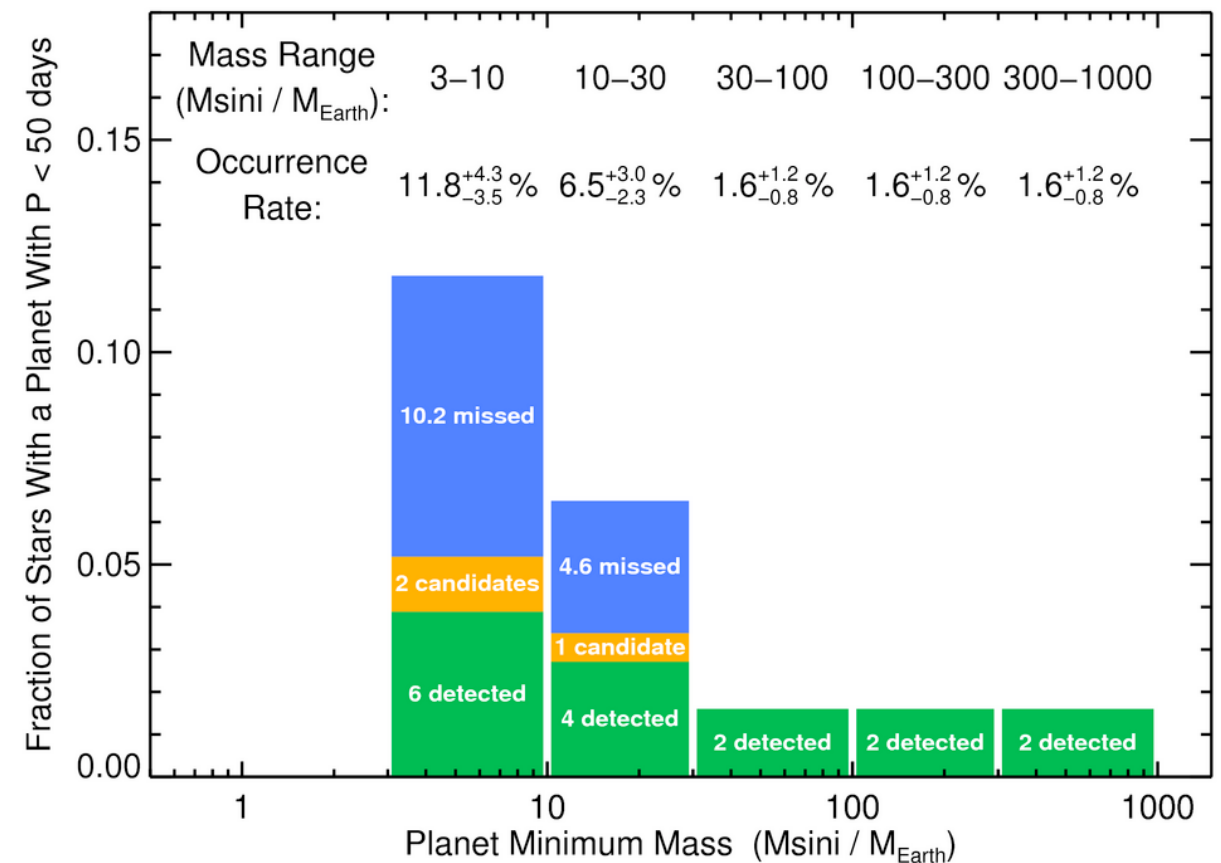
## Power Law Radius Function

$$df/d\log R = kR^\alpha$$

$$k = 2.9 \pm 0.5, \alpha = -1.92 \pm 0.5$$

# Planet Mass Distribution Eta-Earth Survey (*Doppler*)

Howard et al. 2010, Science, 330, 653



## Power Law Mass Function

$$df/d\log M = kM^\alpha$$

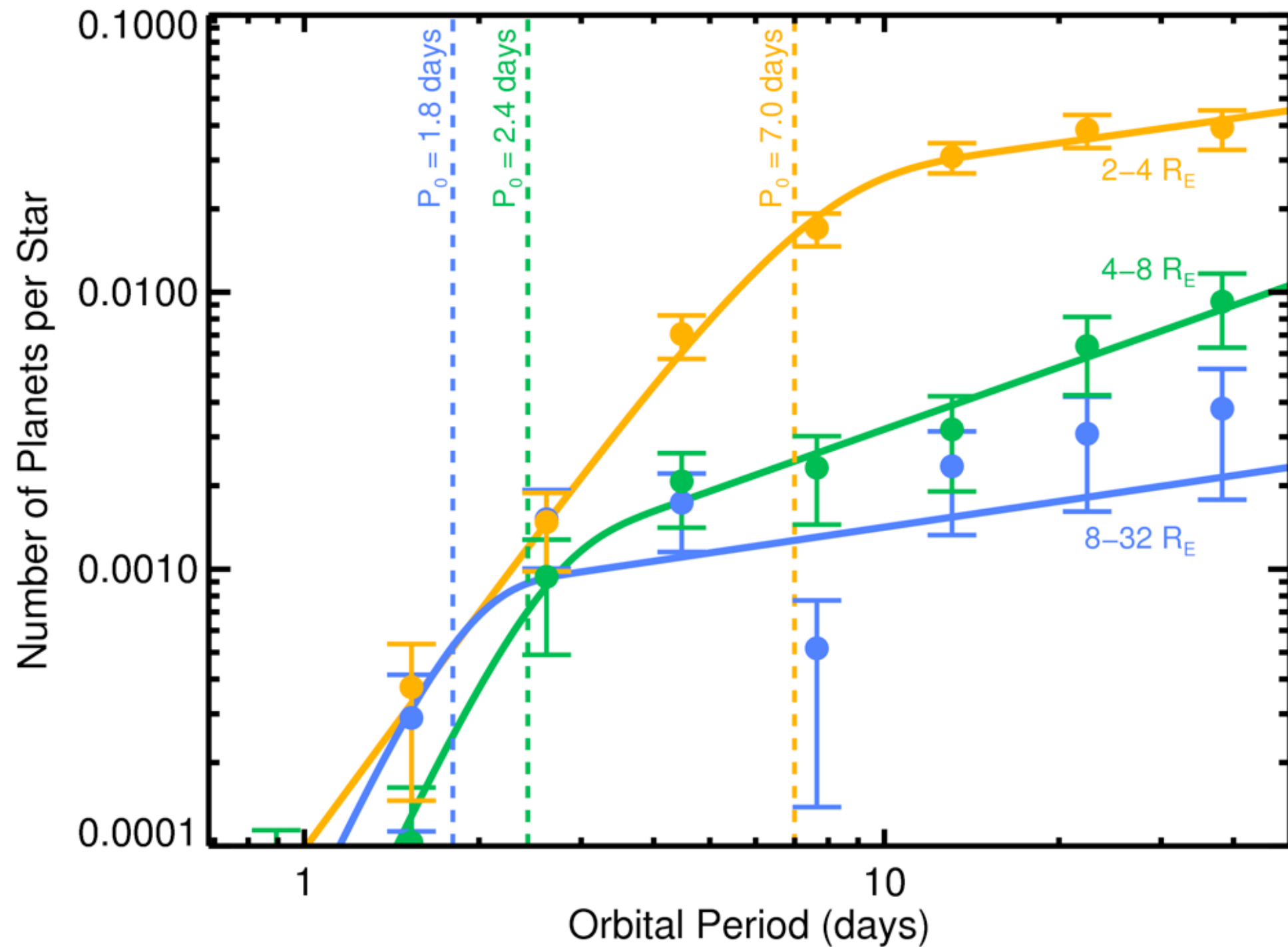
$$k = 0.39^{+0.27}_{-0.16}, \alpha = -0.48^{+0.12}_{-0.14}$$

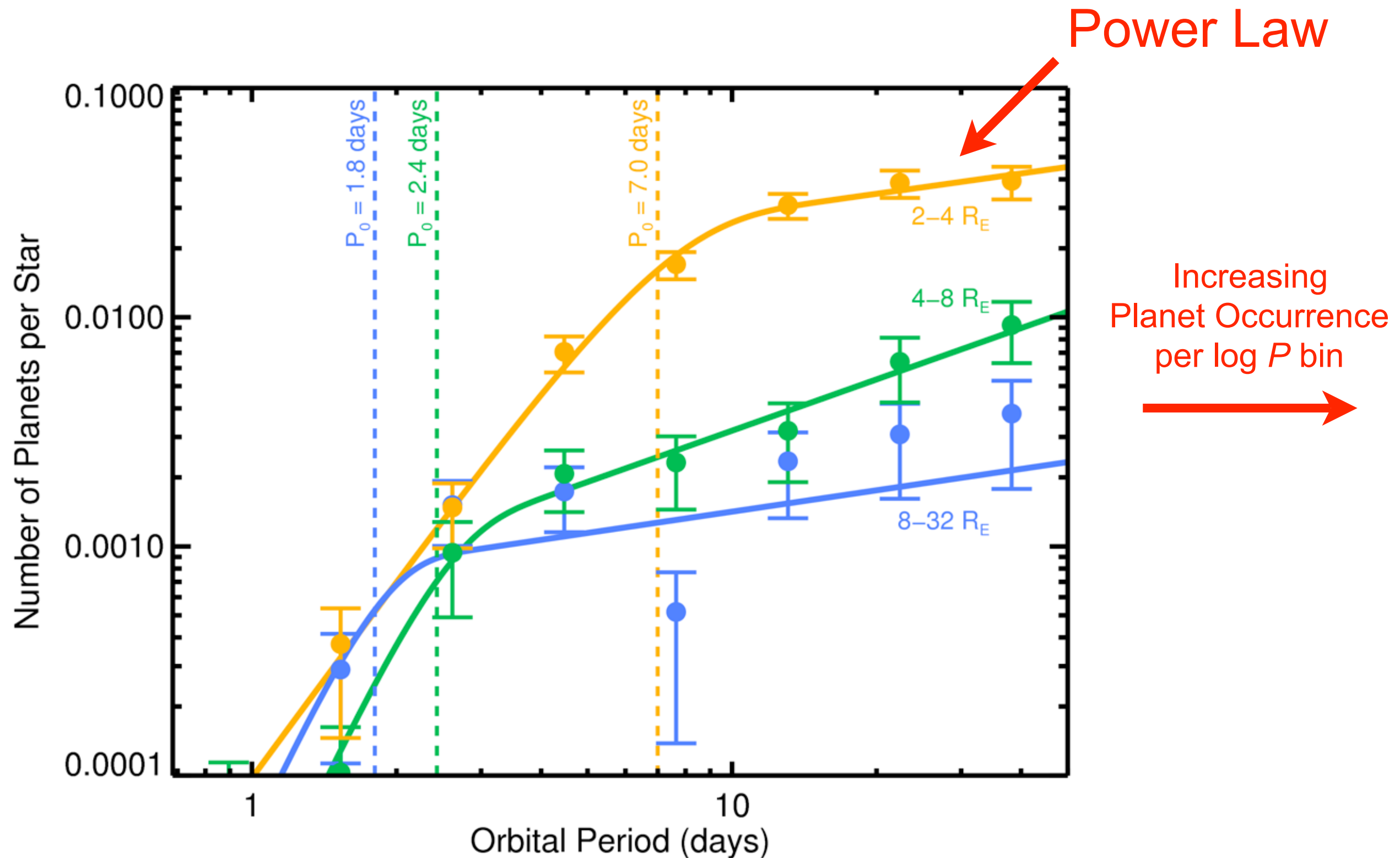
## Extrapolate:

$$\eta_{\text{Earth}} = 23^{+16}_{-10} \%$$

$$M_{\text{sini}} = 0.5-2.0 M_E, P < 50 \text{ days}$$

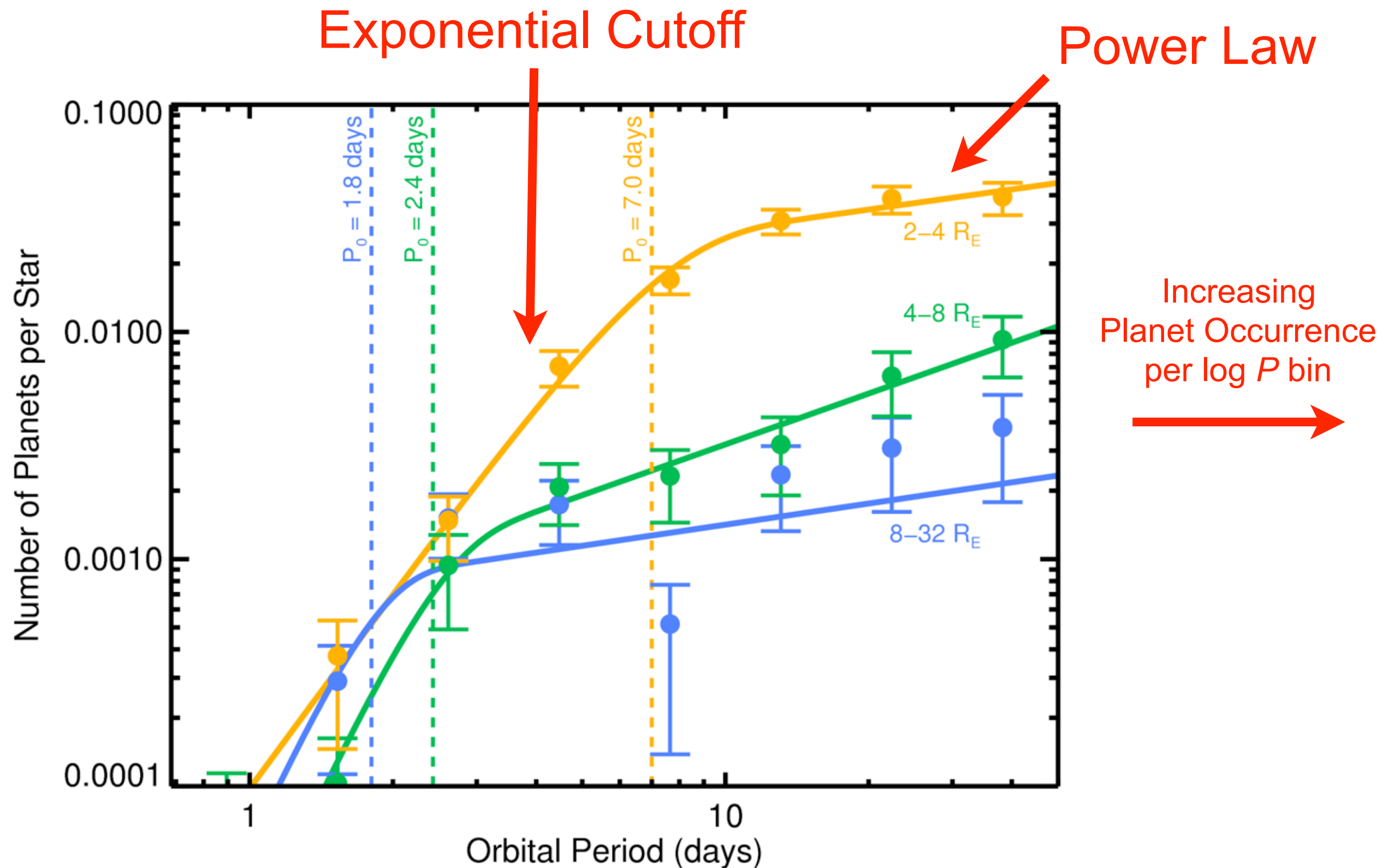
# Back to *Kepler*: Planet Period Distribution



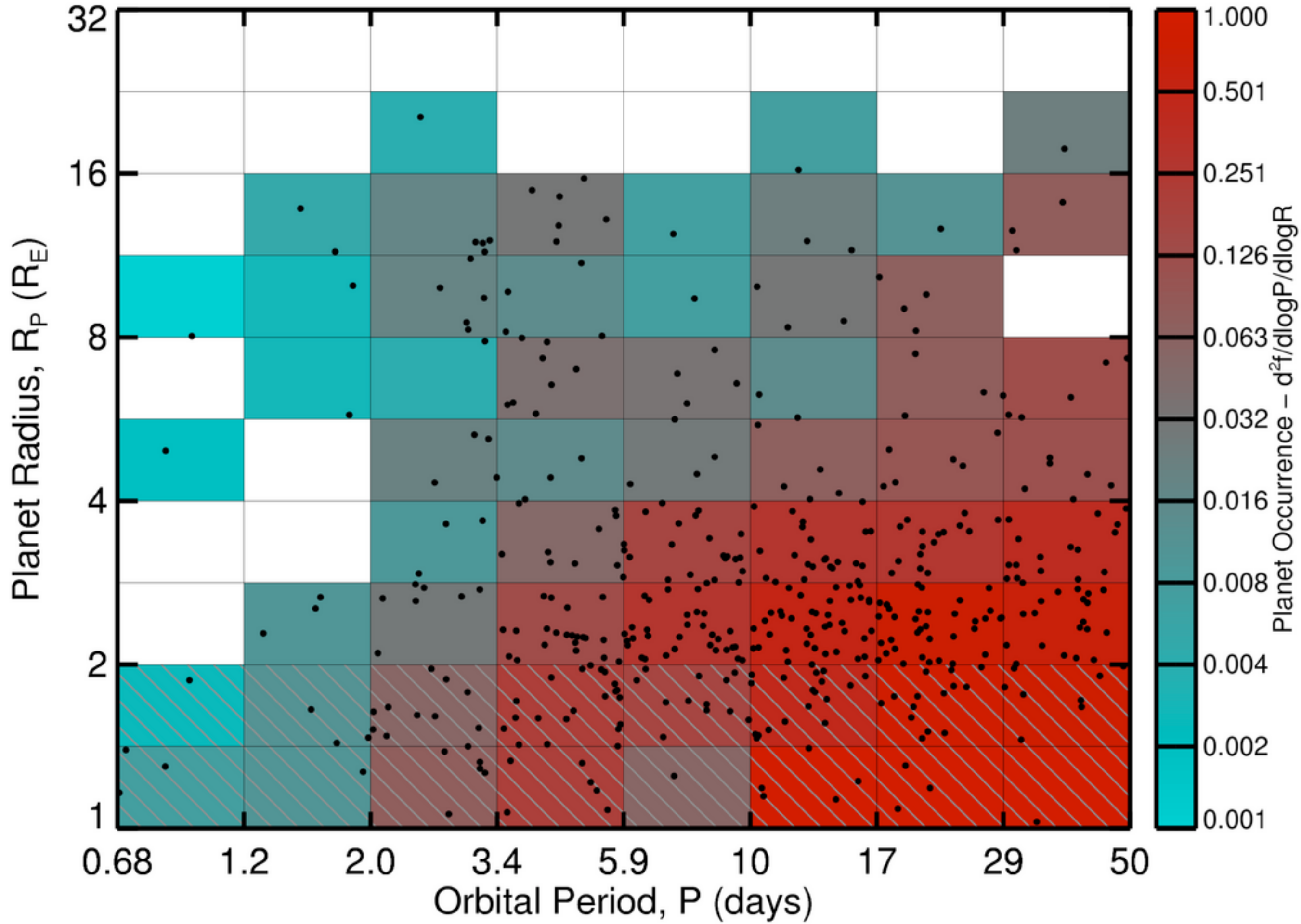


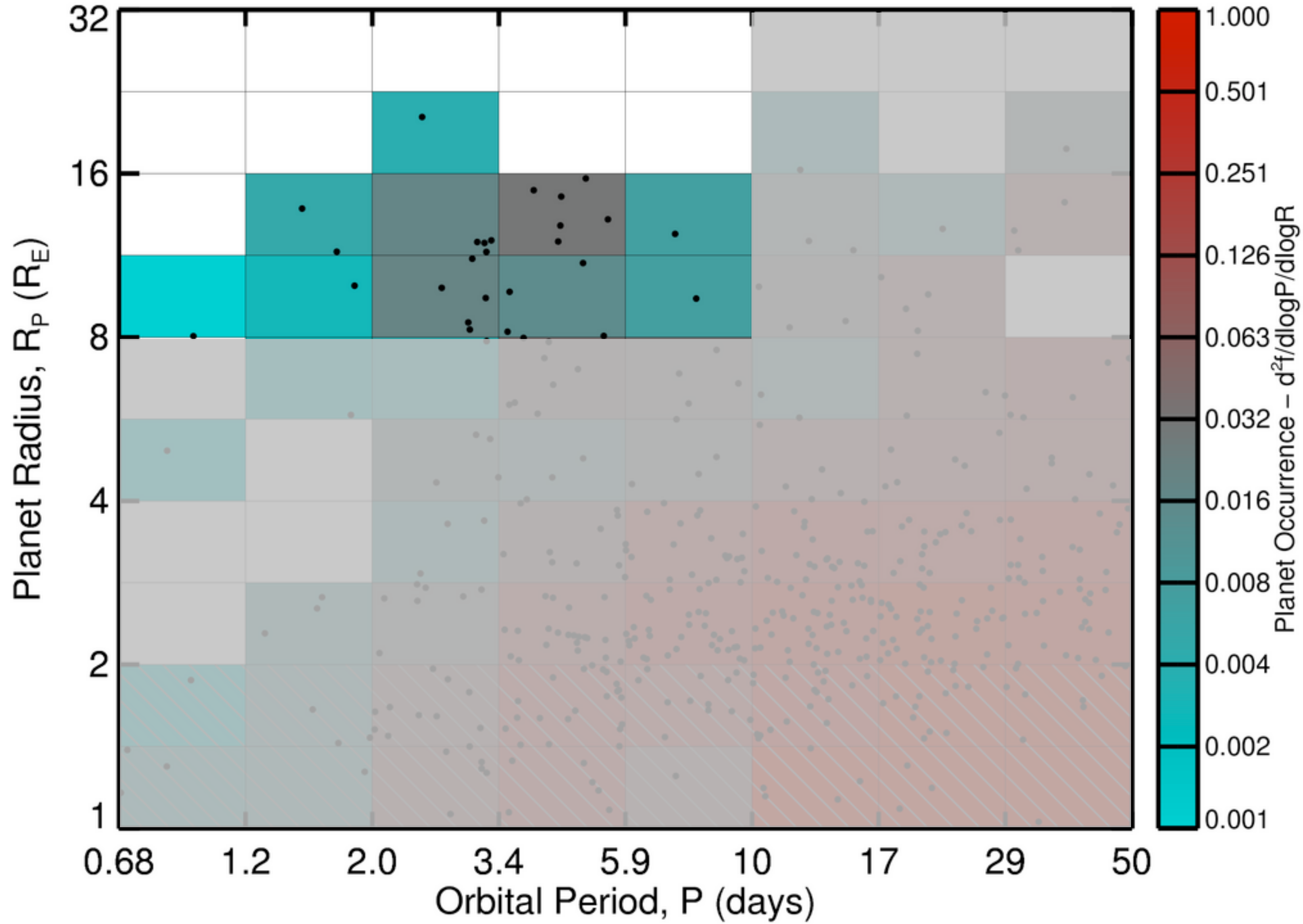
Power law with Exponential Cutoff:  

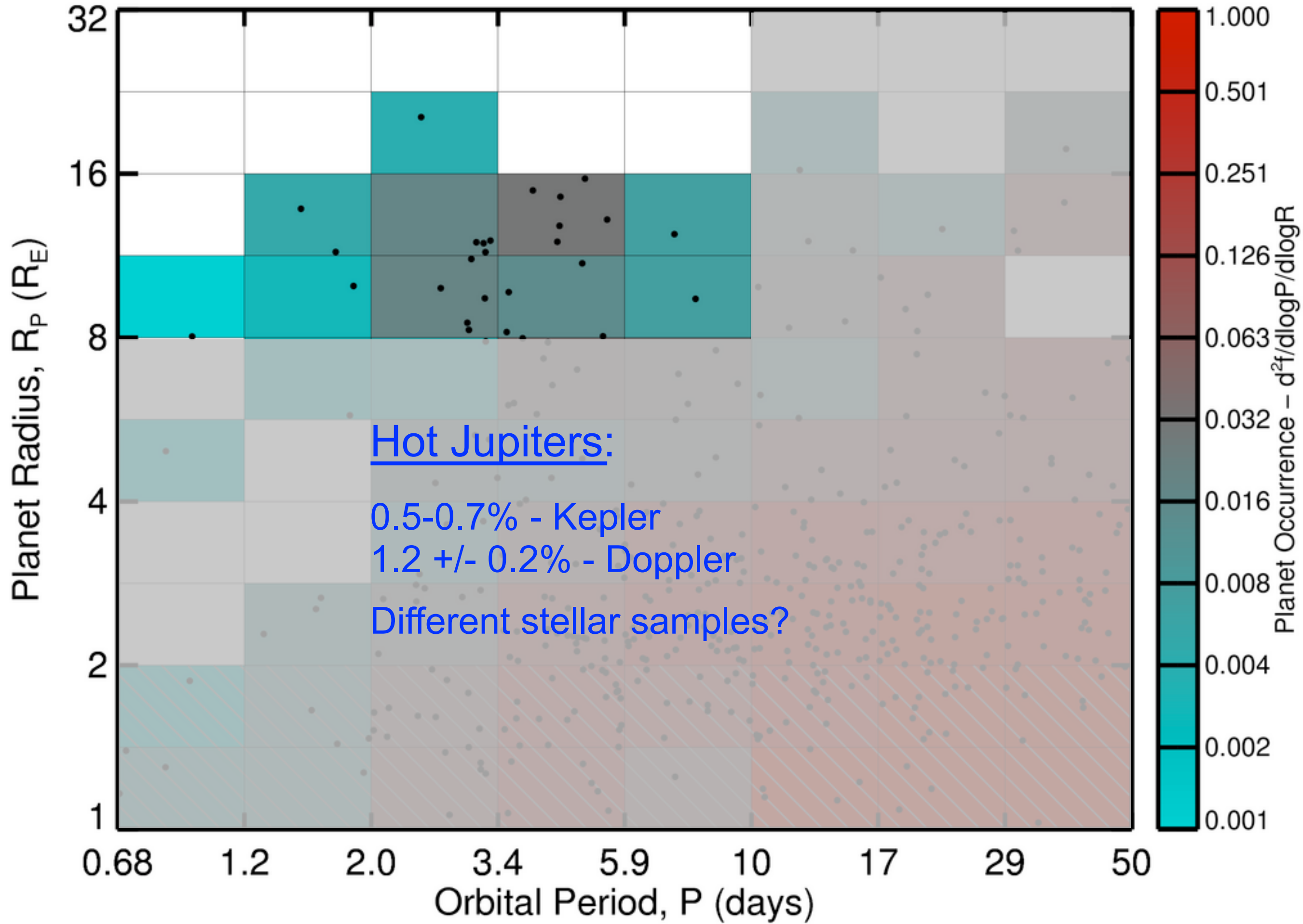
$$dN/d\log P = kP^\beta (1 - \exp(-(P/P_0)^\gamma))$$



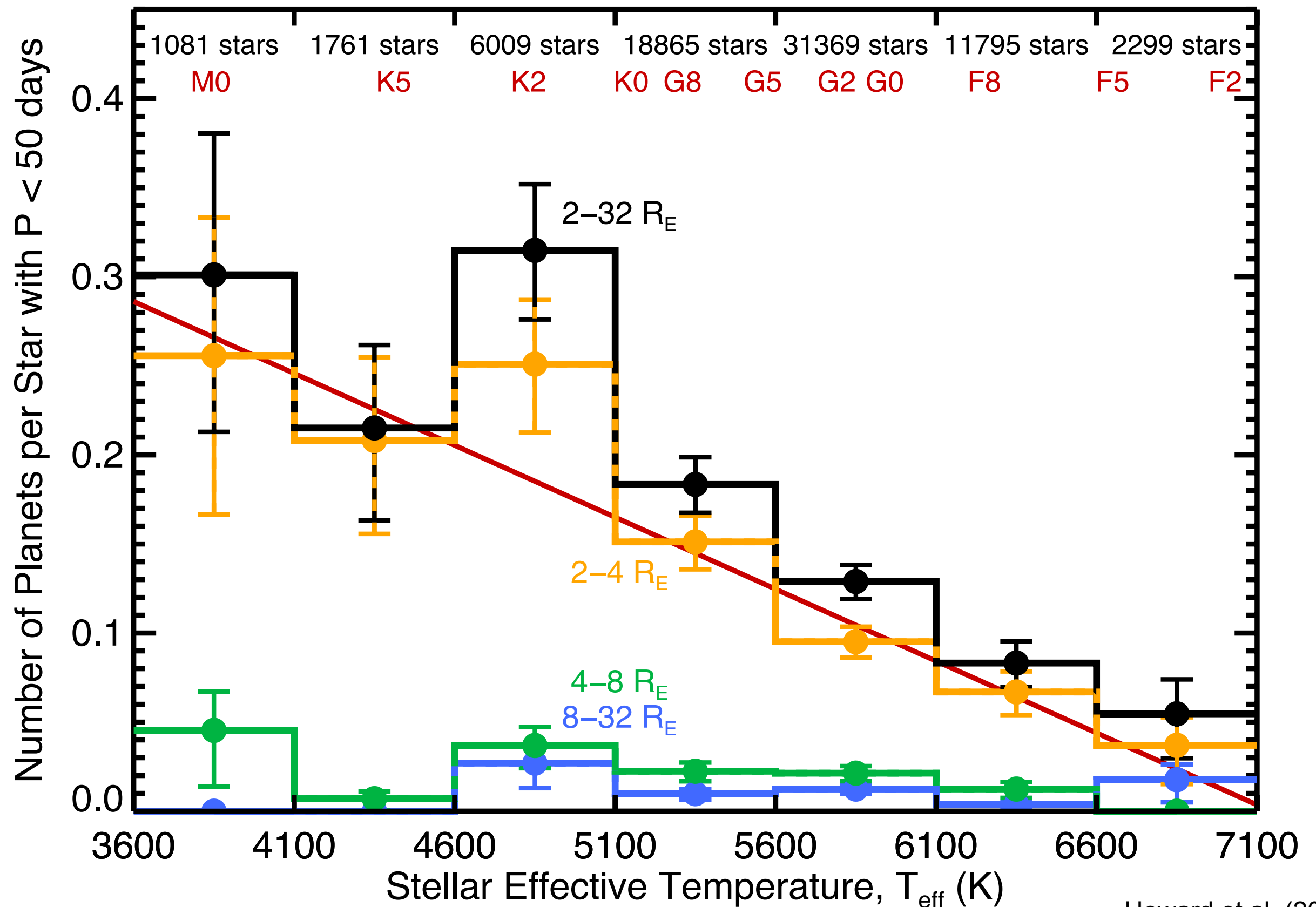
Power law with Exponential Cutoff:  
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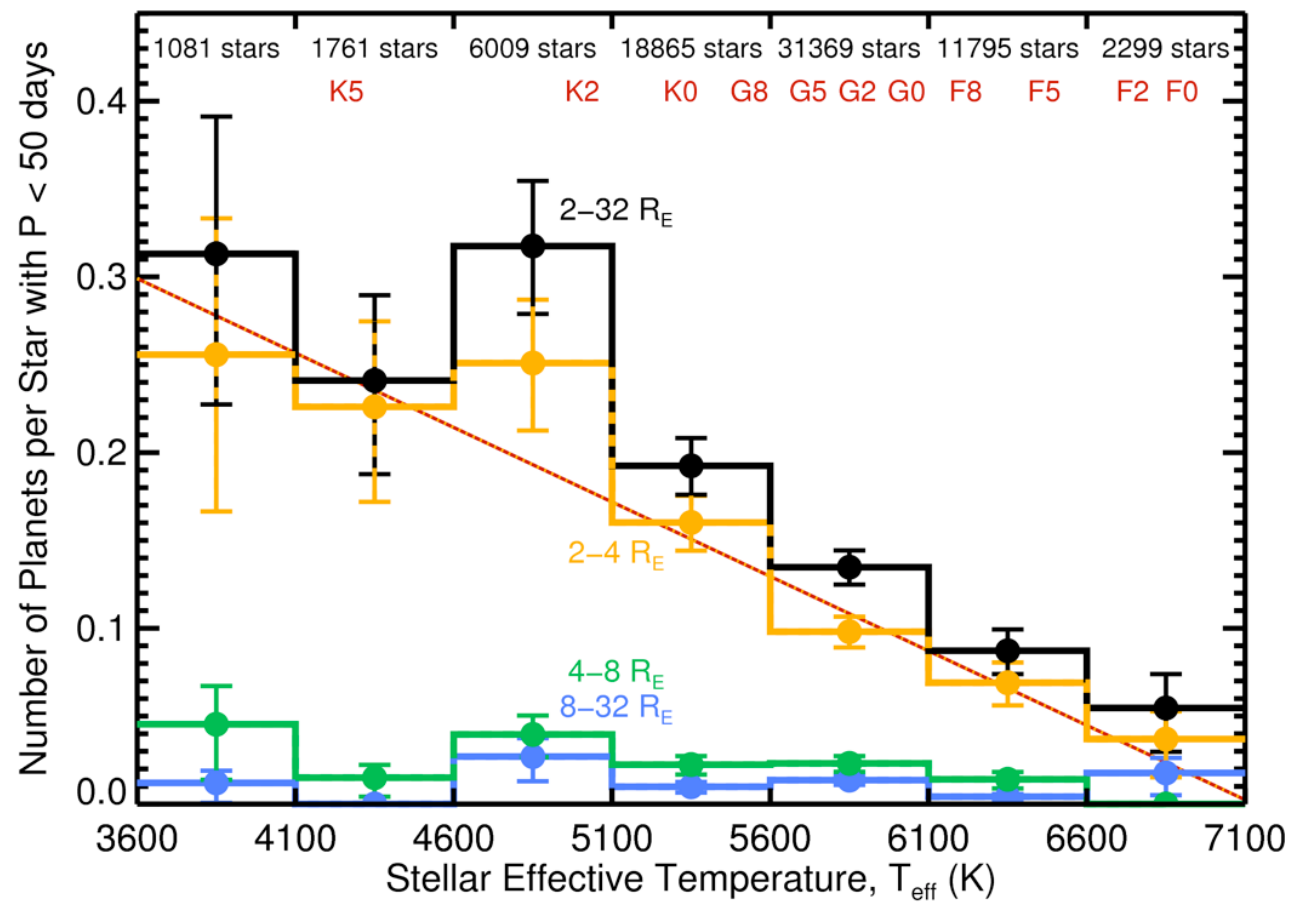




# Planet Occurrence vs. Stellar Temperature



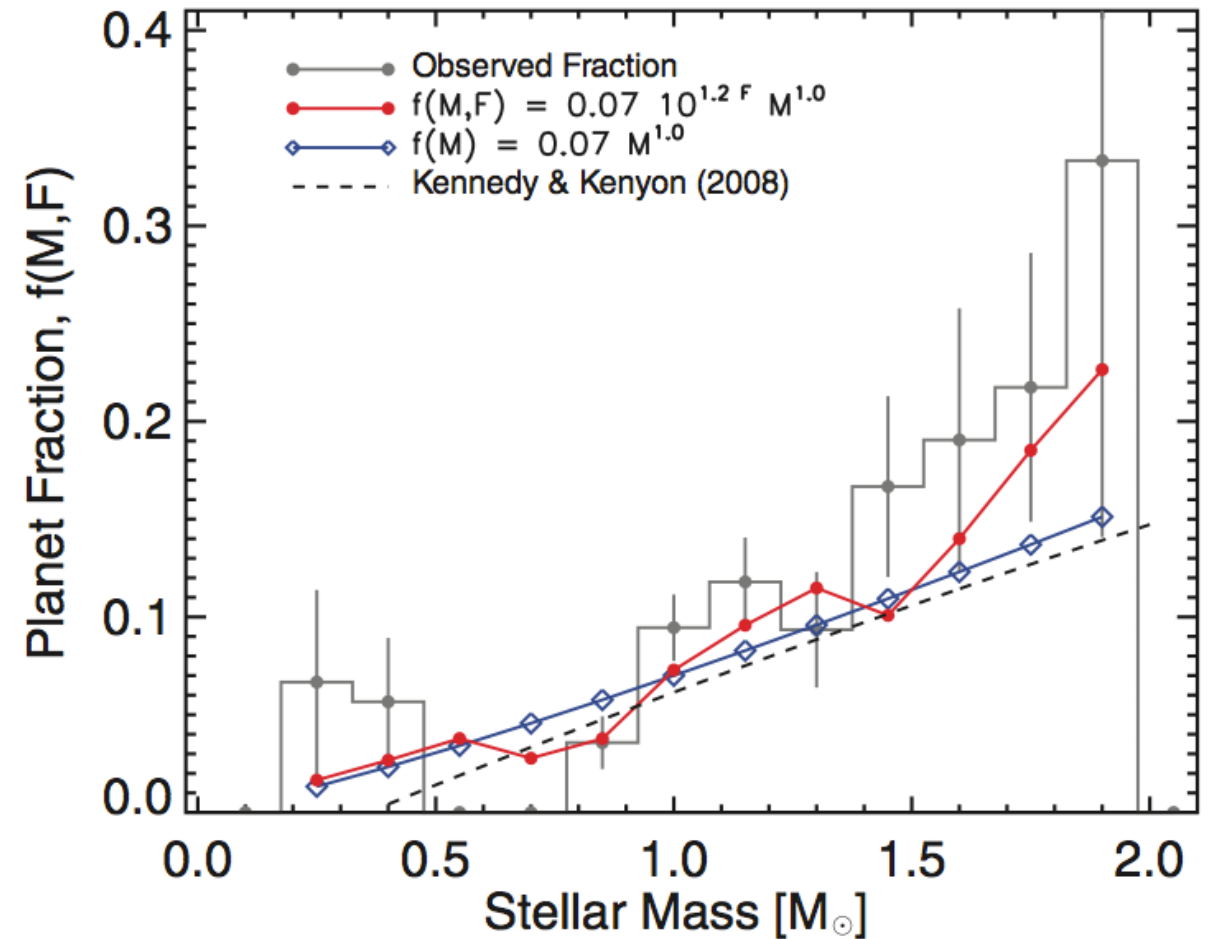
# Planet Occurrence vs. $M_\star$



Occurrence within 0.25 AU  
of small planets  
*decreases* with  $M_\star$



Howard et al. (2011c)

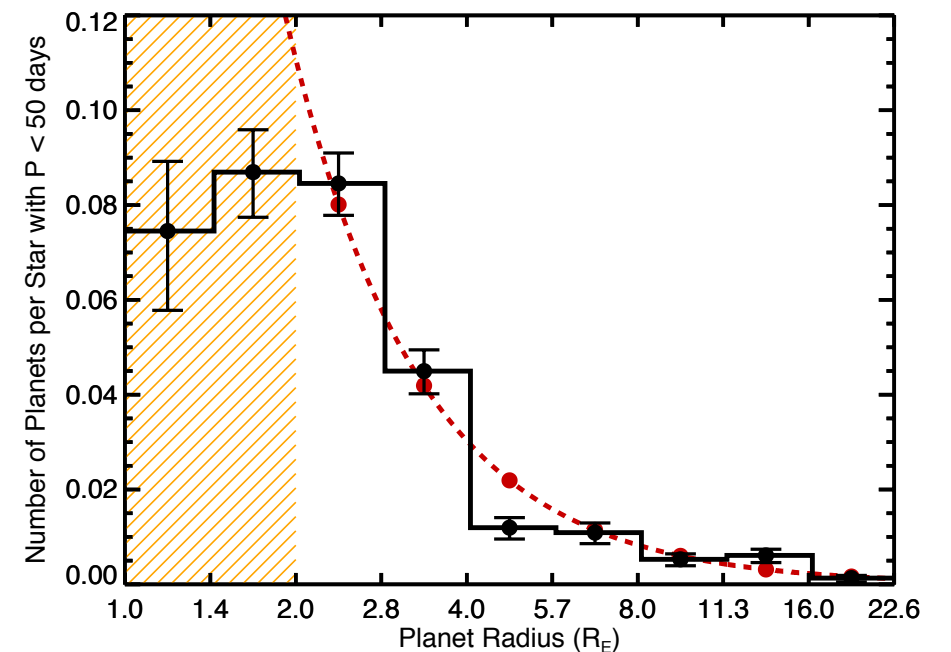
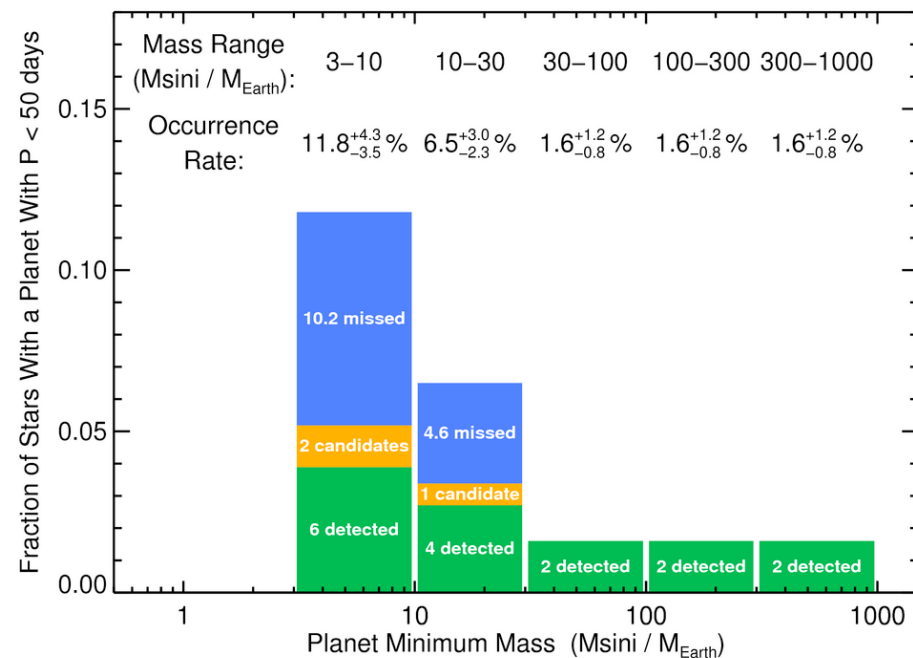
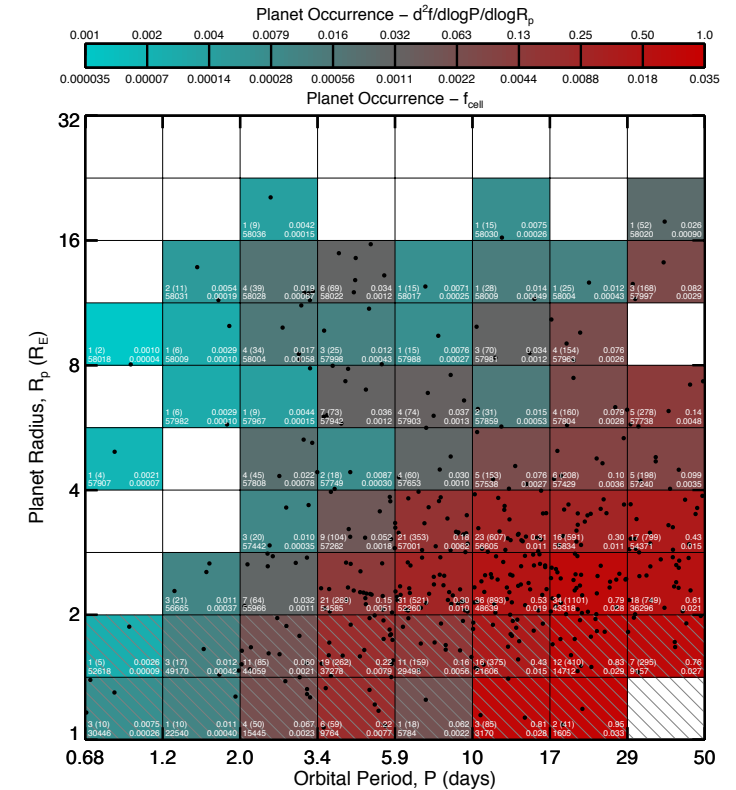
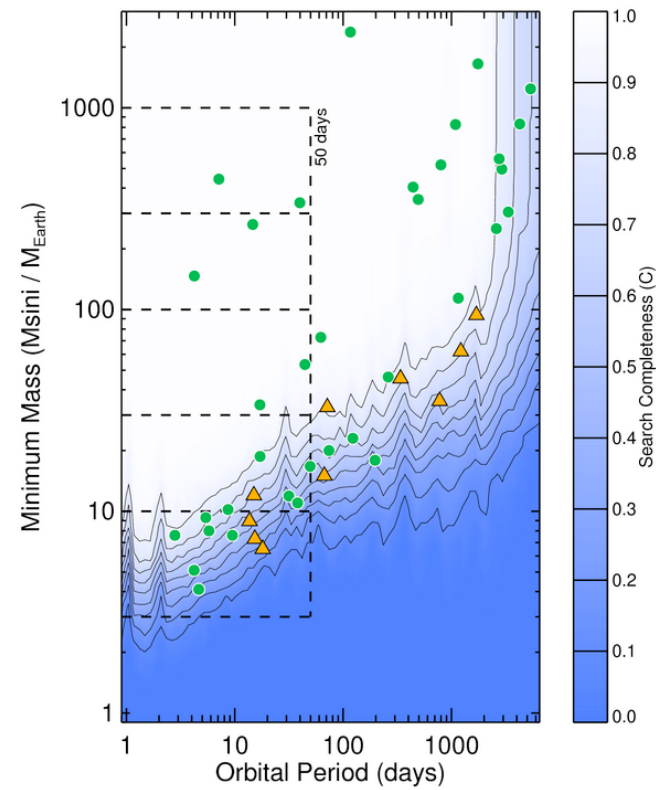


Occurrence within 2.5 AU  
of giant planets  
*increases* with  $M_\star$

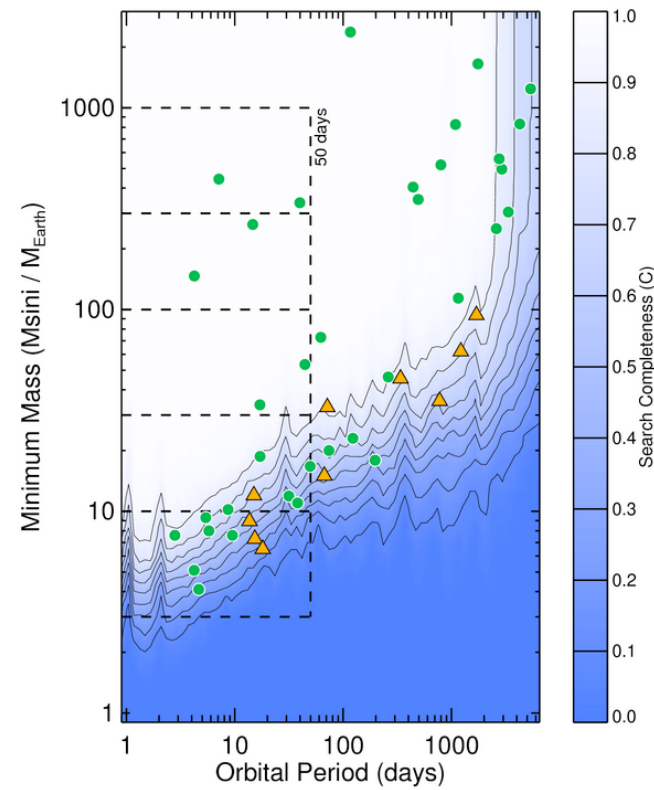


Johnson et al. (2010)

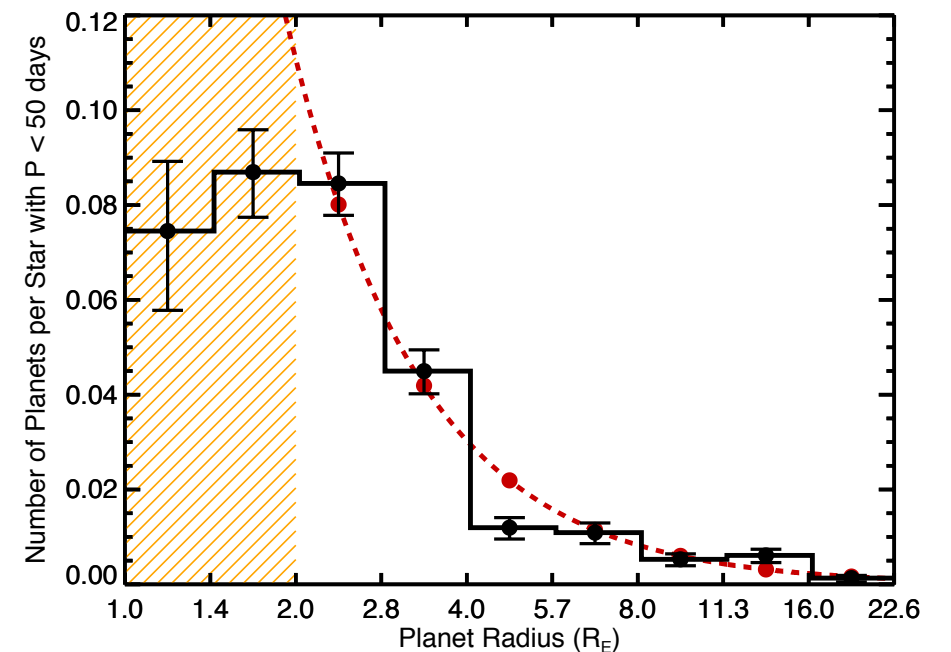
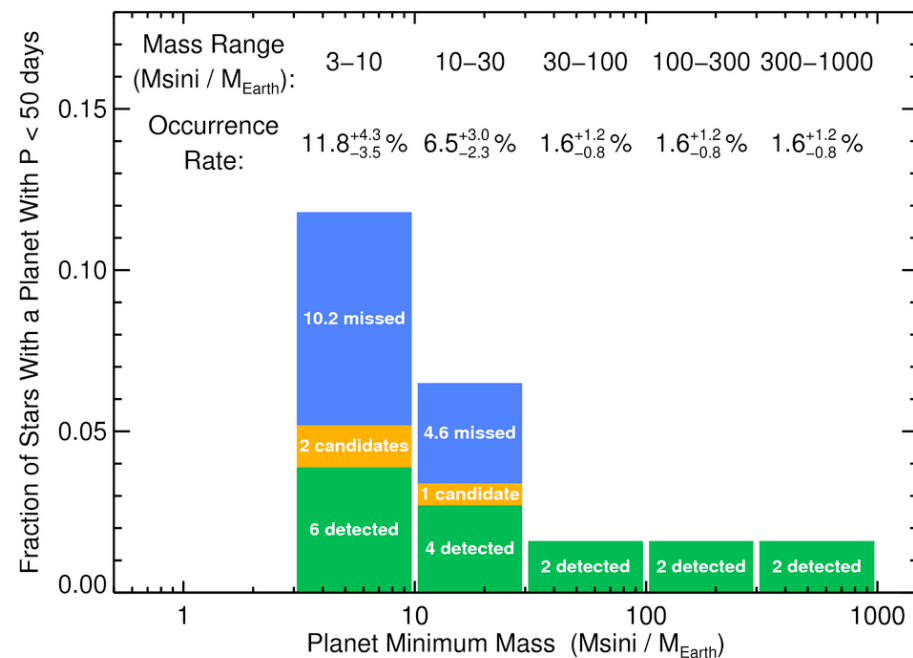
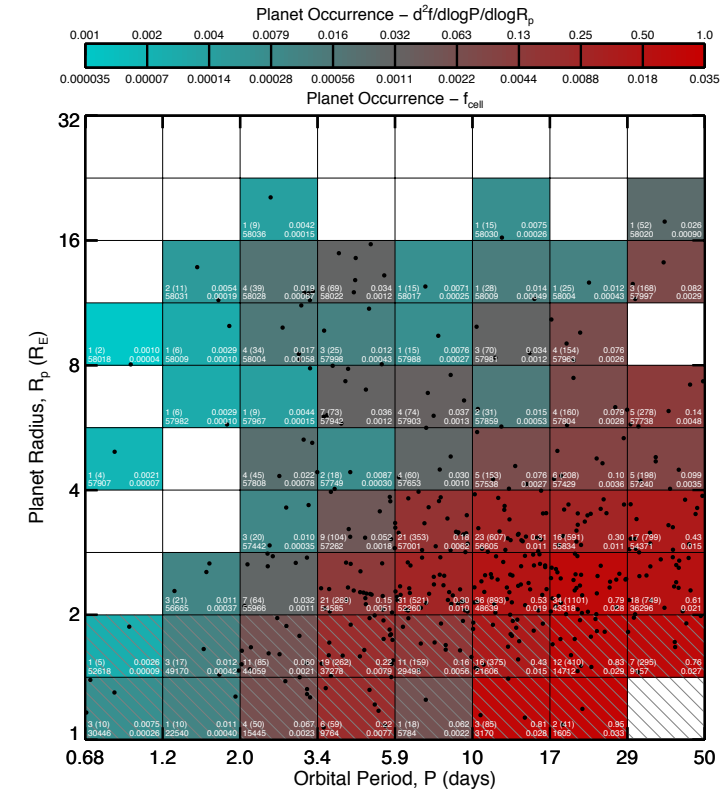
# Take Home Message:



# Take Home Message:



Small  
Planets  
are  
Common



Questions?

